

The information on halo properties contained in spectroscopic observations of late-type galaxies

Tariq Yasin,^{1*} Harry Desmond,² Julien Devriendt,¹ and Adrienne Slyz¹

¹*Astrophysics, University of Oxford, Denys Wilkinson Building, Keble Road, Oxford, OX1 3RH, UK*

²*Institute of Cosmology & Gravitation, University of Portsmouth, Dennis Sciama Building, Portsmouth, PO1 3FX, UK*

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ABSTRACT

Rotation curves are the key observational manifestation of the dark matter distribution around late-type galaxies. In a halo model context, the precision of constraints on halo parameters is a complex function of the properties of the measurements as well as properties of the galaxy itself. Forthcoming surveys will resolve rotation curves to varying degrees of precision, or measure their integrated effect in the H I linewidth. To ascertain the relative significance of the relevant quantities for constraining halo properties, we study the information on halo mass and concentration as quantified by the Kullback-Leibler divergence of the kinematics-informed posterior from the uninformative prior. We calculate this divergence as a function of the different types of spectroscopic observation, properties of the measurement, galaxy properties and auxiliary observational data on the baryonic components. Using the SPARC sample, we find that fits to the full rotation curve exhibit a large variation in information gain between galaxies, ranging from ~ 1 to ~ 11 bits. This range of information gain is equivalent to the difference between the prior shrinking to a flat posterior that is only a factor of 2 smaller compared to a factor of ~ 2000 smaller. It is predominantly caused by the vast differences in the number of data points and the size of velocity uncertainties between the SPARC galaxies. Out of the galaxy properties, only the degree of dark matter dominance is important. We also study the relative importance of the minimum H I surface density probed and the size of velocity uncertainties on the constraining power on the inner halo slope, finding the latter to be significantly more important. We spell out the implications of these results for optimising the strategy of galaxy surveys aiming to constrain galaxies’ dark matter distributions, highlighting spectroscopic precision as the most important factor.

Key words: dark matter – galaxies: kinematics and dynamics – galaxies: statistics

1 INTRODUCTION

In the standard model of cosmology, dark matter clusters under the action of gravity to form virialised, approximately spherical structures known as haloes. Galaxy formation occurs when baryons fall into their deep potential wells, cooling and condensing to form stars. Constraining the relationship between galaxies and their dark matter haloes, the “galaxy–halo connection”, is therefore a necessary step to constructing a complete theory of galaxy formation.

The most direct method to measure dark matter halo density profiles is through spectroscopic observations of galaxies’ kinematics. The mass distribution is then inferred by comparing the observed motions of the gas and stars to the motions expected from the observed luminous matter through Newtonian gravity. Late-type galaxies are ideal systems to study, as their stars and gas follow approximately circular orbits that directly trace the underlying potential. Resolved observations are able to measure rotation curves (RCs), the rotational velocity of the stars and gas as a function of radius (e.g. [Walter et al. 2008](#); [Ponomareva et al. 2016](#); [Lelli et al. 2016](#)).

It has been argued that observed correlations between the dark matter and baryons implied by the RCs (e.g. [McGaugh et al. 2019](#))

cannot be explained in the cold dark matter paradigm. This has led to an abundance of proposed extensions to the standard model, either in the form of new dark matter physics (e.g. [Khelashvili et al. 2022](#); [Adhikari et al. 2022](#)) or modifications to the theory of gravity (e.g. [Milgrom 1983](#); [Burrage et al. 2017](#); [Naik et al. 2019](#)). Improving the precision of constraints on halo properties is therefore also vital to assess the viability of dark matter models. One approach is to compare the halo properties inferred from kinematics to cosmological expectations (e.g. [Katz et al. 2017](#); [Li et al. 2020](#); [Posti & Fall 2021](#); [Mancera Piña et al. 2022](#); [Yasin et al. 2022](#)).

The density profile of dark matter is constrained by fitting parameterised functions to the measured RC. The parameters, which we refer to as the *properties* of the halo, are typically the virial mass, the concentration (a measure of the correlation of dark matter within the halo) and sometimes additional parameters describing the shape of the profile, such as the steepness of the inner slope. The tightness of the constraints on the free parameters is a complex function of the properties of the measurements (e.g. the number of measured RC data points, the uncertainties on the measured velocities, the maximum radius probed), properties of the galaxy (e.g. stellar surface density, gas fraction) and the precision of auxiliary data on the galaxy (inclination, distance, baryon content). The SPARC ([Lelli et al. 2016](#)) database, the largest of its type, contains H I rotation curves for 175

* tariq.yasin@physics.ox.ac.uk

late-type galaxies. Collated from archival observations, the measurement properties vary greatly between different RCs. For example, the best sampled RC has 120 data points, and the worst sampled has 5. In this paper we use SPARC to study the precision of constraints on halo parameters as a function of measurement and galaxy properties. This will reveal the ways in which future surveys ought to be designed to maximise their constraining power on the dark matter distributions around galaxies.

Resolving RCs in detail out to large radius is a luxury afforded only by studying the closest galaxies with targeted observations using HI interferometry. Future and current surveys are seeking to study larger, statistically representative samples of galaxies (e.g. [Maddox et al. 2021](#); [Oosterloo et al. 2010](#)) and galaxies at larger distances, including to high redshift (e.g. [Stott et al. 2016](#)). These surveys will measure rotation velocities with varying precision and will probe different parts of the RC. For example, the MIGHTEE-HI survey is a blind survey that will measure rotation curves out to $z = 0.5$, and has already done out to $z = 0.08$ ([Ponomareva et al. 2021](#)). Due to the further distance to the galaxies, the RCs are sampled less finely than nearby catalogues such as SPARC. IFU surveys such as MANGA ([Bundy et al. 2014](#)) and SAMI ([Croom et al. 2021](#)) use optical emission from ionised gas as a tracer, which extends only as far as the galaxy’s stellar component (see also the recently released PROBES rotation curve compendium, [Stone et al. 2022](#)). Current surveys such as ALFALFA ([Haynes et al. 2018](#)), and future surveys such as the SKA ([Yahya et al. 2015](#)) measure the width of the global HI 21-cm emission line of a galaxy, providing a cheap but less constraining measure of rotational velocity for orders of magnitude more galaxies.

It is currently not known what constraining power these different types of kinematic measurement bring to halo properties. To begin investigation of this question, we study here the information content contained in different parts of the RCs of the SPARC galaxies. This is achieved by constraining halo properties using the RC summary statistics recorded in SPARC, each of which characterise a particular part of the RC. These statistics are V_{flat} , the speed of the flat part of the rotation curve, V_{max} , the peak rotational speed and $V_{2.2}$, the circular velocity at 2.2 times the disc exponential scale length and W_{p20} the width of the spatially integrated HI line profile at 20% of the peak flux.

RC summary statistics have previously been utilised extensively in the study of the Tully–Fisher relation (TFR), the tight empirical relationship between the mass or luminosity of a galaxy and some measure of its rotational velocity. The TFR was first defined using HI linewidth ([Tully & Fisher 1977](#)), and later studies have used V_{max} ([Conselice et al. 2004](#)), $V_{2.2}$ ([Courteau 1997](#)) or V_{flat} ([McGaugh 2005](#); [Ponomareva et al. 2018](#)). Recently [Lelli et al. \(2019\)](#) (henceforth L19) compared the baryonic TFR produced by different summary statistics of the RC, as well as the HI line width, for the SPARC galaxies. The study of the dark matter halo constraints offered by these summary statistics is also interesting in this context.

We quantify the precision of constraints on halo properties using the Kullback–Leiber divergence ([Kullback & Leibler 1951](#)) of the posterior from the prior, a measure of information gain based on information theory. We study the information gain of the 2D total mass–concentration posterior (where total mass is equal to the halo mass plus the galaxy mass). We use total mass rather than halo mass, because the lower bound of the prior on total mass can be naturally set to the galaxy mass, whereas the halo mass, when sampled logarithmically (as is computationally necessary), has no natural lower bound. Our method could equally be applied to additional parameters describing shape. By quantifying the information gain on these halo

properties when using either the full RCs, HI line widths or summary statistics to constrain the kinematic model, we aim to answer the following questions:

- How does the information content depend on properties of the measurement such as velocity uncertainties, the minimum HI surface density probed and auxiliary data on galaxy parameters.
- How does the information content depend on galaxy properties?
- How much information is contained in different summary statistics compared to the full RC?
- How much information is there in cheap integrated HI measurements relative to expensive resolved measurements?

To answer the first two questions we will construct a predictive model for D_{KL} given the galaxy and measurement properties as input. The paper is structured as follows. Section 2 describes the the SPARC data. Section 3 describes the Bayesian models for inferring halo properties from the observations. In Section 4 we present the information content as a function of different types of observations, measurement properties and galaxy properties. We discuss the implications of our results in Section 5, and conclude in Section 6. We define the halo mass M_{halo} using the overdensity condition $\Delta_{\text{vir}} = 178$ of [Bryan & Norman \(1998\)](#). All logarithms are base-10 unless stated otherwise.

2 DATA

SPARC¹ ([Lelli et al. 2016](#)) is a database of rotation curves for 175 late-type galaxies, collated from archival resolved HI observations. In addition, 56 galaxies have hybrid rotation curves with high resolution $H\alpha$ data in the inner parts. Each galaxy has Spitzer photometry at $3.6\mu\text{m}$, a band in which the mass-to-light ratio is relatively constant ([Schombert et al. 2018](#)), which reduces the disk-halo degeneracy in kinematic analyses. The distributions of the stars and gas (*mass models*) are provided in the form of the contribution of each component to the circular velocity as a function of radius. We also utilise the HI surface density (Σ_{HI}) as a function of radius (F. Lelli, private communication) in our analysis. The SPARC galaxies span a wide range in luminosity (10^7 to $10^{12} L_{\odot}$), surface brightness (~ 5 to $\sim 5000 L_{\odot}\text{pc}^{-2}$), gas mass ($\sim 10^7$ to $10^{10.6} M_{\odot}$) and morphological type (S0 to Im/BCD).

L19 calculate the RC summary statistics V_{flat} , V_{max} , $V_{2.2}$ and their associated errors for a subset of SPARC sample with cuts on properties such as i and number of data points. We use their definitions to calculate V_{max} and $V_{2.2}$ for the whole sample, but we calculate V_{flat} using our own definition, which we describe in Section 3.2. The definitions of the velocity measurements are summarised in Table 1. L19 also compile HI linewidths from archival data for various different definitions. We choose to use W_{p20} , the width at 20% of the peak flux, as it is available for the most galaxies. Unlike L19, we do not include the contribution from inclination to the observational uncertainties on the summary statistics and W_{p20} , as inclination is a free parameter in our inference.

Table 1. The summary statistics of the full rotation curve and H I linewidth definition used in this paper.

Measurement	Interpretation	Sample size
Full RC	The full rotation curve.	175
V_{flat}	The velocity of the flat part of the rotation curve. The algorithm to identify the flat part of the rotation curve is listed in Section 3.2 and differs slightly from the definition used in the SPARC database.	123
V_{max}	The maximum velocity of the rotation curve. If the rotation curve is continuously rising in the observed range then this is the outermost measured point	175
$V_{2.2}$	The velocity at twice the exponential disk scale length. This is found by linearly interpolating between datapoints on either side of the required radius.	167
W_{p20}	The width of the global H I 21-cm emission line of a galaxy measured at 20% of the peak flux.	148

Table 2. The free parameters in our kinematic model, their physical definitions and their Bayesian priors. We sample all parameters in logarithmic space except inclination and distance.

Parameter	Units	Definition	Prior
M_{tot}	M_{\odot}	Total mass $M_{\text{tot}} = M_{\text{halo}} + M_{\text{bar}}$	Flat in range $\log(M_{\text{bar}}/M_{\odot}) < \log(M_{\text{tot}}/M_{\odot}) < 15.5$
$c_{0.1}$	-	Halo concentration, as defined in Eq. (9)	Flat in range $0.5 < \log c_{0.1} < 2$
Y_{disc}	M_{\odot}/L_{\odot}	Disc mass-to-light ratio	Lognormal ($\mu = \log(0.5), \sigma = 0.1$)
Y_{bulge}	M_{\odot}/L_{\odot}	Bulge mass-to-light ratio	Lognormal ($\mu = \log(0.7), \sigma = 0.1$)
D	Mpc	Physical distance to galaxy	Gaussian prior from SPARC value and its uncertainty
i	deg	Inclination (0° face on; 90° edge on)	Gaussian prior from SPARC value and its uncertainty

3 METHODS

3.1 Rotation curve model

Dark matter halo properties are inferred by fitting a parameterised halo profile to the observational data. Different halo profiles have been studied extensively in literature (Katz et al. 2017; Li et al. 2020), but a clear picture has yet to emerge of the relationship between the properties of a galaxy and the shape of its halo. The haloes in dark matter-only simulations were found to have a universal profile by Navarro-Frenk-White (NFW) (Navarro et al. 1997). However profiles motivated by hydrodynamical simulations that interpolate between a cusp and a core (Di Cintio et al. 2014; Read et al. 2016) have been found to fit the SPARC data better than the NFW profile. However cores have been observed in high mass systems (Collett et al. 2017), including for the SPARC galaxies (Li et al. 2019, 2022), which is against the expectations of the mass-dependant behaviour of the supernova-induced core flattening observed in some hydrodynamical simulations. On the other hand many studies have found NFW haloes to be good fits to clusters and weak lensing data for high mass galaxies (e.g. Mandelbaum et al. 2016), and some have argued that the inference of cores from 21-cm rotation curve observations may be due to systematics (Roper et al. 2022). In light of this uncertainty we choose to study the NFW and Burkert (Burkert 1995) profiles as representative examples of a cusped and cored profiles respectively. Due to the preference of most SPARC galaxies for cores, we present our results primarily for the Burkert profile. Our conclusions are unaffected by using an NFW profile instead.

The NFW density profile is

$$\rho_{\text{NFW}}(r) = \frac{\rho_s}{\left(\frac{r}{r_s}\right) \left[1 + \left(\frac{r}{r_s}\right)\right]^2}, \quad (1)$$

where r_s is a scale radius and ρ_s a characteristic density. The enclosed mass at radius r is

$$M_{\text{NFW}}(r) = 4\pi\rho_s r_s^3 \left[\ln(1+x) - \frac{x}{1+x} \right], \quad (2)$$

where $x \equiv r/r_s$. The Burkert density profile is

$$\rho_{\text{Burkert}}(r) = \frac{\rho_s}{\left(1 + \frac{r}{r_s}\right) \left[1 + \left(\frac{r}{r_s}\right)^2\right]}, \quad (3)$$

and the enclosed mass is given by

$$M_{\text{Burkert}}(r) = 2\pi\rho_s r_s^3 \left[\frac{1}{2} \ln(1+x^2) + \ln(1+x) - \arctan(x) \right]. \quad (4)$$

We also analyse how well observations can constrain a shape parameter for the inner halo by studying the generalized-NFW profile (gNFW), with density

$$\rho_{\text{gNFW}}(r) = \frac{\rho_s}{\left(\frac{r}{r_s}\right)^{\alpha} \left[1 + \left(\frac{r}{r_s}\right)\right]^{3-\alpha}}. \quad (5)$$

This reduces to NFW for $\alpha = 1$. The mass enclosed is given by

$$M_{\text{gNFW}}(r) = 4\pi\rho_s r_s^3 [B(x/(1+x), 3-\alpha, 0)], \quad (6)$$

where $B(z; a, b) \equiv \int_0^z u^{a-1} (1-u)^{b-1} du$.

Typically rotation curve measurements extend to only a small

¹ <http://astroweb.cwru.edu/SPARC/>

fraction of the virial radius. Therefore constraining halo mass requires a large extrapolation. Whilst gNFW, NFW and Burkert differ in shape towards the centre, at large radii they all decline as $\rho \propto 1/r^3$, so the comparison between the inferred masses is fair.

The circular speed due to the dark matter at radius r is $V_{\text{DM}} = \sqrt{GM_{\text{DM}}(r)/r}$. It is conveniently expressed in terms of M_{halo} , the virial velocity (V_{halo}) and the virial radius (R_{halo}),

$$\frac{V_{\text{DM}}(r)}{V_{\text{halo}}} = \sqrt{\frac{M_{\text{DM}}(r)}{M_{\text{halo}}} \frac{R_{\text{halo}}}{r}}, \quad (7)$$

where $V_{\text{halo}} \equiv \sqrt{GM_{\text{halo}}/R_{\text{halo}}}$. Concentration is commonly defined based on the radius at which the logarithmic slope of the density profile is -2 (r_{-2} , which for NFW is equal to r_s)

$$c = \frac{R_{\text{halo}}}{r_{-2}}. \quad (8)$$

Following [Yasin et al. \(2022\)](#) we instead use the definition

$$c_{0.1} = \frac{R_{\text{halo}}}{r_{0.1}}, \quad (9)$$

where $r_{0.1}$ is the radius enclosing 10% of M_{halo} . This definition has three advantages: 1) it can be calculated in simulations by simply counting dark matter particles, rather than fitting a profile; 2) $r_{0.1}$ is defined for all halo profiles, as it does not require a particular slope; 3) it is not based on the halo's shape, and hence is more intuitive.

As the late-type galaxies studied are rotationally dominated, we assume the rotational speed is equal to the circular speed. We try a non-fiducial model where we add 10km/s in quadrature to all velocities as a crude "asymmetric drift" correction, and find it does not affect our results. The total circular speed, $V_c(r)$ is equal to the sum in quadrature of the circular speed due to the dark matter and each baryonic component (dark matter, gas, stellar disc, stellar bulge)

$$V_c^2(r) = V_{\text{DM}}^2 + Y_{\text{bulge}} V_{\text{bulge}}^2 + Y_{\text{disc}} V_{\text{disc}}^2 + V_{\text{gas}}^2, \quad (10)$$

where each V is also a function of r . The latter are tabulated in the SPARC database for each galaxy. The baryonic mass models depend on the assumed distance as

$$V_{\text{disc, bulge, gas}}(r) \propto \sqrt{D}, \quad (11)$$

and the radius depends on the assumed distance as

$$r \propto D. \quad (12)$$

The predicted rotational speed is found by correcting for the inclination i of the galaxy ($i = 0^\circ$ face-on; $i = 90^\circ$ edge-on)

$$V_{\text{pred}}(r) = V_c(r)/\sin i. \quad (13)$$

D , i , Y_{disc} and Y_{bulge} are free parameters in the inference. When fitting to the full RC, $V_{\text{pred}}(r)$ can be compared directly to the observed RC. For the summary statistics, $V_{\text{pred}}(r)$ is evaluated at the same radii as the observed data points, and then the same algorithm that was used to calculate each summary statistic from the observed RC is applied.

3.2 Definition of V_{flat}

The algorithm to calculate V_{flat} in the SPARC database (see [Lelli et al. 2015](#)) starts by defining the outermost observed data point as being the flat part of the RC, and then adds additional points to it

iteratively. At each step the next innermost data point at radius r_{i-1} is added if its speed is within 5% of the mean of the data points already included:

$$\Delta = \frac{|\bar{V} - V_{i-1}|}{\bar{V}} < 0.05. \quad (14)$$

If the difference is greater than 5% the process terminates, and V_{flat} is the mean of the points already included. A galaxy is only considered to have a defined V_{flat} if the flat part constitutes at least three points when the algorithm terminates. The definition depends on the distance between the points, which means finely sampled RCs can still be considered flat even if they are much steeper than less finely sampled RCs. To lessen this bias we change the condition to a limit on Δ per stellar disk scale length (R_{disc}),

$$\frac{\Delta}{(R_i - R_{i-1})/R_{\text{disc}}} < 0.10. \quad (15)$$

The condition is set to 10% per disc scale length so V_{flat} is defined for a similar number of galaxies (123) as the original definition (133). Our results are not sensitive to the exact value. For galaxies which meet both the old and new V_{flat} definition, the difference between the two values is negligible.

3.3 HI linewidth model

The summary statistics V_{max} , $V_{2.2}$ and V_{flat} can be calculated from the RC alone. To calculate a model W_{p20} that can be compared to the observed value, we must calculate a model HI integrated spectrum from the RC and the HI surface density profile. Our method is described and validated in detail in [Yasin et al. \(2022\)](#), but we describe it here in brief. We use the method of [Obreschkow et al. \(2009\)](#), treating the HI disc as a series of concentric infinitely thin rings each with circular velocity given by $V_{\text{pred}}(r)$. We want to calculate the flux observed at the wavelength λ that corresponds to gas with a radial velocity V_λ relative to the galactic centre. The flux from a single infinitely thin ring of gas is given by

$$\tilde{\psi}(V_\lambda, V_{\text{pred}}) = \begin{cases} \frac{1}{\pi \sqrt{V_{\text{pred}}^2 - V_\lambda^2}} & \text{if } |V_\lambda| < V_{\text{pred}} \\ 0, & \text{otherwise.} \end{cases} \quad (16)$$

We assume the gas has a constant velocity dispersion of 10km/s (based on observations of local galaxies: [Leroy et al. 2008](#); [Mogotsi et al. 2016](#)), which broadens the flux distribution from each ring

$$\psi(V_\lambda, V_{\text{pred}}) = \frac{\sigma_{\text{HI}}^{-1}}{\sqrt{2\pi}} \int_{-\infty}^{+\infty} dV \exp\left[-\frac{(V_\lambda - V)^2}{2\sigma_{\text{HI}}^2}\right] \tilde{\psi}(V, V_{\text{pred}}). \quad (17)$$

The total, integrated flux profile of the galaxy is then obtained by integrating across the whole HI disc:

$$\Psi_{\text{HI}}(V_\lambda) = \frac{2\pi}{M_{\text{HI}}} \int_0^\infty dr r \Sigma_{\text{HI}}(r) \psi(V_\lambda, V_{\text{pred}}(r)). \quad (18)$$

The resulting flux profile is symmetric, so the model W_{p20} summary statistic can be trivially calculated by finding a peak of the distribution, and moving outwards to larger speed until the flux drops to 20% of the maximum. Although the observed profile may be asymmetric, the important quantity is the width, so this should not bias the results. This is verified for the SPARC sample in [Yasin et al. \(2022\)](#).

3.4 Inference

Bayes' theorem is used to calculate the probability of our parameters θ condition on our data D given the model \mathcal{M} ,

$$\mathcal{P}(\theta|D, \mathcal{M}) = \frac{\mathcal{P}(D|\theta, \mathcal{M})\pi(\theta|\mathcal{M})}{p(D|\mathcal{M})}, \quad (19)$$

where $\mathcal{P}(D|\theta, \mathcal{M})$ is the likelihood of the data, $\pi(\theta|\mathcal{M})$ is the prior probability density and $p(D|\mathcal{M})$ the marginalised likelihood. For $V_{\text{flat}}, V_{\text{max}}, V_{2.2}, W_{\text{p20}}$, which consist of a single observation, the likelihood of the data is

$$\mathcal{P}(W_{\text{obs}}|\theta, \mathcal{M}) = \frac{\exp\{-(W_{\text{obs}} - W_{\text{pred}})^2 / (2\delta W_{\text{obs}}^2)\}}{\sqrt{2\pi}\delta W_{\text{obs}}}, \quad (20)$$

where W_{obs} is the observed quantity and W_{pred} the model prediction. For the full rotation curve the likelihood is

$$\mathcal{P}(D|\theta, \mathcal{M}) = \prod_i \frac{\exp\{-(V_{i,\text{obs}} - V_{\text{pred}}(r_i))^2 / (2\delta V_{i,\text{obs}}^2)\}}{\sqrt{2\pi}\delta V_{i,\text{obs}}}. \quad (21)$$

When fitting with summary statistics we find some galaxies have non-zero posterior probability at $M_{\text{halo}} = 0$. Therefore in order to raise the lower limit of the posterior to a finite value we sample $\log M_{\text{tot}} = \log(M_{\text{halo}} + M_{\text{bar}})$ rather than M_{halo} itself, setting the lower bound on its flat prior to be $\log M_{\text{bar}}$. The total free parameters are $\{M_{\text{tot}}, c_{0.1}, Y_{\text{disc}}, Y_{\text{bulge}}, i, D\}$. The priors on Y_{disc} and Y_{bulge} are lognormal with means of $\log(0.5)$ and $\log(0.7)$ respectively, and 0.1 dex scatter. The priors on i and D are normal with mean given by the observed values and scatter by the observational uncertainties from SPARC. The posterior is sampled using the emcee ensemble sampler (Foreman-Mackey et al. 2013). We set the number of walkers to 200 and the stretch move to $a = 2$. To ensure the chain is converged we run the sampler until the chain is at least 50 times the autocorrelation length in all parameters, or a minimum of 10,000 steps to ensure the posterior is densely sampled to aid in the calculation of D_{KL} . The first 25 autocorrelation lengths are discarded as burn-in.

3.5 Goodness-of-fit

We wish to examine the dependence of constraining power on the type and precision of the measurements. A nuisance effect is that constraints can be very tight for models that are a poor fit to the data, as a small fraction of parameter space can still have high likelihood relative to the rest of it, even if the absolute value of the likelihood is low for that region. To exclude galaxies that are bad fits to a particular profile, we examine the distribution of normalised residuals, which should be drawn from a standardised normal distribution for a perfect fit model (Andrae et al. 2010; Zentner et al. 2022). We consider models a bad fit if the probability that the distribution of residuals is drawn from a standardised normal is $p_{\text{fit}} < 0.05$, as calculated by the Kolmogorov–Smirnov test. We define a galaxy as overfit if the standard deviation of their residuals is less than 1, or underfit if it is greater than 1. This procedure finds 14 galaxies are overfit for NFW and 11 underfit for the Burkert profile in the fiducial model.

3.6 Abundance matching

For comparison we also show the information gain on halo properties from abundance matching, an empirical model that matches the

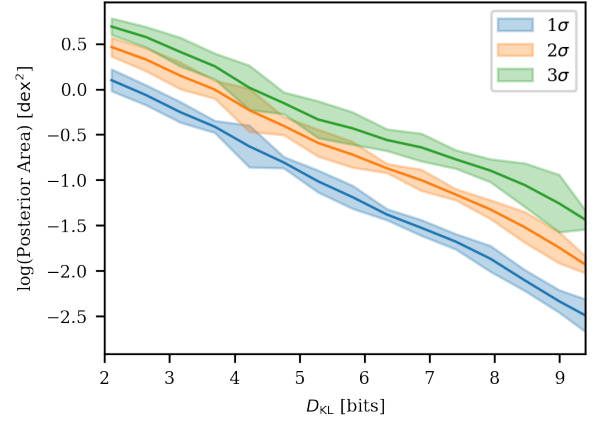


Figure 1. The correlation between the KL divergence (D_{KL}) of the posterior from the prior in $M_{\text{tot}} - c_{0.1}$ space and the area of the 2D posterior (for varying confidence intervals) from fitting Burkert haloes to the full RC for the whole sample of galaxies. The solid lines is the mean posterior area in bin of D_{KL} (15 bins total), with the band showing the 1σ spread in each bin. D_{KL} and posterior size are strongly correlated, with a small scatter due to changing shapes of the posteriors and the prior on M_{tot} that is a function of the galaxy's M_{bar} .

haloes in simulations to observed galaxies by positing an approximately monotonic relationship between a halo property (the *proxy*) and a galaxy observable (e.g. Conroy et al. 2006). In the simplest model the proxy is halo mass and the galaxy observable is stellar mass or luminosity. We use the proxy of Lehmann et al. (2016) which models assembly bias through the hybrid proxy $v_\beta := V_{\text{halo}} \left(\frac{v_{\text{max}}}{V_{\text{halo}}} \right)^\beta$, where v_{max} is the maximum circular velocity of the halo, and V_{halo} is the velocity at the virial radius. The free parameters in the model are the AM scatter σ_{AM} and β , which they constrain by clustering to $\beta = 0.57^{+0.20}_{-0.27}$ and $\sigma_{\text{AM}} = 0.17^{+0.03}_{-0.05}$. We calculate the posterior on mass-concentration using the Bayesian inverse subhalo abundance matching scheme of Yasin et al. (2022) (section 3.3), sampling over σ_{AM} and β .

3.7 The Kullback-Leibler divergence

The Kullback-Leibler divergence (Kullback & Leibler 1951) can be used to quantify the information gain in an experiment in going from the prior distribution to the posterior in units of bits

$$D_{\text{KL}}(P \parallel \pi) = \int_{\theta} \mathcal{P}(\theta) \log_2 \left(\frac{\mathcal{P}(\theta)}{\pi(\theta)} \right) d\theta. \quad (22)$$

It quantifies the similarity between $\mathcal{P}(\theta)$ and the reference distribution $\pi(\theta)$. In information theory terms it is the excess surprise when using $\mathcal{P}(\theta)$ compared to $\pi(\theta)$. It is the appropriate metric to use when comparing the improvement on precision in constraints between two experiments (Buchner 2022), as it takes into account the full probability distributions, as opposed to comparing a summary statistic such as the 2σ credible interval. However this comes at the expense of ease of interpretation. An intuitive example is an experiment with a flat prior that produces a flat posterior that is k times smaller. In this case $D_{\text{KL}} = \log_2(k)$.

To calculate D_{KL} , kernel density estimation (KDE) is used to estimate the posterior probability distribution. We use the fastKDE algorithm (O'Brien et al. 2014, 2016), which selects the kernel and

bandwidth based on the criteria of [Bernacchia & Pigolotti \(2011\)](#). We ensure that D_{KL} is converged with respect to the number of MCMC samples by checking that our results do not change when using a shorter or longer chain.

In Fig. 1 we plot D_{KL} against the 2D credible interval size for one of our runs, showing the strong correlation between them. Two factors cause a scatter between D_{KL} and the size of a contour. Firstly two posteriors with the same size 2σ contours will have a different D_{KL} if the rest of their contours are different. The prior on M_{tot} is also a function of the galaxy's M_{bar} . D_{KL} is dependent on the size of the prior. Our prior bounds are well motivated for mass, and the lower bound of concentration. However the upper bound of concentration is arbitrary. However we are interested in the relative D_{KL} between different measures, so our conclusions are not sensitive to the choice of prior.

3.8 The predictive model for D_{KL}

We aim to study the dependence of D_{KL} on the properties of the measurement and the properties of the galaxy. To do this we build a predictive model for D_{KL} using the ExtraTrees algorithm ([Pedregosa et al. 2011](#)). We optimise the hyperparameters through a grid search with 5-fold cross-validation. The features are the set of galaxy properties listed in the SPARC database. In addition, when fitting using the full RC, we add the following features that describe the details of the galaxy and RC: N (the number of RC data points); R_{out} (the radius of the outermost data point); $R_{\text{out}}/R_{\text{eff}}$ (the ratio of the radius of the outermost data point to the effective radius of the galaxy); $\frac{1}{N} \sum_i V_{\text{obs},i}/V_{\text{bar},i}$ (the mean ratio of the observed velocity to the baryonic circular velocity, which quantifies the mean dark matter dominance); $V_{\text{obs}}(r)/V_{\text{bar}}(r)$ at the radii R_{out} , R_{disk} and $R_{2.2}$ (this quantifies the dark matter dominance at different points in the galaxy); the summary statistics and their uncertainties; $\delta V_{\text{out}}/V_{\text{out}}$ (the uncertainty on the outermost data point); $\frac{1}{N} \sum_i \delta V_{\text{obs},i}/V_{\text{obs},i}$ (the mean velocity uncertainty). When fitting using individual summary statistics, as most of the above features are not relevant, we only add the summary statistic and its uncertainty, as well the ratio of the summary statistic to the baryonic circular velocity at the corresponding radius e.g. $V_{\text{max}}/V_{\text{bar}}(R_{\text{max}})$.

To find which features are most important in determining D_{KL} , we use the feature importance analysis method of [Stiskalek et al. \(2022, section 3.6\)](#). Features are added to the list of features used to train the ExtraTrees regressor one at a time, with the feature added at each increment the one that yields the greatest improvement in accuracy. This produces a list of features, ranked from most important (added first) to least important (added last), and the new accuracy after their inclusion. This method avoids the ambiguities associated with correlated features. Due to the small sample size, we divide the sample into 10 and calculate predictions for each subsample using a regressor trained on the rest of the samples. The accuracy of the model's predictions are assessed using the coefficient of determination

$$R^2 = 1 - \frac{\sum_i (y_{i,\text{true}} - y_{i,\text{pred}})^2}{\sum_i (y_{i,\text{true}} - \hat{y}_{\text{true}})^2}, \quad (23)$$

where $y_{i,\text{true}}$ is the test set value, $y_{i,\text{pred}}$ the corresponding prediction and \hat{y}_{true} the mean test set value. $R = 1$ corresponds to perfect accuracy, and $R = 0$ to a model that always predicts \hat{y}_{true} irrespective of the data.

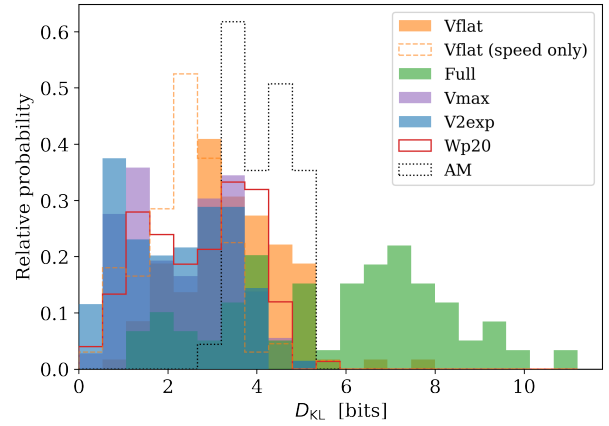


Figure 2. The distribution of KL divergence (D_{KL}) for each kinematic summary statistic, and for abundance matching (AM). The full rotation curve (Full) contains the most information. The single point summary statistics ($W_{\text{p20}}/V_{\text{max}}/V_{2.2}$) contain much less information and are similar to each other. V_{flat} has much less information than the full RC, albeit more than other summary statistics, showing the importance of the inner parts of the RC in constraining the shape of the halo and breaking the degeneracy between mass and concentration. Abundance matching has more information than any measure except the full RC.

4 RESULTS

4.1 Summary statistics

4.1.1 Overview

In our primary analysis we wish to study the dependence of the KL divergence (D_{KL}) on the type of measurement, properties of the measurement, and properties of the galaxy. In Fig. 2 we show the distribution of D_{KL} when fitting to the different types of measurement in the kinematic inference: the full RC, W_{p20} , or the summary statistics (see Table 1). The full rotation curve produces the tightest constraints, with a fairly flat distribution of D_{KL} between 4 and 10 bits (corresponding roughly to posteriors that are 16 and 1000 times smaller than the prior). The broad distribution of D_{KL} is due to the massive variation both in measurement properties and galaxy properties across the sample. The summary statistics and W_{p20} produce similar distributions in D_{KL} , with a minimum of 0 bits (meaning the posterior is similar to the prior) and a maximum of 6 bits (approximately half that of the full RC fits). V_{flat} has a higher mean D_{KL} than the other summary statistics. We also show the distribution of D_{KL} from abundance matching, which is narrow and has a mean 1-2 bits higher than the summary statistics.

We use the ExtraTrees algorithm to construct a predictive model for D_{KL} and carry out a feature importance analysis (as described in Section 3.8) and show the results in Fig. 3. The model is moderately predictive, with an accuracy of $R_2 = 0.77$. The most important features for predicting D_{KL} are, in descending order of importance: the number of data points N , the uncertainty on the outermost measured velocity $\delta V_{\text{out}}/V_{\text{out}}$, the fractional uncertainty on inclination $\delta i/i$, the mean ratio of the observed velocity to baryonic circular velocity $V_{\text{obs}}/V_{\text{bar}}$ and the ratio of the radius of the outermost data point to the effective radius $R_{\text{out}}/R_{\text{eff}}$. We reiterate that the uncertainties on V_{obs} do not include a contribution from inclination - which is treated separately.

As illustrative examples, in Fig. 4 we show the RCs and posteriors of F574-1, a low surface-brightness galaxy, and NGC4157, an intermediate mass spiral galaxy. F574-1 is an example of a galaxy that is

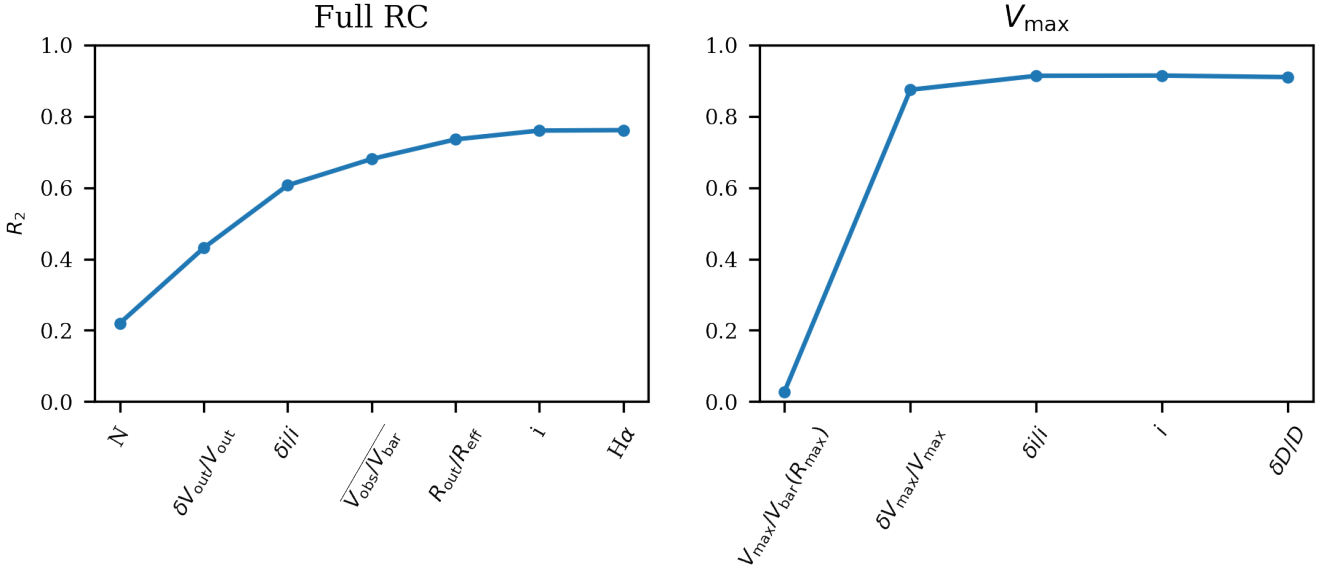


Figure 3. The important features for predicting the KL divergence (D_{KL}) from fitting to the full rotation curve (left panel) and fitting to V_{\max} (right panel) using an ExtraTrees regressor. Features are added left to right in the order which maximises accuracy at each increment, as described in Section 3.8. No features are predictive on their own, but for the full rotation curve the combination of number of data points N , the fractional uncertainty on the outermost measured velocity $\delta V_{\text{out}}/V_{\text{out}}$, the fractional uncertainty on inclination $\delta i/i$, and the mean ratio of observed rotational velocity to baryonic circular velocity $V_{\text{obs}}/V_{\text{bar}}$ (a measure of DM dominance) give reasonable accuracy, with $R^2 = 0.77$. For V_{\max} a combination of $V_{\max}/V_{\text{bar}}(R_{\max})$ and the fractional uncertainty $\delta V_{\max}/V_{\max}$ give good accuracy ($R^2 = 0.9$).

dark matter dominated. Its RC gradually rises and levels off to a flat part close to the last measured point. NGC4157 is baryon-dominated in its inner parts, with a RC that sharply rises to a maximum velocity that corresponds to the peak in the stellar velocity, before declining slightly to the flat part. Both galaxies have a similar number of data points, and uncertainties on observed distance and inclination. In general dark matter dominated galaxies have tighter constraints, as when V_{bar} is low relative to V_{obs} , the uncertainties on V_{bar} (which are set by the uncertainties on the mass-to-light ratios) are less important. The RC of F574-1 is also sampled further out into the halo (relative to the virial radius) than NGC4157, contributing to its tighter constraints.

4.1.2 W_{p20} , $V_{2.2}$, V_{\max}

Using either W_{p20} , $V_{2.2}$ and V_{\max} in the inference produces posteriors that are very similar in shape for most galaxies. We show the posteriors for W_{p20} in Fig. 4. As we have assumed the halo is spherically symmetric, the circular velocity due to the halo depends solely on its enclosed mass. This results in a perfect degeneracy between the halo mass and concentration for $W_{\text{p20}}/V_{2.2}/V_{\max}$, which do not constrain the shape of the RC. For the dark matter dominated F574-1, the posteriors are simply a band corresponding to the additional circular velocity required from the dark matter to generated the observed $W_{\text{p20}}/V_{2.2}/V_{\max}$, thickened by the its observational uncertainty and the uncertainties on i , D and $Y_{\text{disk/bulge}}$.

For NGC4157 the constraints on halo properties from $W_{\text{p20}}/V_{2.2}/V_{\max}$ are extremely weak, as the baryons alone can generate the observed values of these summary statistics. Therefore a large range of haloes are compatible with observation, as long as they do not significantly change: the mass enclosed within the HI disk for W_{p20} ; the maximum observed velocity for V_{\max} ; the velocity at $R_{2.2}$ for $V_{2.2}$. These three criteria result in similar constraints: a halo must

have mass or concentration low enough such that there is no significant halo mass at lower radii. In the case of W_{p20} , the degeneracy between mass and concentration can be broken by fitting the full HI flux profile rather than just the linewidth, as a long flat RC produces a very different HI profile to a RC that peaks and then declines (as occurs with very low mass / concentration haloes). We leave this to future work.

The mean D_{KL} for $W_{\text{p20}}/V_{\max}/V_{2.2}$ is 2.76/2.41/2.27. In Fig. 3 we present the D_{KL} feature importance analysis for V_{\max} only. V_{\max} is chosen because it is available for more galaxies than W_{p20} and $V_{2.2}$, but all three give similar results. The important features are, in descending order: $V_{\max}/V_{\text{bar}}(R_{\max})$ (which measures the dark matter-dominance at R_{\max}), $\delta V_{\max}/V_{\max}$, its fractional uncertainty, and $\delta i/i$. We interpret the ordering of the mean D_{KL} for the three measurements as being due to the dark matter dominance of the region probed by each quantity. The HI disc extends beyond $R_{2.2}$, and so probes the RC in the more dark matter-dominated outer regions. For most galaxies in our sample, V_{\max} coincides with the outer point of the RC (as in F574-1), which is typically beyond $R_{2.2}$. However for the galaxies with baryon-dominated inner regions such as NGC4157 V_{\max} roughly coincides with $R_{2.2}$. Hence the mean D_{KL} for V_{\max} is between W_{p20} and $V_{2.2}$.

4.1.3 V_{flat}

V_{flat} has higher mean D_{KL} than W_{p20} , V_{\max} and $V_{2.2}$. Its posteriors (see Fig. 4) are either a band similar to W_{p20} (NGC4157) or a truncated band (F574-1). There are two distinct contributors to the constraining power on halo properties from the V_{flat} statistic. The first is that M_{halo} and $c_{0.1}$ must generate a RC that meets the flatness criterion. If this is met, then the velocity of the flat part of the model RC must also be equal to the observed V_{flat} . The second criteria is

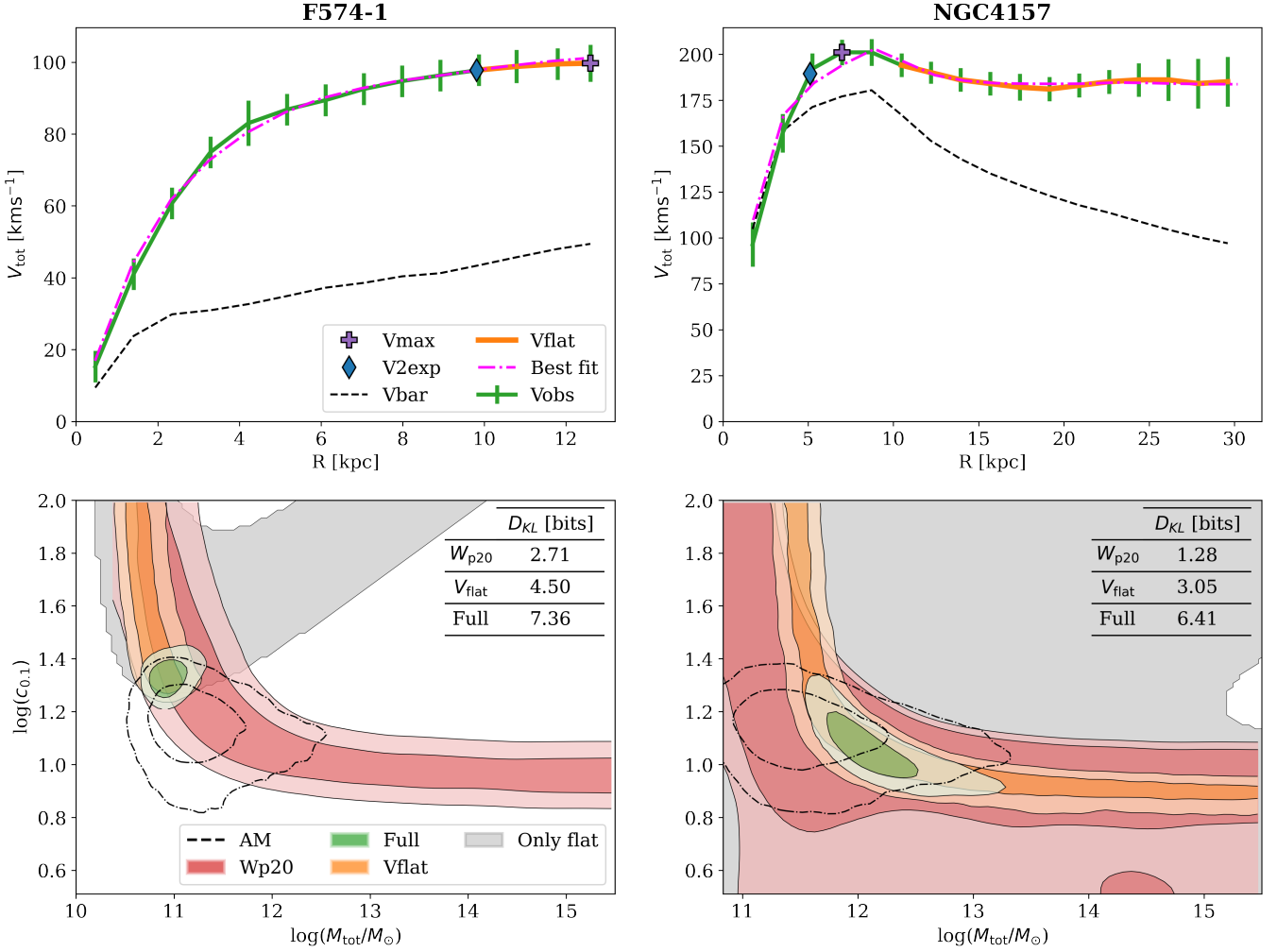


Figure 4. A graph showing the posteriors on mass-concentration for NGC4157 (left panel), an intermediate mass spiral galaxy, and F574-1 (right panel) a low-surface brightness galaxy, from fitting to the full rotation, fitting to the HI linewidth W_{p20} and fitting to the V_{flat} summary statistic (see Section 3.2). The grey shows the posterior from just requiring that the rotation curve flatness condition be met, without matching the actual value of V_{flat} . The D_{KL} of each posterior is shown on the right on the right. The posteriors for V_{max} and $V_{2.2}$ are not shown, but are similar to W_{p20} . The abundance matching posterior for each galaxy is also shown for comparison.

similar to the summary statistics $V_{2.2}$ and V_{max} , in that it simply requires one part of the RC to be a certain value, resulting in a degenerate band posterior. It is the flatness criterion that can truncate the band, as it does for F574-1.

To demonstrate the behaviour of the flatness criterion in isolation, in Fig. 4 we show the regions for which the RC is considered flat in grey. The shape of the RC depends on the parameters describing the baryons, so we consider a given $M_{\text{tot}}, c_{0.1}$ to meet the flatness criterion if the probability of the RC being flat is $> 34\%$ (i.e. 1σ) when marginalising over D and Y . A much smaller region of the $M_{\text{tot}}, c_{0.1}$ prior is considered flat for F574-1 than NGC4157. This is because for NGC4157 the RC with baryons alone is considered flat, as it is gently declining over many disc scale lengths. But generating a flat RC for F574-1 requires a dark matter halo that is both dominant over the baryonic component and has high enough concentration to have reached the gently declining “flat” part by the outer most RC point.

To separate the second criteria out from the flatness requirement, we calculate D_{KL} for a new summary statistic: V_{flat} (speed only). This

is a single data point that is the mean speed of the flat part of the RC, occurring at its mean radius, without any flatness requirement. We see in Fig. 2 that the D_{KL} for V_{flat} (speed only) are similar to $W_{\text{p20}}, V_{\text{max}}$ and $V_{2.2}$. This exercise demonstrates the extra constraining power that comes from observing a flat rotation curve over its length, compared to just measuring a single point from it.

We train an ExtraTrees regressor on D_{KL} for V_{flat} , but found it to be poorly predictive. This is due to difficulty in predicting the size of the region for which the flatness criterion is met, which depends on the detailed shape of the circular velocity due to the baryons.

4.2 D_{KL} as a function of measurement properties

In the feature importance analysis the uncertainties on inclination and velocity were found to be important predictors of D_{KL} . We

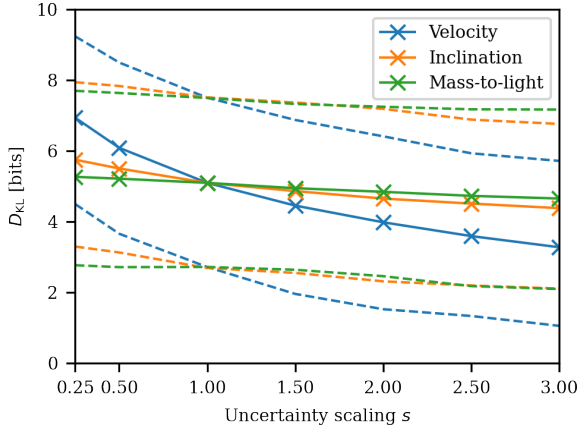


Figure 5. A plot showing the dependence of D_{KL} on the uncertainties on velocity, inclination and mass-to-light ratio. At each point, the uncertainties are scaled on either the velocities, the inclination, or the mass-to-light ratios on the bulge and disk by the factor s (such that $\delta' = s\delta$) and the D_{KL} from fitting to the full rotation curve is recalculated for each galaxy. Any galaxy that is considered underfit (see 3.5) for any value of s for any of velocity/inclination/mass-to-light is removed for this analysis. Therefore D_{KL} is calculated for the same sample of 98 galaxies for all points. The solid line shows the mean of D_{KL} at each value of s (marked by crosses), and the dashed lines show the 16th and 84th quantiles of the distribution. D_{KL} is most dependent on the velocity uncertainties.

now isolate their effect on D_{KL} for the full RC by scaling their uncertainties by a constant factor s , i.e.

$$\delta V_{\text{obs, scaled}} = s \delta V_{\text{obs}}, \quad (24)$$

and repeating the inference. We do this for both δV_{obs} and δi separately, for a range of values of s . We also apply the same procedure to the scatter on the prior of $Y_{\text{disc/bulge}}$ (changing it for both disc and bulge simultaneously).

We exclude bad fits using the residual analysis described in Section 3.5. To ensure we use the same sample for all three quantities, galaxies are only included if they are not bad fits for any value of s for all of V_{obs} , i and $Y_{\text{disc/bulge}}$. With a minimum value of $s = 0.25$ (the fits are worst for lower s), this leaves 98 galaxies in the sample. We show their D_{KL} as a function of s in Fig. 5. D_{KL} shows the greatest dependence on the velocity uncertainties, and is relatively flat for the rest. The rate of increase in D_{KL} steepens as s decreases for δV_{obs} and i .

R_{out} was identified as an important feature for predicting D_{KL} . It is set by the minimum HI surface density probed by the observation. We study the effect of varying the minimum HI density on D_{KL} , by repeating the inference with modified RCs that only include data points at radii where the HI surface density is above a chosen minimum value, which we vary. In this analysis we include only galaxies that each have a $\Sigma_{\text{HI}}(r)$ that fully spans the range 1 to 8 $\text{M}_{\odot}/\text{pc}^2$, and which do not have $H\alpha$ observations, leaving 45 galaxies. The D_{KL} for this sample is shown in Fig. 6. If one knew the uncertainty on Σ_{HI} , an alternative approach would be to vary the minimum signal-to-noise ratio rather than the surface density.

4.3 Halo profile comparison

We now study the constraining power of observations on whether a halo has a cusp or a core. In the above analysis δV_{obs} was found

to be the most important measurement uncertainty. Therefore we study the dependence of the likelihood ratio test between a NFW and Burkert profile on δV_{obs} , using the same uncertainty scaling procedure as above. We plot the resulting distribution of likelihood ratios in Fig. 7. We interpret a likelihood ratio greater than 100 as one halo profile being significantly favoured over the other. For $s = 1$ (no scaling) this occurs for around 30% of galaxies, with most favouring Burkert. For $s = 2$ this drops to 15%, and for $s = 0.5$ it rises to 70%.

Another way of looking at this is to examine how much constraining power an observation has on the shape parameter of 3 parameter halo profiles. The α parameter of the gNFW profile controls to the gradient of the inner slope. A cusped profile has $\alpha = 1$ and a cored profile $\alpha = 0$. We repeat the uncertainty scaling procedure above, but this time fitting a gNFW profile instead. The uncertainty on the marginalised α parameter (which we take to be its standard deviation, $\text{std}(\alpha)$) measures how well the inner slope is constrained. We study the distribution of $\text{std}(\alpha)$ for the sample as a function of s in Fig. 7. The mean scatter on α only drops below 0.2 for $s \sim 0.6$.

In the right column of Fig. 7, we repeat the same halo profile comparison analysis as above, but this time varying the minimum HI density instead s . We find the dependence is much weaker than for δV_{obs} .

5 DISCUSSION

5.1 Predicting information gain

Constraining halo mass and concentration from fits to rotation curve data is routine procedure in the study of late-type galaxies. However, to our knowledge, this study is the first to formally quantify the precision of the constraints and study their variation with galaxy and measurement properties. For the SPARC sample, we found massive variation in precision on M_{tot} and $c_{0.1}$ when fitting to the full RC, ranging from 1 to 11 bits of information gain. This range is equivalent to the difference between a flat prior shrinking to a flat posterior by only a factor of 2 (2^1) compared to ~ 2000 (2^{11}). We created a predictive model for D_{KL} using the ExtraTrees algorithm, and conducted a feature importance analysis to identify the galaxy and measurement properties that are the strongest predictors of information gain. The measurement properties are, in descending order of importance: the number of data points N , the fractional uncertainty on the outermost measured velocity $\delta V_{\text{out}}/V_{\text{out}}$, the fractional uncertainty on inclination $\delta i/i$, and the radius of the outermost measured velocity normalised by the effective radius $R_{\text{out}}/R_{\text{eff}}$. The only important galaxy property is $V_{\text{obs}}/V_{\text{bar}}$, a measure of the dark matter dominance. N ranks more highly than both the maximum radius of the RC and whether or not a galaxy has $H\alpha$ data. This shows the importance of sampling at many points across the RC to constrain the shape.

The moderate predictivity of our model ($R^2 = 0.77$) is part due to the small sample size, and part due to the input features not fully capturing the full details of the RC and the distribution of the baryons. For example N does not account for the autocorrelation of the RC, and $V_{\text{obs}}/V_{\text{bar}}$ is only an average. Our model is more predictive for the fits to the single data point summary statistics, W_{p20} , $V_{2.2}$ and V_{max} , with the latter giving $R^2 = 0.9$ and depending only on the uncertainty on V_{max} and $V_{\text{max}}/V_{\text{bar}}(R_{\text{max}})$, a measure of dark matter dominance.

Our feature importance analysis only ranks the measurement properties by importance. In order to quantify the size of their effect on D_{KL} , we varied the uncertainties on velocity and inclination whilst holding the rest of the inference constant. We also varied uncertain-

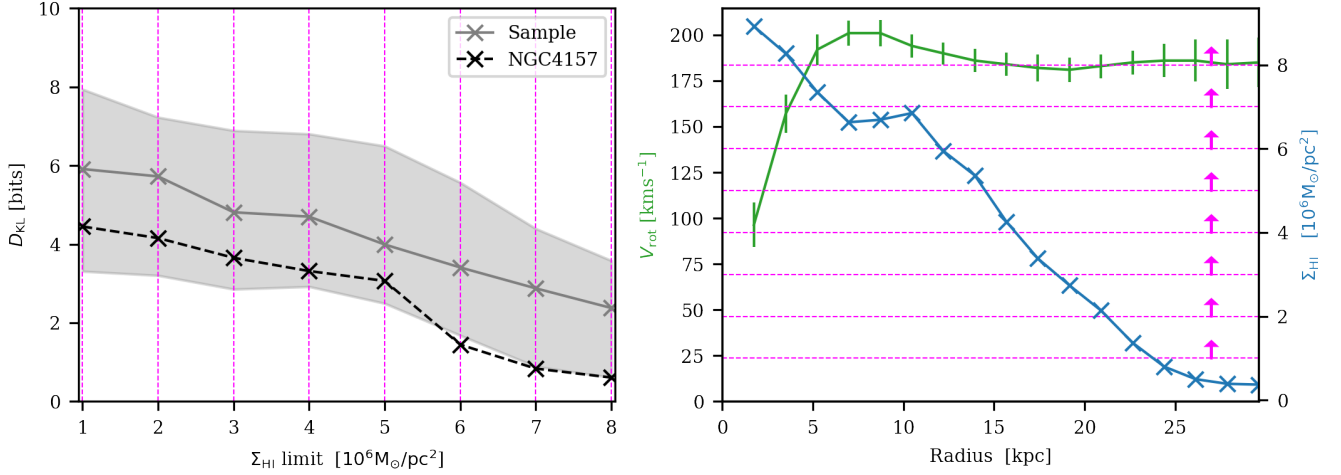


Figure 6. A plot showing the dependence of D_{KL} (grey line, sample mean; band, 16th and 84th quantiles) on the minimum HI surface density probed $\Sigma_{HI, \min}$ (left panel). The procedure for varying $\Sigma_{HI, \min}$ is illustrated for a single galaxy (NGC4157) in the right panel: for each value of $\Sigma_{HI, \min}$ (pink lines), only the data points of the rotation curve (blue) at radii where $\Sigma_{HI} > \Sigma_{HI, \min}$ (i.e. green > pink) are used in the inference. The D_{KL} for NGC4137 is shown as a dashed line in the left panel. This analysis is only applied to the subsample of 54 galaxies that each have an HI surface density profile that spans all the way from 1 to 8 $\Sigma_{HI}/M_{\odot}pc^{-2}$. For this sample, D_{KL} increases strongly (with an approximately linear relationship) as $\Sigma_{HI, \min}$ is reduced and the lower surface density regions towards the outskirts of the galaxy are added to the observation.

ties on the mass-to-light ratios Y_{disc} and Y_{bulge} , applied in our model through the priors. We find that the constraints are most dependent on the velocity uncertainties, with the inclination and mass-to-light showing only weaker dependence. In Fig. 6 we calculated the dependence of D_{KL} on the minimum HI density probed (which sets the maximum radius of the RC). Reducing the minimum density by $1 M_{\odot}/pc^2$ yields an additional 0.4 bits of information gain.

Our results can inform future survey design, by highlighting which features of the measurement should be prioritised for optimisation. The main way to improve constraints is of course to use longer integration times or higher instrument sensitivity, which would reduce the uncertainty on velocity and increase the number of data points and the maximum radius probed. However specific optimisations are possible. The beam size sets the maximum resolution. At fixed sensitivity/integration time, there is a trade off between minimum HI density that can be probed (i.e. the maximum radius) and resolution, although this can be altered with adaptive smoothing techniques in post-processing. The uncertainty on velocity is also dependent on the spectral resolution on the instrument. SPARC inclinations are produced as output of tilted-ring fits to the 2D velocity field, so the uncertainties would be reduced by improving the velocity map. Inclination can also be calculated using the ellipticity of the HI zeroth-moment map (Ponomareva et al. 2021), including forward modelling it to the datacube (Mancera Piña et al. 2021), or more imprecisely using optical data. Kourkchi et al. (2022) recently used a combination of machine learning and citizen science to improve inclinations from optical data. Schombert et al. (2022) use stellar population models to study the variation of the uncertainty on Y with the passband, and the available morphology and colour information.

5.2 Summary statistics

We compared the information gain when constraining halo parameters using the full RC, summary statistics or the HI linewidth W_{p20} . Our feature important analysis found the degree of dark matter dom-

inance at the radius of the velocity measurement was the most important factor in determining D_{KL} . The ranking of mean D_{KL} for the summary statistics from highest to lowest is $\{V_{flat}, W_{p20}, V_{max}, V_{2.2}\}$, which reflects how far out into the halo each measurement probes, and hence the degree of dark matter domination. L19 calculated the intrinsic scatter of the BTFRs constructed using the different summary statistics as the velocity measure. For the ones studied in this paper, they found $\{V_{flat}, W_{p20}, V_{max}, V_{2.2}\}$ had an intrinsic orthogonal scatter of $\{0.026, 0.035, 0.040, 0.070\}$ respectively, which is the inverse ordering of D_{KL} . Lelli et al. interpret the amount of scatter to be negatively correlated with the closeness of the summary statistic to the true flat value of the rotation curve (as opposed to the observed V_{flat} , which is limited by maximum radius probed in some galaxies). Hence it is not surprising that the mean information content on halo properties from each measurement is negatively correlated with the scatter of their respective BTFRs.

5.3 Constraining the inner halo shape

We studied the ability of observations to constrain the inner halo shape in two ways. Firstly we studied the impact of the δV_{obs} uncertainties on the ability to distinguish between the cored Burkert profile and cusped NFW profile using a likelihood ratio test. We found that with the unmodified velocity uncertainties, one profile was decisively favoured over another in 40% of cases, with this rising to 70% when the uncertainties are halved. On the other hand, when the uncertainties are doubled, one halo is only strongly favoured in only 20% of galaxies. We also studied the uncertainty of the inner slope parameter α when fitting a 3-parameter generalized NFW profile. With the normal SPARC uncertainties the mean $std(\alpha)$ is ~ 0.3 (the change in α is 1 between a cored gNFW profile and a normal NFW profile), but there are a significant number of galaxies with $std(\alpha) > 0.5$. This suggests a survey with velocity measurements more precise than SPARC is necessary to precisely constrain the inner halo shape for whole samples of galaxies. We repeated the analysis varying the

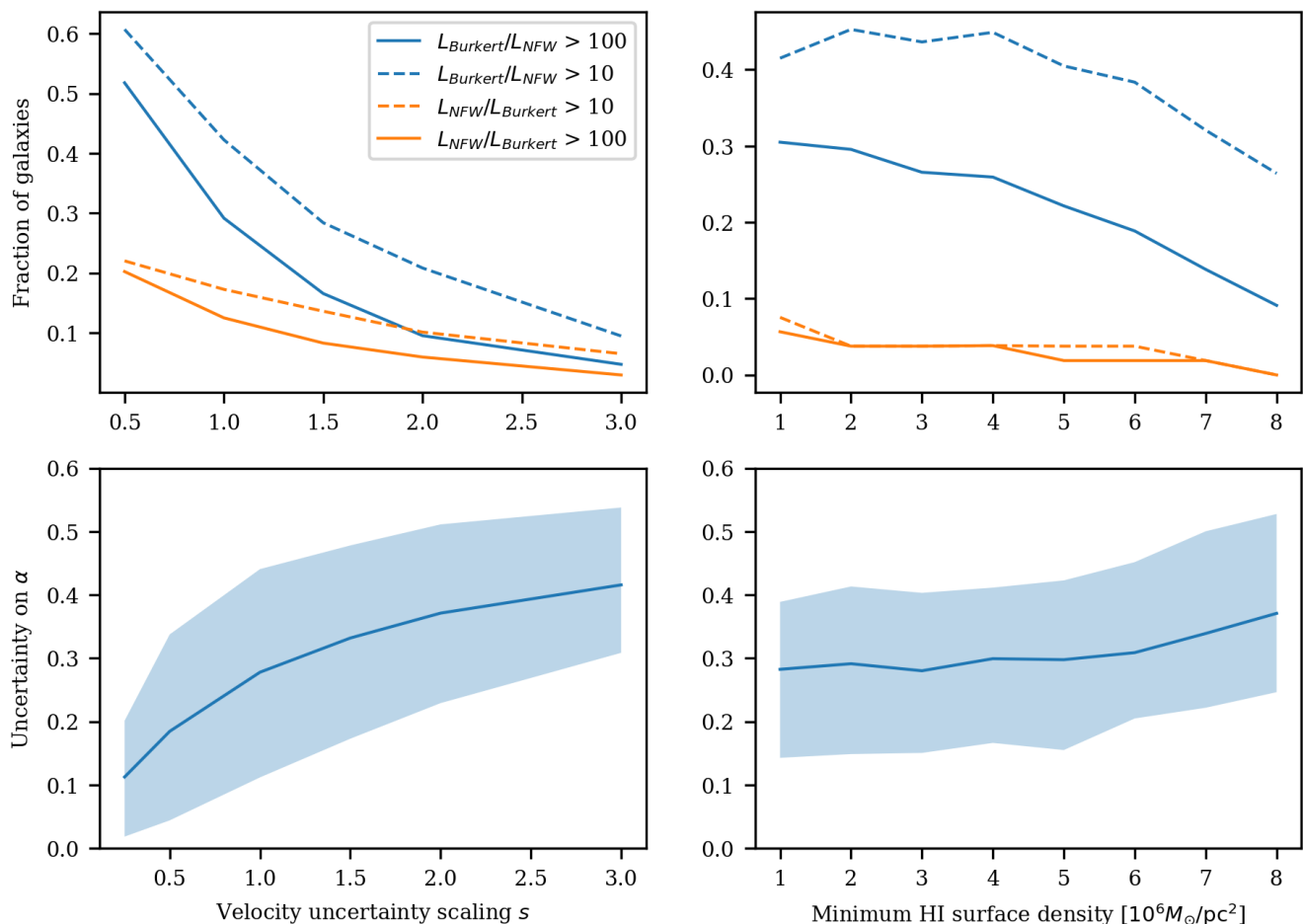


Figure 7. *Top panels:* The fraction of galaxies in the sample for which $L_{\text{Burkert}}/L_{\text{NFW}}$ (or the inverse) is above a certain value, as a function of a multiplicative scaling s of the velocity uncertainties (*left panels*, see Fig. 5) and the minimum HI surface density probed (*right panels*, see Fig. 6). We interpret $L_{\text{Burkert}}/L_{\text{NFW}} > 100(10)$ as strongly (moderately) favouring Burkert over NFW (and vice-versa). At $s = 1$ (no scaling) one halo is favoured over another in around ~40% of galaxies, with most favouring Burkert. If the velocity uncertainties are halved ($s = 0.5$), this rises to 70%.

Bottom panels The distribution of the uncertainties on the α shape parameter from fitting a gNFW halo to the full rotation curve (solid line, mean; band, 16th and 84th quantiles in bins of s). The uncertainty on α for each galaxy is calculated by taking the standard deviation of its marginalised distribution. A galaxy with a smaller uncertainty on α has a better known inner halo shape, with $\Delta\alpha = 1$ the difference between a cored and a cusped ($1/r$) inner profile. Both the uncertainties on α and the likelihood ratios show a strong dependence on the velocity uncertainties, but a weaker dependence on the HI surface density. Probing lower surface densities (below $\sim 4M_{\odot}/\text{pc}^2$) does not further constrain the inner halo shape but reducing the velocity uncertainties below the SPARC catalogue values does, showing that these should be targeted to improve our knowledge of galaxies' dark matter distributions.

minimum HI density probed, and found a much weaker dependence. Surveys examining the cusp-core problem should therefore target velocity precision over probing lower HI surface densities.

5.4 Comparison with the literature

We have highlighted the important measurement properties which should be targeted by future surveys. Identifying the observational parameters (such as integration time) required to achieve the desired measurement properties is beyond the scope of this work. Recent work has simulated spectroscopic HI observations of simulated late-type galaxies (Oman 2019), which in theory allows an end-to-end determination of the effect of observational properties such as integration time on the constraints on halo properties. However, our analysis of real observations is an important complementary approach, as simulations still struggle to produce realistic samples of

rotation curves (Roper et al. 2022). In light of increasingly expensive observations, but comparatively cheap computational resources, there are an increasing number of studies examining optimal observational strategies. For example, two recent studies have used the Fisher-matrix formalism to quantify the information content in stellar streams (Bonaca & Hogg 2018) and the cosmic web (Kostić et al. 2022) in order to identify the best observational strategy.

6 CONCLUSION

We have used the Kullback-Leibler divergence (D_{KL}) of the posterior on total mass-concentration (where total mass is equal to the halo mass plus the galaxy mass) from the prior to quantify the gain in information obtained from spectroscopic observations of the late-type galaxies of the SPARC database. We set the observable in the

kinematic inference to be either the full rotation curve, summary statistics of the rotation curve (V_{\max} , $V_{2.2}$, V_{flat}), or the linewidth of the integrated 21-cm spectrum, $W_{\text{p}20}$, in order to quantify the information contained in different parts of the rotation curve and different types of measurement. Further, to determine the properties of the measurements that are most important for the information gain, we study the variation on D_{KL} as we modify properties of the rotation curve observations such as the uncertainties on velocity or the minimum HI surface density probed. Our conclusions are as follows:

(i) The full RC fitting offers a wide range of information gain for the SPARC galaxies, ranging from ~ 1 bit to ~ 11 bits. This is predominantly due to the massive range in the number of data points each rotation curve has, and the large variation in velocity uncertainties.

(ii) Fits to the summary statistics of the RCs offer much smaller gains, ranging from ~ 0 to ~ 6 bits, as the posteriors are degenerate in mass–concentration and run up against the prior bounds. V_{flat} offers a modest increase due to the flatness constraint. For most SPARC galaxies $W_{\text{p}20}$, V_{flat} , $V_{2.2}$ and V_{\max} all probe regions of the rotation curve which are dark matter dominated, and hence contain similar information on the halo.

(iii) We measured D_{KL} as a function of the minimum HI surface density probed, and the uncertainties on velocity, inclination and mass-to-light ratios. Its dependence is strongest on the minimum surface density and the velocity uncertainties. These results can be used to weigh up the increase in precision on halo constraints afforded by improving each aspect of the measurement against the associated cost.

(iv) The tightness of the constraints on the inner halo shape are strongly dependent on the velocity uncertainties, but have a much weaker dependence on the minimum HI surface density. This suggests that whilst both sensitivity and spectroscopic precision are important for obtaining tight constraints on halo properties, surveys specifically targeting e.g. the cusp-core problem should prioritise spectroscopic precision.

Our study has identified the most important variables for improving the constraints on dark matter halo properties from spectroscopic observations of late-type galaxies. With forthcoming instruments set to greatly enhance our ability to probe the dark matter distribution around galaxies, in terms of number of galaxies, increasing redshift and measurement precision, these results should inform future survey design to maximise the return of knowledge about the galaxy–halo connection.

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DATA AVAILABILITY

The data underlying this article will be made available on reasonable request to the corresponding author.

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