

Bayesian galaxy shape measurement for weak lensing surveys – III. Application to the Canada–France–Hawaii Telescope Lensing Survey

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ABSTRACT

A likelihood-based method for measuring weak gravitational lensing shear in deep galaxy surveys is described and applied to the Canada–France–Hawaii Telescope (CFHT) Lensing Survey (CFHTLenS). CFHTLenS comprises 154 deg² of multi-colour optical data from the CFHT Legacy Survey, with lensing measurements being made in the i' band to a depth $i'_{\text{AB}} < 24.7$, for galaxies with signal-to-noise ratio $v_{\text{SN}} \gtrsim 10$. The method is based on the *lensfit* algorithm described in earlier papers, but here we describe a full analysis pipeline that takes into account the properties of real surveys. The method creates pixel-based models of the varying point spread function (PSF) in individual image exposures. It fits PSF-convolved two-component (disc plus bulge) models to measure the ellipticity of each galaxy, with Bayesian marginalization over model nuisance parameters of galaxy position, size, brightness and bulge fraction. The method allows optimal joint measurement of multiple, dithered image exposures, taking into account imaging distortion and the alignment of the multiple measurements. We discuss the effects of noise bias on the likelihood distribution of galaxy ellipticity. Two sets of image simulations that mirror the observed properties of CFHTLenS have been created to establish the method's accuracy and to derive an empirical correction for the effects of noise bias.

Key words: gravitational lensing: weak – methods: data analysis – methods: statistical – cosmology: observations.

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1 INTRODUCTION

Weak gravitational lensing allows measurement of the matter content of the Universe, which, when used in combination with other probes, should enable tests of dark energy models of the Universe and tests of general relativity on cosmological scales. However, despite the detection of weak lensing around galaxy clusters more than 20 years ago (e.g. Tyson, Wenk & Valdes 1990) and the detection of cosmological large-scale lensing (‘cosmic shear’) more than 10 years ago (Bacon, Refregier & Ellis 2000; van Waerbeke et al. 2000; Wittman et al. 2000), it has proved difficult to make measurements that are free from systematic error at the level required to measure cosmological parameters accurately. To date, imaging data obtained from the *Hubble Space Telescope* (*HST*) have been the most successful in this respect, but have yielded accurate cosmological results only relatively recently (e.g. Schrabback et al. 2010). Extension of these studies to large sky areas is currently only feasible from the ground, but then the weak lensing measurements need to be made using images of faint galaxies that are only marginally resolved and for which the lensing signal needs to be disentangled from the effects of an image point spread function (PSF) which varies both with position within an image and between images.

Fig. 1 shows the sizes of disc galaxies in the ‘Groth Strip’, as measured by Simard et al. (2002) using data from the *HST* Wide Field Planetary Camera 2 (WFPC2). These data are used in Appendix B to obtain an estimate of the Bayesian prior distribution of galaxy size used in this paper; however it is useful to show the distribution here to illustrate the level of difficulty imposed by ground-based measurements. The figure shows the fitted values of semi-major axis disc exponential scalelength obtained for galaxies which had fitted bulge-to-total flux ratios (B/T) less than 0.5.¹ We have measured the median values of that scalelength in bins of *i*-band (WFPC2 filter F814) apparent magnitude, excluding values $r_d < 0.02$ arcsec, and these are shown as large filled circles in Fig. 1. In this paper, we describe how we have made weak lensing measurements using ground-based data from the Canada–France–Hawaii Telescope (CFHT) MegaCam camera (Boulade et al. 2003) obtained for the CFHT Legacy Survey (CFHTLS) and analysed as the CFHT Lensing Survey (CFHTLenS): Fig. 1 also shows the MegaCam typical pixel size of 0.187 arcsec and a typical CFHTLS PSF half-width half-maximum (HWHM) of 0.35 arcsec. It can be seen that at the faintest magnitudes (CFHTLenS analysed galaxies to $i'_{AB} < 24.7$) the median semi-major axis exponential scalelength of galaxies is comparable to the pixel size, and that very few galaxies have scalelengths that are larger than the PSF HWHM. Given that the faint galaxies are also noisy, with the faintest galaxies having signal-to-noise ratio as low as 7, it is clear that highly accurate methods need to be adopted for measuring the systematic distortions of galaxies at the level of a few per cent in ellipticity that are expected from weak lensing.

We describe in this paper a Bayesian model-fitting method aimed at measuring the shapes of faint galaxies in optical galaxy surveys. Current weak lensing surveys such as CFHTLenS comprise about 10^7 galaxies, so one design aim of the algorithms is that they should be fast. However, as survey sizes increase, the requirements on acceptable levels of non-cosmological systematics signals become

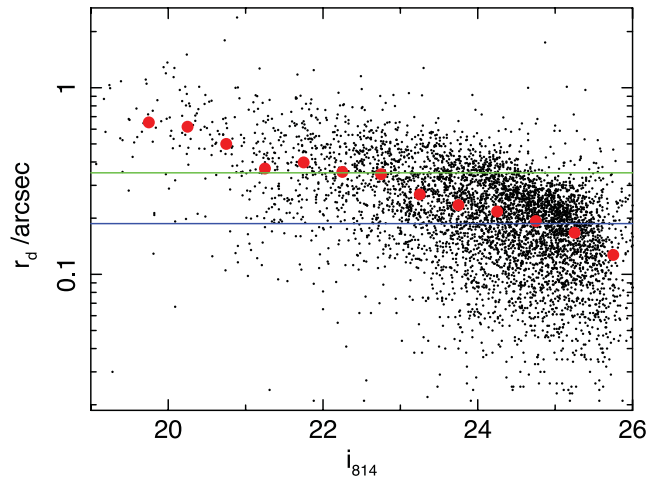


Figure 1. Fitted disc semi-major axis scalelength, plotted as a function of i' -band magnitude, for galaxies identified as being disc-dominated, from the fits of Simard et al. (2002) to *HST* WFPC2 data. The large filled circles indicate the median size measured in bins of apparent magnitude. The horizontal lines indicate the CFHT MegaCam pixel size of 0.187 arcsec (lower line) and a typical CFHTLenS PSF HWHM of 0.35 arcsec (upper line).

more stringent, meaning that the new generation of measurement algorithms must be simultaneously faster and more accurate than previous methods. The principles that were adopted for measuring galaxy ellipticity are described by Miller et al. (2007, hereafter Paper I), and implemented in the ‘*lensfit*’ algorithm. Tests on simulated images of a previous implementation of the basic shear-measurement algorithm are described by Kitching et al. (2008, Paper II).

This paper describes a number of significant improvements to the algorithm which proved necessary for working on data from a real survey, rather than simplified simulations.

- (i) PSF modelling took account of spatial variation of the PSF and possible discontinuities between CCDs in the mosaic camera (MegaCam, Boulade et al. 2003, in the case of CFHTLenS).
- (ii) Galaxy measurement used multiple, dithered exposures, optimally combined in the measurement process, taking account of pointing and camera distortion variations and varying image quality.
- (iii) Camera distortion was measured directly from images and corrected in the measurement process.
- (iv) Fitted galaxy models allowed joint fitting of bulge and disc components.
- (v) Full Bayesian marginalization over nuisance parameters was carried out.
- (vi) Additional measures were used to account for bad and saturated pixels and columns, cosmic rays, image ghosts, blended stars and galaxies and recognition of morphologically complex galaxies.

Information on these various improvements is described below. Section 2 describes CFHTLenS; Sections 3–6 provide details of the method and the shear measurement pipeline; Section 7 provides an overview of tests for systematic errors, which are described in more detail by Heymans et al. (2012), and Section 8 describes extensive image simulations that have been used to test the method and to calibrate the effects of noise bias. Appendices A and B provide additional details of the models and priors used in the method, and a mathematical illustration of the effect of noise bias is given in Appendix C.

¹ A population of objects that were fitted by Simard et al. (2002) with scalelengths less than 0.02 arcsec have been excluded: these very small scalelengths are not correlated with apparent magnitude and are likely fits to stars or are otherwise spurious measurements, as noted by Simard et al.

2 THE CANADA–FRANCE–HAWAII TELESCOPE LENSING SURVEY

The data used in this paper are from the CFHT Wide Legacy Survey,² analysed by the CFHTLenS team.³ The CFHTLS Wide consisted of 171 pointings covering a sky area of 154 deg² in five bands, $u^*g'r'i'z'$, with the MegaCam camera (Boulade et al. 2003), which has a field of view of approximately 1 deg². The pointings were distributed over four separate patches on the sky, and each patch was contiguous, covering an area of 63.8, 22.6, 44.2 and 23.3 deg², respectively for the W1, W2, W3 and W4 patches. Each pointing comprised typically seven i' -band exposures, often with many of those taken during the same night. In some fields, one or two exposures were judged to be of too poor quality, based on the seeing or apparent degree of PSF variation, reducing the number of usable exposures to 5 or 6. In some other fields, two entire sets of exposures were obtained if there was concern about the quality of the first set, but in a number of cases all 14 exposures were subsequently judged to be usable and were included in the CFHTLenS analysis. Exposures were dithered by up to 90 arcsec in declination and 24 arcsec in right ascension in patches W1, W3 and W4, and by up to 400 arcsec in W2.

The CFHTLenS analysis of the data was designed and optimized for weak lensing science, both photometry and shape measurements. The photometric analysis of the survey is described by Hildebrandt et al. (2012) and Erben et al. (2012). Galaxy shapes have been measured on the i' band which reaches a 5σ point source limiting magnitude of $i'_{AB} \sim 25.5$. A magnitude limit of $i'_{AB} < 24.7$ was imposed for shape measurement, noting that photometric redshift accuracy also is poor at fainter magnitudes than this limit (Hildebrandt et al. 2012). Exposures obtained prior to 2007 June were obtained with the i' -band filter MP9701. After 2007 October an i' -band filter with a slighter different spectral response, MP9702, was used.

All fields had a full-width half-maximum (FWHM) seeing condition which was required to be better than 0.8 arcsec. Unlike previous lensing analyses, galaxy shape measurement was performed on the individual exposures, photometrically calibrated with Elixir (Magnier & Cuillandre 2004). The CFHTLenS data were then processed with the *THELI* pipeline (see Erben et al. 2005). The specific algorithms and procedures were described by Erben et al. (2009). In that work, the astrometric and photometric calibration was made independently in each 1 deg² MegaPrime pointing. The most important refinement to the calibration for CFHTLenS was a simultaneous treatment of all data from a CFHTLenS patch. This typically led to an improvement of a factor of 2–3 in astrometric and photometric accuracy compared with the earlier analysis. Further details of the construction of the survey are given by Heymans et al. (2012) and Erben et al. (2012).

3 THE SHEAR MEASUREMENT METHOD

3.1 Principles

Key points to emphasize about the approach of Bayesian model fitting are as follows.

(i) It enables a rigorous statistical framework to be established for the measurement, resulting in optimally weighted use of the

available data, regardless of variations in signal-to-noise ratio or PSF.

(ii) Quantities other than ellipticity that form part of the problem may be marginalized over and, provided an accurate prior is available, those uninteresting quantities may be eliminated from the problem without bias. The prime example is that of the unknown galaxy position: all other methods published to date require a fixed centroid position for a galaxy, and centroid errors introduce spurious ellipticity, with the possibility of bias (e.g. Bernstein & Jarvis 2002). In model fitting we can treat the unknown galaxy position as a free parameter and marginalize.

(iii) The weak lensing signal is carried by the faintest galaxies; often the integrated signal-to-noise ratio may be as low as 10. A model-fitting approach adds extra information to the problem, such as prior knowledge of possible surface brightness distributions. In principle, this leads to more precise measurements than a method that does not use such information, thereby allowing measurements to lower signal-to-noise ratio.

(iv) The models must be unbiased, however, in the sense that they should span the range of surface brightness profiles of real galaxies, and the priors adopted for the parameters of those models should be an accurate reflection of the true distribution. If these conditions are not satisfied, the model fitting may be systematically biased (Bernstein 2010; Voigt & Bridle 2010).

(v) A full Bayesian approach should be able to correct for the effects of noise bias, and hence avoid the need for any independent calibration. We shall see in Section 3.5 that we have not yet achieved that goal.

Point (iv) is worth considering in more detail. Bernstein (2010) has argued that model basis functions need to be truncated to some finite set to avoid the fitting results being dominated by noise, and that truncation leads to bias (see also Ngan et al. 2009). Thus, even when specific physical models are not assumed, the use of any so-called ‘model-independent’ basis set also results in biased results, unless the truncated basis set employed is capable of accurately modelling the full observed range of surface brightness profiles. In methods that use simple statistical measures of the data, such as moments (e.g. Kaiser, Squires & Broadhurst 1995, hereafter KSB, or Melchior et al. 2011) it is necessary to use weighted moments because of image noise, and the use of a finite set of weighted moments can lead to a bias similar to the use of truncated basis sets. Thus, all methods that aim to measure galaxy shapes from noisy, PSF-convolved data have the risk of producing biased results, and choosing a so-called ‘model-independent’ method does not alleviate this concern (see also Melchior & Viola 2012). The approach adopted here is that if the model basis sets are a close match to the true galaxy surface brightness distributions, and in particular if the PSF-convolved model distributions span the range of true PSF-convolved distributions, then it should be possible to avoid model truncation bias. By choosing model surface brightness distributions that match the bulge and disc components of galaxies at low redshift, we are making a choice of a basis set that should minimize bias. That does leave open the question of whether such models can span the range of galaxy surface brightness distributions at high redshift, given the expected morphological evolution of the galaxy population, but whether or not the fitting of composite disc and bulge models to galaxies at high redshift results in bias has not yet been established. This concern should be explored in future work, especially given the ever-increasing accuracy required for future surveys.

² www.cfht.hawaii.edu/Science/CFHTLS

³ www.cfhtlens.org

The models fitted in CFHTLenS comprised two-component bulge plus disc galaxies (Appendix A). There were seven free parameters: galaxy ellipticity, e_1 , e_2 ; galaxy position x , y ; galaxy size r ; galaxy flux; and bulge fraction. Ultimately we are interested only in the ellipticity, so the other parameters were regarded as nuisance parameters and marginalized over. Bayesian marginalization was used, adopting prior distributions for each parameter, which are described in Appendix B. With the given priors, each parameter was marginalized over either analytically, if possible, or otherwise numerically. The bulge fraction marginalization was analytic, galaxy flux and size were numerically marginalized over by sampling multiple values, fitting a function and integrating. Galaxy position was numerically marginalized over using the Fourier method described in Paper I: the process of model fitting is equivalent to cross-correlation of the model with the data, and marginalization over unknown position is equivalent to integrating under the peak of the cross-correlation function. Thus the likelihood surface was obtained as a function of only the apparent galaxy two-component ellipticity. Finally, that likelihood surface was converted to weighted estimates of ellipticity for use in shear measurement (Section 3.5). The following sections give more detail of the galaxy models and of the calculation of the likelihood function and its marginalization.

3.2 Galaxy models

Adopting the same conventions as in Paper I, observed galaxy ellipticity e is related to the intrinsic galaxy ellipticity e^s by

$$e = \frac{e^s + g}{1 + g^* e^s} \quad (1)$$

(Schramm & Kayser 1995; Seitz & Schneider 1997), where e and g are represented as complex variables, g , g^* are the reduced shear and its complex conjugate, respectively, and e is defined in terms of the major and minor axes and orientation a , b , θ , respectively, as $e = (a - b)/(a + b) \exp(2i\theta)$. In this formalism, we expect

$$\langle e \rangle = g \quad (2)$$

for an unbiased sample with $g < 1$, where $\langle e^s \rangle = 0$. Note that this result differs from the other commonly used formalism which instead uses ellipticity, or polarization, defined as $e = (a^2 - b^2)/(a^2 + b^2) \exp(2i\theta)$, and which requires measurement of the distribution of galaxy ellipticities in order to calibrate the response to lensing shear (e.g. Bernstein & Jarvis 2002).

The model-fitting approach to shear measurement of Paper I was adopted, but extended to include multiple galaxy components, which we assumed to comprise a disc and bulge component. In reality, galaxies may have more complex morphologies than are describable by this simple phenomenological model. In non-Bayesian methods, models of greater complexity than required by the data lead to ‘overfitting’ and its associated biases (e.g. Bernstein 2010). In principle, a Bayesian method could include additional components and model parameters, provided that prior probability distributions are available, so that parameters not of interest to the shear measurement may be marginalized over. In practice, we wish to avoid using models of greater complexity than are required by the data, particularly as more complex models increase the computational time required to make the measurements. In CFHTLenS, early experiments with the systematic tests described in Section 7 and by Heymans et al. (2012) indicated that significant improvements could be obtained by extending the single-component model of Paper I to a simple two-component galaxy model in which only one additional free parameter was allowed.

The generation of the galaxy models is described in Appendix A. We assumed that a galaxy survey may contain two fundamental galaxy types. A fraction f_B were bulge-dominated galaxies with a single morphological component, which was assumed to be described by a de Vaucouleurs profile (Sérsic index 4). A fraction $(1 - f_B)$ were disc galaxies which comprised both a bulge component and a disc component. The disc component was assumed to be a pure exponential, Sérsic index 1, with the bulge component again having Sérsic index 4. The model components were truncated in surface brightness at a major axis radius of 4.5 exponential scale-lengths (disc component) or 4.5 half-light radii (bulge component), which was both computationally convenient and also in general agreement with observations of disc galaxies (e.g. van der Kruit & Searle 1982).

To avoid prohibitive computational cost by introducing another model parameter, which would not be analytically marginalizable, the half-light radius of the bulge was set equal to the exponential scalelength of the disc for the composite disc-dominated galaxies (i.e. the bulge half-light radius was 0.6 times the value of the disc half-light radius). In reality, galaxies display a wide range in their ratios of the sizes of these components, although that distribution itself depends on the assumed surface brightness profiles adopted (Graham & Worley 2008). The adopted ratio of bulge half-light radius to exponential scalelength is towards the upper end of the range encountered in the literature. As the CFHTLenS galaxies are generally only marginally resolved, smaller bulges would have half-light radii significantly smaller than the pixel scale, and the final PSF-convolved models were largely insensitive to the choice of this ratio, although we did find that significantly smaller values of bulge half-light radius led to larger PSF–galaxy cross-correlations (see Section 7) for the smallest fitted galaxies, and were therefore disfavoured.

Given the low signal-to-noise ratio and the small number of pixels associated with the faintest galaxies in the CFHTLenS data, a more complex set of galaxy models, with more free parameters, was not justified. We might ask, however, whether the particular models chosen are the ‘best’, in the sense of being the least biased, or whether some other choice might be better. To answer such a question, we need to have a good understanding of the actual population of galaxies being measured. Recently, attempts have been made to create realistic simulations of the galaxy population based on high-resolution *HST* observations (Mandelbaum et al. 2012), and these form the basis of the GREAT3 challenge simulations.⁴ In the current paper, we take the view that the bulge-plus-disc parametrization is known to fit well to the local galaxy population, and therefore the surface brightness distributions of such models are already a necessary ingredient of any successful models. Even though high-redshift galaxies tend to have more complex morphologies than local galaxies, the effects of convolution with the PSF may render such differences unimportant in ground-based data of CFHTLenS quality, where the high-redshift galaxies are only marginally resolved. We discuss in Section 8.3 the possible amplitude of biases that might arise from inappropriate model choice. The requirements for larger, future surveys, with correspondingly tighter requirements on the acceptable level of systematic biases, will need to be established through more extensive simulations.

⁴ great3challenge.info

3.3 The likelihood function and marginalization over nuisance parameters

We define the likelihood \mathcal{L} and the χ^2 goodness-of-fit statistic as

$$\chi^2 \equiv -2 \log \mathcal{L} = \sum_i \left[\frac{y_i - SBg_i - S(1-B)f_i}{\sigma_i} \right]^2, \quad (3)$$

where y_i is the image data value of pixel i , σ_i is the statistical uncertainty of that data value, f_i is the disc component model value for that pixel, convolved with the PSF, g_i is the bulge component model amplitude, convolved with the PSF, B is the bulge fraction defined as the ratio of the flux in the bulge component to the total flux (often written as B/T) and S is the total galaxy flux. The pixel noise is assumed to be stationary and uncorrelated, which is appropriate for shot noise in CCD detectors in the sky-noise limit. Bright galaxies may make a significant contribution to photon shot noise, but the resulting non-stationary noise would then invalidate the *lensfit* Fourier-space approach which is described in Paper I and used below. Hence this algorithm is appropriate for model fitting to faint galaxies in the sky-limited regime. The method can easily be generalized to the case where the noise is stationary but correlated between pixels, but we do not consider that generalization here.

The prior distribution of bulge fraction B for the disc galaxy population was assumed to be a truncated Gaussian distribution,

$$\mathcal{P}(B) = \mathcal{P}_0 \exp \left[-(B - B_0)^2 / (2\sigma_B^2) \right] \quad 0 \leq B < 1, \quad (4)$$

where \mathcal{P}_0 was a normalizing factor and B_0 and σ_B were parameters of the prior distribution. The Gaussian form of the prior was a convenient choice to allow easy marginalization over B , given the quadratic dependence of $\log \mathcal{L}$ on B . We adopted $B_0 = 0$, $\sigma_B = 0.1$, $f_B = 0.1$ as a reasonable representation of the distribution of bulge fraction found at low and high redshift (e.g. Schade et al. 1996; Simard et al. 2002). We marginalized over the Gaussian part of the prior by collecting terms in equation (3) and multiplying by the prior, writing the log of the posterior probability density for the disc population, $\log P_{\text{disc}}$, as

$$-2 \log P_{\text{disc}} = \sum y^2/\sigma^2 + 2B^2\theta^2 + 2Br_0 + S^2 \sum f^2/\sigma^2 - 2S \sum fy/\sigma^2 + B_0^2/\sigma_B^2 \quad (5)$$

where

$$2\theta^2 = S^2 \left[\sum f^2/\sigma^2 + \sum g^2/\sigma^2 - 2 \sum fg/\sigma^2 \right] + 1/\sigma_B^2,$$

$$r_0 = S \left(\sum fy/\sigma^2 - \sum gy/\sigma^2 \right) + S^2 \left(\sum fg/\sigma^2 - \sum f^2/\sigma^2 \right) - B_0/\sigma_B^2,$$

and where the summations are understood to be over the pixels i . We marginalized over B to obtain the log of the marginalized disc posterior probability density as

$$-2 \log P_{\text{disc}} = S^2 \sum f^2/\sigma^2 - 2S \sum fy/\sigma^2 - \frac{r_0^2}{2\theta^2} + 2 \log \theta - 2 \log \left[\text{erf} \left(\theta + \frac{r_0}{2\theta} \right) - \text{erf} \left(\frac{r_0}{2\theta} \right) \right], \quad (6)$$

neglecting any normalization terms that do not depend on a galaxy's model parameters.

The bulge-fraction-marginalized total probability $P = (1 - f_B)P_{\text{disc}} + f_BP_{\text{bulge}}$ then needed to be marginalized over both flux and position. To marginalize over galaxy position at a given model galaxy flux, we followed an improved version of the method described in Paper I. For stationary noise, the terms $\sum f^2/\sigma^2$ and

$\sum g^2/\sigma^2$ are invariant under shifts of the model position, and following Paper I, the terms $\sum fy/\sigma^2$ and $\sum gy/\sigma^2$ may be regarded as the spatial cross-correlation of each model component with the data, so we may write

$$h_f(\mathbf{x}) = \sum fy/\sigma^2, \quad h_g(\mathbf{x}) = \sum gy/\sigma^2, \quad (7)$$

where the cross-correlations h_f or h_g depend on the vector position shift \mathbf{x} of the model. As pointed out in Paper I, to marginalize over \mathbf{x} we need to adopt a prior $\mathcal{P}(\mathbf{x})$: a priori, the position of the galaxy is unknown, and is equally likely to exist at any position on the sky, so we adopt a uniform prior. Any other choice would lead to a bias in the position of the galaxy, and hence a bias in the measured ellipticity, arising from the correlation between fitted ellipticity and position. However, as discussed in Paper I, $\mathcal{L} \rightarrow \text{constant}$ as $|\mathbf{x}| \rightarrow \infty$ and the marginalized likelihood would not be finite. This problem arises because, no matter how large a pixel value, it always has a finite chance of being due to random noise, with the true galaxy being positioned elsewhere. At large position offsets, $\mathbf{x} \rightarrow \infty$, $h_f, h_g \rightarrow 0$ and the posterior probability density tends towards a 'pedestal' value $P_0 \equiv (1 - f_B)P_{0,\text{disc}} + f_BP_{0,\text{bulge}}$, where

$$\begin{aligned} -2 \log P_{0,\text{disc}} &= S^2 \sum f^2/\sigma^2 - \frac{r_0'^2}{2\theta^2} + 2 \log \theta \\ &\quad - 2 \log \left[\text{erf} \left(\theta + \frac{r_0'}{2\theta} \right) - \text{erf} \left(\frac{r_0'}{2\theta} \right) \right], \\ r_0' &= S^2 \left(\sum fg/\sigma^2 - \sum f^2/\sigma^2 \right) - B_0/\sigma_B^2, \end{aligned}$$

$$-2 \log P_{0,\text{bulge}} = S^2 \sum g^2/\sigma^2.$$

Thus we need to identify a maximum bound for the position offset over which to marginalize.

As argued in Paper I, galaxies generally have smooth centrally concentrated surface brightness distributions which are convolved with centrally concentrated PSFs in an observed image (the assumption of a smooth convolved galaxy profile is a good approximation for galaxies that are not well resolved). The model is also smooth, centrally concentrated and convolved with the same PSF. As in the derivation of the central limit theorem, such a cross-correlation should be well represented by a two-dimensional Gaussian distribution, and hence we approximated $\log P$ as being a 2D Gaussian added to the pedestal value $\log P_0$. To marginalize over the galaxy's position, we adopted a more exact integration of this form than assumed in Paper I, and integrated out to some maximum position uncertainty, r_{max} . To establish the value of r_{max} , we considered that the only information that we have about the galaxy is on the images being analysed: the galaxy probably was initially detected, and its position measured, with some uncertainty, from the same data being used to measure the shape. We are seeking to use the data in this region of the image to measure the shape of a galaxy that exists at this location, not of any arbitrary galaxy anywhere else in the image: although the likelihood is non-zero at large position offsets, such a galaxy would not correspond to the one which had been detected. Accordingly we set r_{max} to be the position beyond which the detection of this galaxy becomes statistically insignificant, which we define to be given by $\Delta \log P = P(r_{\text{max}}) - P_0 < \Delta \chi_{\text{crit}}^2/2$ and where we chose $\Delta \chi_{\text{crit}}^2 = 6$, corresponding to the 95 per cent confidence region for the location of the galaxy. The integral is given by an exponential integral which was evaluated numerically using the GNU Scientific Library (Galassi et al. 2009), rather than using the approximation of Paper I.

In Paper I, the flux marginalization was carried out analytically by noting the Gaussian dependence of P on S . For the CFHTLenS

analysis, it was decided to adopt a prior distribution for S which was given by a power law in S : $\mathcal{P}(S) \propto S^{-\alpha}$ where $\alpha = 2$ was chosen to match the flatter-than-Euclidean slope of the faint i -band galaxy number counts (e.g. Gabasch et al. 2008). The resulting posterior probability distribution cannot then be analytically marginalized over S , so this was achieved numerically, by evaluating P at a set of points close to the expected maximum likelihood, fitting a cubic spline function to the variation of $\log P$ with S and numerically integrating the spline-fitted values of P .

After these marginalization steps, the marginalized likelihood had been reduced to being a function of ellipticity and galaxy scale-length alone. The final step was to adopt a prior for the scalelength distribution (described in Appendix B) and to numerically marginalize over this parameter. As for the flux marginalization, $\log P$ was measured for a set of values of scalelength r , a cubic spline fitted to those values, and P was numerically integrated with respect to r . The set of r values measured included a point source, $r = 0$, and it was ensured that the set of $\log P$ values to be fitted varied smoothly towards $r = 0$. To distinguish between stars and galaxies, objects were classified using an F-ratio test of the ratio of the chi-squared values of the best-fitting galaxy ($r > 0$) and the best-fitting star ($r = 0$) models, taking into account the differing numbers of degrees of freedom of these two fits. A star classification was adopted if the probability of the F-ratio exceeded a magnitude-dependent limit that was matched to visual classification of objects.⁵ Of the objects that were fitted, typically 95 per cent were classified as galaxies, as expected at faint magnitudes.

3.4 Sampling the likelihood surface

In order to obtain unbiased estimates of shear, and to measure the statistical uncertainty in ellipticity for each galaxy, we need to not only find the model ellipticity whose marginalized probability is a maximum, but also we need to sample around the maximum to measure the likelihood surface, at least in regions where the probability is not insignificant. The ellipticity surface is two-dimensional, and bounded by the requirement $|e| < 1$, so it is relatively straightforward to sample the surface appropriately.

However, galaxies of low signal-to-noise ratio or small size are expected to have broad likelihood surfaces, whereas high signal-to-noise ratio galaxies should have sharp, well-defined surfaces. Thus the sampling needed depends on the properties of each galaxy. The algorithm adopted here was an adaptive grid sampling, in which ellipticity measurements were first made on a coarse grid of ellipticity, at intervals of 0.16 in ellipticity. At each sampled value, a full marginalization over the other parameters was made. Then, the algorithm counted the number of sampled points above a threshold of 5 per cent of the maximum posterior probability, and if that number was below a minimum value of 30 points, the sampling in ellipticity was reduced by a factor of 2, and the region around the maximum, only, sampled at the higher resolution, until either 30 points had been measured above the threshold, or until a resolution of 0.02 in ellipticity was reached (for the faint galaxies that dominate the weak lensing signal this limiting resolution in ellipticity

was adequate, and is much smaller than the ‘shape noise’ caused by the broad distribution of intrinsic ellipticity).

In a later improvement, two passes were made through the ellipticity surface, the first to find only the maximum using the adaptive sampling algorithm, and the second to map out the surrounding region. This algorithm was found to be more efficient (in that the fraction of sampled points that were retained above the threshold was higher) but was not implemented in time for the CFHTLenS analysis reported here. Monte Carlo Markov Chain (MCMC) methods could also be used, but given the low dimensionality and finite bounds of the parameter space, are unlikely to be significantly more efficient than this adaptive sampling algorithm.

3.5 Shear estimation and noise bias

Papers I and II discussed at length the advantages of Bayesian estimation, but also noted that a bias would arise in shear estimates if the likelihood distribution for each galaxy were multiplied by an ellipticity prior before the estimation process. The existence of that bias demonstrates that a Bayesian method of measuring the *ellipticity* of an individual galaxy is not in itself a complete method for measuring *shear* in a sample of galaxies. A solution proposed at that time was to calculate the ‘sensitivity’, which measured the biasing effect on a shear estimator of applying an ellipticity prior. In Paper I, the sensitivity was defined as $s \equiv \partial \langle e_\alpha \rangle / \partial g_\alpha$, a calibration of how the mean ellipticity of a sample $\langle e_\alpha \rangle$, for ellipticity component α , responds to a cosmological reduced shear g_α in that component. Cross-terms were ignored. Analogous approaches to the sensitivity, albeit not in this Bayesian context, have been discussed by Bernstein & Jarvis (2002) and Kaiser (2000).

An immediate difficulty with this approach is that the sensitivity should be a rank 2 tensor, but in fact even if only the diagonal terms are considered, the sensitivity is not isotropic, and application of sensitivity as a correction can lead to an orientation bias in the measured shear. As the sensitivity is calculated from the measured likelihood surface, and as those surfaces are biased (Appendix C), poor results were obtained in CFHTLenS when this approach was used.

An improved alternative, used for CFHTLenS, is a likelihood-based estimator of ellipticity, in which all nuisance parameters were marginalized over using a fully Bayesian approach as previously described, but where the ellipticity estimate is based solely on the likelihood distribution rather than a posterior probability distribution. It was argued in Paper I that such an estimator should be an unbiased estimator of shear, even though the statistical distribution of galaxy ellipticities is broadened by the measurement noise: i.e. the likelihood estimator has no correction for noise, but should respond linearly to a cosmological shear, when averaged over many galaxies (when ellipticity is defined as in Section 3.2). However, the non-linearity of the transformation from a pixel basis to ellipticity does result in a ‘noise bias’ in shear estimates (Melchior & Viola 2012; Refregier et al. 2012) which is discussed below.

We should also consider which estimator to use for the ellipticity value of each galaxy. If we consider the (simplified) shear measurement process of averaging galaxy ellipticities, where each galaxy has its own PDF of ellipticity, the PDF of the mean is given by the convolution of the individual PDFs, suitably rescaled. According to the central limit theorem, the PDF of the mean should be Gaussian, with expectation value given by the mean of the individual galaxy PDF means. Thus to make a fully probabilistic estimator of shear, we should use the mean ellipticity as our estimate for each galaxy, not the maximum likelihood estimate. Unfortunately, such an

⁵ A Bayesian procedure that took account of the relative probabilities of stars and galaxies as a function of magnitude could avoid needing to specify the value of the F-ratio statistic; however given the very small numbers of stars expected in the faint galaxy, weak lensing, regime, $i \gtrsim 23$, this was not deemed to be an important improvement in this survey.

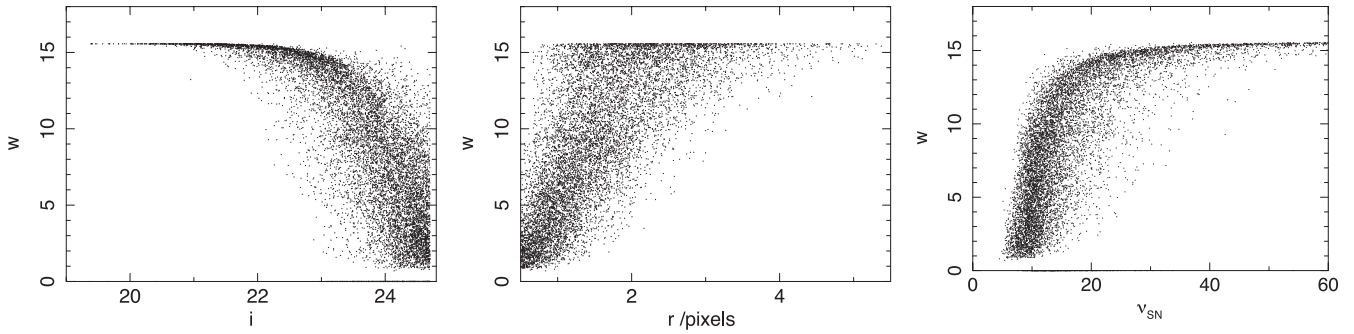


Figure 2. An example of the distribution of lensing weights for 10 000 measured galaxies in one typical CFHTLenS field, W1m2m3. Galaxies were measured on seven exposures to a limiting magnitude $i' < 24.7$. Weights are shown as a function of (left) apparent magnitude i , (centre) fitted semi-major axis disc scalelength r and (right) signal-to-noise ratio v_{SN} .

estimator is biased, as discussed in Appendix C. For the case where the galaxy size is known precisely, the maximum likelihood ellipticity should be unbiased and may be a better estimator to use, but this is not true if the galaxy's size is a free parameter: there is an inevitable degeneracy in the fits, between size and ellipticity, which causes a further bias in the likelihood surfaces (see also Melchior & Viola 2012; Refregier et al. 2012). Unlike the case of a pure likelihood estimator, Bayesian marginalization over the size parameter mitigates this effect, because small galaxy sizes are downweighted by the size prior, but it does not eliminate it. A fully Bayesian estimate of ellipticity should control the noise bias, but currently we do not have a satisfactory method of correcting for the bias on the shear introduced by the application of an ellipticity prior discussed at the start of this subsection.

To partially address the problem of not being able to use a Bayesian estimate for the ellipticity, we downweighted likelihood values at ellipticities where the Bayesian posterior probability was low. We formed a posterior probability surface by multiplying by the ellipticity prior, and rejected any posterior probability values less than some threshold. The threshold was set to 1 per cent of the maximum posterior probability. Applying our prior knowledge of the ellipticity distribution in this way allowed consistent rejection of the ‘small, highly elliptical’ tail on the likelihood surface described above. The galaxy's ellipticity was then estimated from the mean of the *likelihood* distribution of the ellipticity samples that were above the posterior threshold. This somewhat ad hoc approach is not a substitute for a fully Bayesian approach, but it did reduce the residual systematic signals discussed in Section 7. We still expect bias in the estimator at low signal-to-noise ratio, although the lowest signal-to-noise ratio galaxies also had low weights in the analysis. An empirical correction for the noise bias was derived from the simulations described in Section 8. One might hope that, in future, a better ellipticity estimator could be constructed from the likelihood or Bayesian posterior probability surfaces which is immune to noise bias and which leads to an unbiased estimate of shear. To date, we have not achieved such an estimator, but have found an approach that yields satisfactory results for CFHTLenS.

3.6 Galaxy weighting

Faint galaxies have more noisy ellipticity measures than bright galaxies, or, equivalently, have broader likelihood surfaces. Given that the full likelihood surfaces for the galaxies' ellipticities are not used in making cosmological shear estimates, each galaxy should be weighted according to the width of its likelihood surface in the

cosmological analysis. We calculated an inverse-variance weight, defined as

$$w = \left[\frac{\sigma_e^2 e_{\text{max}}^2}{e_{\text{max}}^2 - 2\sigma_e^2} + \sigma_{\text{pop}}^2 \right]^{-1} \quad (8)$$

where σ_e^2 is the 1D variance in ellipticity of the likelihood surface, σ_{pop}^2 is the 1D variance of the distribution of ellipticity of the galaxy population and e_{max} is the maximum allowed ellipticity, assumed here to be the maximum disc ellipticity as in equation (B2). In the limit where $e_{\text{max}} \rightarrow \infty$, this definition of the weight tends towards a conventional form, $w \rightarrow (\sigma_e^2 + \sigma_{\text{pop}}^2)^{-1}$. A finite value of e_{max} was included to ensure that a galaxy with a likelihood surface which is uniform for $e < e_{\text{max}}$, and which therefore conveys no information about the galaxy shape, has zero weight. For a Gaussian likelihood surface and prior (and in the limit $e_{\text{max}} \rightarrow \infty$), w has the same value as the sensitivity (Section 3.5), except for its normalization (Paper I, Section 2.5). The expressions differ in the (actual) case of non-Gaussian probability surfaces, but w has the advantage that it is not by definition anisotropic. In the limit of low-noise measurements, the weight becomes dominated by the ‘shape noise’ of the galaxy population, as expected for weak-lensing studies (Bernstein & Jarvis 2002). Fig. 2 shows an example of the distribution of weights in one of the CFHTLenS fields, as a function of i' -band magnitude, fitted galaxy major-axis scalelength, and signal-to-noise ratio, v_{SN} . The weight falls off sharply for $v_{\text{SN}} \lesssim 20$. The maximum value of the weight is determined by the value of $\sigma_{\text{pop}} = 0.255$, obtained from the ellipticity prior described in Appendix B.

However, although the weight is designed to be isotropic, it is possible for there to exist a form of ‘orientation bias’ similar to the selection bias discussed by Kaiser (2000) and Bernstein & Jarvis (2002). In the case of the selection bias, we are concerned about galaxies being omitted from the sample if they are cross-aligned with the PSF. However, even for detected galaxies, the weight also may be larger for galaxies that are aligned with the PSF than for those that are cross-aligned, because the likelihood surface may be more sharply peaked. We did not make any attempt to correct explicitly for this effect, but it may manifest itself as a correlation between weighted galaxy shapes and the PSF, as shown in Section 8.4.

3.7 The treatment of large, blended or complex galaxies

The galaxy models were fitted to each galaxy in subregions extracted around each galaxy (‘postage stamps’) of size 48 pixels (approximately 9 arcsec), this choice being a compromise between having a large enough region to correctly model the galaxies but small enough that the compute time for operations such as the fast

Fourier transform was not prohibitively slow. If each galaxy had n useful exposures, then n postage stamps were created and models jointly fitted, allowing for astrometric offsets and camera distortion as described in Sections 4 and 6. Inevitably, some galaxies had sizes too large to be fitted in this size of postage stamp; such galaxies were excluded from the analysis.

In some cases, two or more neighbouring galaxies appeared within the same postage stamp. The algorithm can only fit one galaxy at a time, so the solution adopted was to first see whether it was possible to mask out one galaxy (set its pixel values equal to the background) without disturbing the isophotes of the galaxy being fitted. To this end, a co-added image postage stamp was created, averaging all the exposures available for that galaxy, shifted so the relative positions agreed to the nearest pixel, which was then smoothed by a Gaussian of FWHM equal to that of the local PSF. Isophotes were created for each smoothed galaxy: if a separate galaxy or other object was identified with non-touching isophotes, at a level of twice the smoothed pixel noise, that other galaxy was masked out and the fitting would proceed. Such close pairs of galaxies are thus included in the output catalogues from CFHTLenS. We note, however, that low-level light leaking below the two-sigma isophote could still contaminate the measurement, and thus we expect the ellipticity measurements of galaxies in close pairs, whose isophotes may be contaminated by their neighbour, to be artificially correlated.

Within each postage stamp, it may be that some pixels should be masked because of image defects. The THELI pipeline provided images of pixel masks to be applied. If such masked pixels occurred within the two-sigma isophote of a galaxy on one individual exposure, that exposure was not used in the joint analysis. If such pixels occurred outside the two-sigma isophote, the pixel values were set equal to the background and that masked exposure was used in the joint fitting.

Other galaxies may be sufficiently close that their smoothed isophotes overlapped, and there may also be individual galaxies with complex morphology, not well described by a simple bulge-plus-disc model. These galaxies were identified using a deblending algorithm, testing for the presence of significant independent maxima in the smoothed surface brightness distribution.⁶ Any such complex or blended galaxies that were found were excluded from the analysis. A further criterion was imposed, that the intensity-weighted centroid of a galaxy, measured from the pixels within the smoothed 2σ isophote, should lie within 4 pixels of the nominal target position: this criterion guarded against any blended galaxies that had been identified as blends in the original input catalogue but that had not been identified by the other tests described in this section. Some examples of images of galaxies excluded by these criteria are shown in Fig. 3, which shows examples of the stacked, smoothed images used for testing for object complexity. Visual inspection indicated that the great majority of galaxies excluded in this way had isophotes that overlapped with neighbouring galaxies.

⁶The algorithm was similar to that of Beard, MacGillivray & Thanisch (1990). Maxima in the smoothed surface brightness distribution associated with the target galaxy were identified, and regions ‘grown’ around those maxima by successively lowering a threshold isophote level from that maximum level. Pixels above the threshold were either identified with the corresponding maximum of any identified pixels that they touched, or otherwise were defined to be a new, secondary, maximum. Regions with fewer than 8 pixels were amalgamated into any touching neighbours. If multiple regions remained after this process, within the limiting 2σ isophote, the galaxy was flagged as ‘complex’.

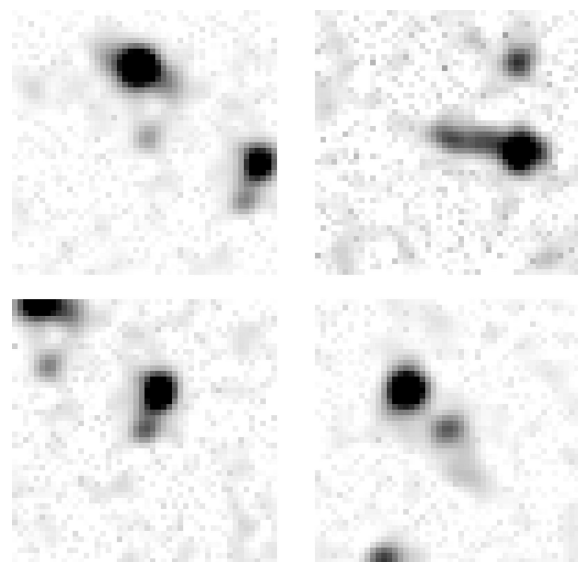


Figure 3. Examples of four galaxies excluded from measurement by the criteria described in Section 3.7, in field W1m0m1. Each panel shows a co-added image 48 pixels (approximately 9 arcsec) square, centred on each target galaxy, and the inverted grey-scale is linear up to some maximum value which varies between images.

The fraction of galaxies that were excluded in this way varied somewhat between fields, as the criteria were affected by the size of the PSF. Typically, 20 per cent of galaxies were excluded. Although this fraction seems high, such a loss of galaxy numbers does not significantly degrade the signal-to-noise ratio of the final cosmological analysis, but it does help ensure that galaxies whose measurements would be poor because of their size, or because they would be poorly modelled, have been excluded. These exclusion criteria are likely to introduce small-scale selection effects into the galaxy distribution (e.g. neighbouring galaxies would have been classed as being blended with greater or lesser probability depending on how they were aligned with respect to the PSF) and so lensing signals on arcsec scales, $\lesssim 5$ arcsec, should be excluded from analyses of this survey, even though nominal measurements are reported in the output catalogues. We note that the exclusion of some fraction of close pairs of galaxies may introduce a bias at a level of a few per cent into cosmological parameters (Hartlap et al. 2011): we do not currently have any way to estimate the size of this bias in an actual survey such as CFHTLenS, without a detailed model of the true distribution of galaxy pairs and of the effect of the measurement process on those pairs.

4 OPTIMAL COMBINATION OF MULTIPLE IMAGES

The algorithm presented in Papers I and II, and also the simulations of the GREAT08 (Bridle et al. 2010) and GREAT10 (Kitching et al. 2012) challenges, assume that each galaxy is measured on a single image. However, actual galaxy surveys use combinations of multiple exposures in the same waveband, or even across different filters. The reasons for having multiple exposures in the same filter are: (i) to increase the dynamic range of the observations; (ii) to prevent an excessive build-up of cosmic ray artefacts on any one image; (iii) to allow dithering of observations, filling in gaps where CCD boundaries or CCD artefacts prevent useful data being obtained and mitigating the effects of the finite pixel sampling. Thus

any shear measurement method should make optimal use of such multiple images. In CFHTLenS typically seven dithered exposures were obtained in each field (Section 2).

It is common practice for multiple images to be averaged together to make a single image on which shear measurement is performed. This is problematic for high-precision galaxy shape measurement, for several reasons.

(i) In order to make the averaged image, the individual images need to be first interpolated on to a common pixel grid. The interpolation causes distortion of the PSF, but because of the slightly differing local plate scale between exposures, the interpolation kernel varies across the image on scales typically less than arcmin, leading to patterns analogous to Moiré fringes and introducing artificial PSF variations on arcmin scales. These are difficult to calibrate from star measurements owing to the finite density of stars in typical observations. Knowing the interpolating kernel, the effect of the interpolation could in principle be corrected for at the location of each galaxy, but this is not a straightforward process.

(ii) The interpolation introduces correlation between pixels into the noise, which is also problematic to correctly allow for in shear measurement methods. That noise correlation also varies on arcmin scales.

(iii) Usually the PSF differs between each observation, so that averaging does not represent an optimal combination of images.

(iv) The co-addition of PSFs of differing shapes and orientations may lead to a complex stacked PSF which might be more problematic to measure and model, depending on the method adopted.

A likelihood-based method provides a straightforward way of overcoming all four issues. The galaxies on each individual image can be fitted by models that have been convolved with the correct PSF for that image, and the likelihoods simply multiplied to obtain the final combined likelihood. Images of poor quality, perhaps with a worse PSF, have a broader distribution of likelihood values than those of better quality, and hence the combined likelihood represents an optimal combination of information. Because models are fitted to each individual image, there is no need for interpolation, avoiding the first two problems listed above. Tests of the *lensfit* algorithm on stacked images, not taking into account the above effects, produced significantly worse results, with large cross-correlation signals between galaxies and the PSF, than were obtained from the joint analysis of individual exposures presented here.

In this method, the likelihoods are calculated on each image as a function of all the model parameters, including the unknown position of a galaxy. Optimum results are obtained by assuming that the galaxy's position should be the same, in celestial coordinates, on each image – provided the relative registration of the images is known, the positions on each exposure should be the same and should not be treated as independent model parameters. This approach requires highly accurate knowledge of the relative astrometric registration of the images; however, otherwise an artificial shear would be introduced which would be highly correlated between neighbouring galaxies. A similar problem arises when averaging images into a single stacked image: astrometric errors degrade the PSF. However, in that case, at least the PSF then is measured from the stacked image, so the degraded PSF is, in principle, still correctly measured (unless astrometric errors vary on small length-scales across the image). In the case of the likelihood-based method, we do not have this luxury, because the PSFs are measured on individual frames. Thus, it is essential to have an accurate measurement of the relative image registrations. Those registration shifts can then be applied to each of the models fitted to the individual images. This

process must be accurate to significantly less than one pixel, since astrometric errors are correlated over large scales across an image, and if uncorrected, would introduce an apparent shear.

The way these relative astrometric registrations were measured for CFHTLenS was, for each image, to cross-correlate the local PSF model with each star, thus measuring, from the maximum of the cross-correlation, the apparent location of each star with respect to its nominal celestial coordinate with maximum precision. A low-order polynomial function was then fitted to those registration shifts across each CCD, and the shift function applied to the positions of the PSF-convolved galaxy models when fitting each exposure of each galaxy. The shifts applied were always sub-pixel, and usually were less than one-tenth of a pixel. A further sub-pixel correction was applied to allow for the fact that an individual galaxy's postage stamp was extracted from each exposure based on the nominal celestial coordinate (using the same coordinate transformation as for the stars used to create the PSF model), but was centred at an integer pixel value (i.e. no interpolation of data was done when extracting each postage stamp).

Similarly to the issue of celestial position of a galaxy, optimum joint analysis of multiple exposures should also assume that a galaxy has an invariant, but unknown, true flux, and that the relative measured flux on each exposure is related to the true flux through a known calibration relation. As the photometric calibration was carried out over many bright stars, we assumed the uncertainty in the photometric calibration was negligible compared with the photometric uncertainty associated with each image of an individual galaxy. Thus a galaxy's flux was treated as a single unknown parameter to be marginalized over, using exposures that had the photometric calibration applied to them.

In principle, the same methodology could include exposures obtained in other filters, provided the filters are sufficiently close in wavelength that the same model provides a good description of each galaxy's structure at all the observed wavelengths. In that case, the galaxy's flux and bulge fraction values in each filter should be added as additional free parameters to be marginalized over. In CFHTLenS, only the deep, good-seeing *i'*-band exposures were used.

5 POINT SPREAD FUNCTION ESTIMATION

As in other lensing methods, we measured the PSF from images of stars on the same exposures as the galaxies being measured. The PSF was represented as a set of pixels at the same resolution as the data, and in the model-generation process, galaxy models were convolved with that PSF model. The PSF varied over the field of view: in MegaCam it varies both over the full optical field and also between the individual CCDs that comprise the full camera, and discontinuities in PSF are detectable between CCDs.

To generate a pixel representation of the PSF, an initial catalogue of candidate stars was selected from a stacked image of all exposures. Stars were selected using a conventional cut in the apparent size–magnitude plane, identifying by eye the locus of stars in the size–magnitude plane and selecting a box around that locus. The selection was further refined by applying a three-waveband *gri* selection. First, a stellar locus was defined in colour space by measuring the colour-space density of the input catalogue and rejecting any object which lies in a low-density region. Images of a subset of the resulting catalogues were inspected by eye and were found to comprise clean samples of stars (note that this procedure was relevant only for the generation of bright star catalogues for PSF creation and was not used for the classification of faint objects in

the lensing measurement). More details of the generation of the PSF star catalogues are given by Heymans et al. (2012).

A pixelized PSF model was then created for each exposure by first centring each star according to a measurement of its centroid, then fitting each centred pixel value as a two-dimensional polynomial function of position in the telescope field of view (the centring process is discussed below). For the interpolation in CFHTLenS, a third-order polynomial was found to produce good results on a random subset of the survey, with no improvement found using fourth-order, where the test used was to measure the residual cross-correlation between the PSF model and galaxy ellipticities (Section 7). To allow for the discontinuities in the PSF across the boundaries between CCDs, some coefficients of the polynomial could be allowed to vary between CCDs, with the remaining coefficients being fixed across the field of view. Good results were found allowing the constant and linear terms to vary in this way.

The polynomial fits were constrained by the images of the stars in the input catalogue, where only stars with integrated signal-to-noise ratio larger than 20 were used, and where each star was weighted by the function $w_s = v_{\text{SN},\text{star}}^2 / (v_{\text{SN},\text{star}}^2 + 50^2)$, where $v_{\text{SN},\text{star}}$ was the star's integrated signal-to-noise ratio on that exposure and where the extra term in the denominator prevents any one very bright star from dominating the polynomial fit. Any star whose central pixel was more than 50 per cent of the full-well limit for the CCD was not used, and any star which had either pixels that were flagged as not usable or pixels that appeared to be part of another object within a radius of eight pixels of the star centroid were also excluded. Pixels were deemed not usable either if they appeared to have been affected by a cosmic ray impact, if they were identified as being defective in a 'bad pixel' mask, or if they appeared to be affected by charge transfer from saturated stars, in the THELI pipeline (Section 2). Any individual CCD images that had fewer than 40 usable stars were not included in the analysis, as it was found that smaller numbers of stars could produce PSF models that were not robust to a test of varying the stars selected.

Star centroids were measured iteratively by cross-correlating with the PSF model and finding the centroid location. Star images were then interpolated so as to be centred on the pixel grid, using sinc-function interpolation, and thus assuming the PSF to be a Fourier band-limited function (i.e. assuming there are no modes above the Nyquist limit of the data). As this assumption is unlikely to be valid, this is one of the main areas to be addressed in future improvements of the algorithm (see Section 9).

The top panels of Fig. 4 show an example of the PSF pixel model created in one of the CFHTLenS exposures, sampled at three locations across the MegaCam field (two of the corners and near the centre). For comparison, the centre panels of the figure also show the mean difference between the model and the stars, averaged over all the stars on each CCD (80, 71 and 105 stars in each of the three CCDs shown). The lower panels show the rms variation in the model, measured by creating 64 bootstrap resamplings of the input star catalogue, recalculating the entire PSF model for each resampling, and measuring the rms variation in those mean residuals between the bootstrap-generated models and the stars from the original catalogue. The PSF model was fitted globally to all 36 MegaCam CCDs simultaneously, and the values shown have been evaluated by averaging over all stars on one CCD. The images have been normalized so that the sum of values in the PSF is unity. Within the central region of the PSF the typical rms is $\sim 3 \times 10^{-4}$. This test shows the reproducibility of the PSF model generation on the scale of each CCD, 6.3×14.3 arcmin².

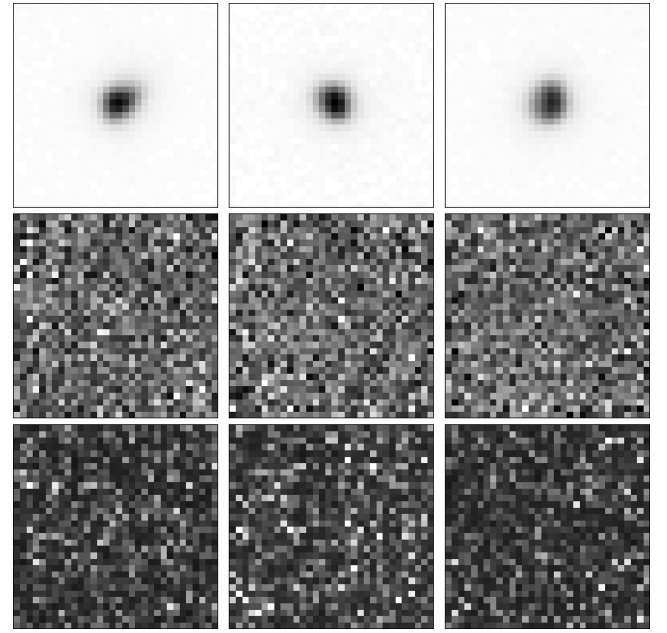


Figure 4. Example PSF models derived in a typical CFHTLenS exposure: ID 720093 in field W1p3m0. The upper panels show the model PSF sampled at three locations across the MegaCam field, with a linear grey-scale varying from 0 to 0.05, where the values are relative to the total flux in the PSF. Each panel is 32 pixels (approximately 6 arcsec) square. The centre panels show the mean residuals, averaged over all the stars selected on that CCD, between the PSF model and the star images, with a linear grey-scale varying over the range $\pm 5 \times 10^{-4}$. No residuals are larger than 7×10^{-4} anywhere in these panels. The lower panels show the rms variation in the model, measured as described in the text, with a linear grey-scale over the range 0 (black) to 5×10^{-4} (white), with maximum rms 5.6×10^{-4} . The PSFs and the residuals are shown: (left) at the centre of the CCD in the highest right ascension, lowest declination part of the MegaCam field; (centre) at the centre of the CCD immediately below the centre of the MegaCam field; (right) at the centre of the CCD in the lowest right ascension, highest declination part of the MegaCam field.

Fig. 5 shows an example of the PSF variation across one of the CFHTLenS fields. The quantities plotted are the PSF ellipticity (orientation and amplitude) and the fraction of light in the central pixel. These quantities have been averaged over seven exposures. As well as a large-scale radial distortion, jumps in PSF are visible at CCD boundaries, especially the horizontal boundaries where there are two larger gaps between detectors, and where those boundaries appear ill-defined on the figure because of dithering of the seven exposures.

6 DISTORTION CORRECTION

Wide-field cameras introduce distortion into the field-of-view which can mimic the cosmological shear signal, and this therefore needs to be corrected for. Because field distortion may vary with telescope position and temperature, and because exposures are dithered, it should ideally be mapped and corrected on each individual exposure. The distortion information is implicitly available in the astrometric calibration of the survey: provided the astrometric solution has good absolute accuracy, the distortion can be measured simply by measuring the relationship between pixel and celestial coordinates as a function of position across the field. The centroid positions of stars, on which the astrometric solution is based, can generally be measured to subpixel accuracy, whereas the scalelengths of

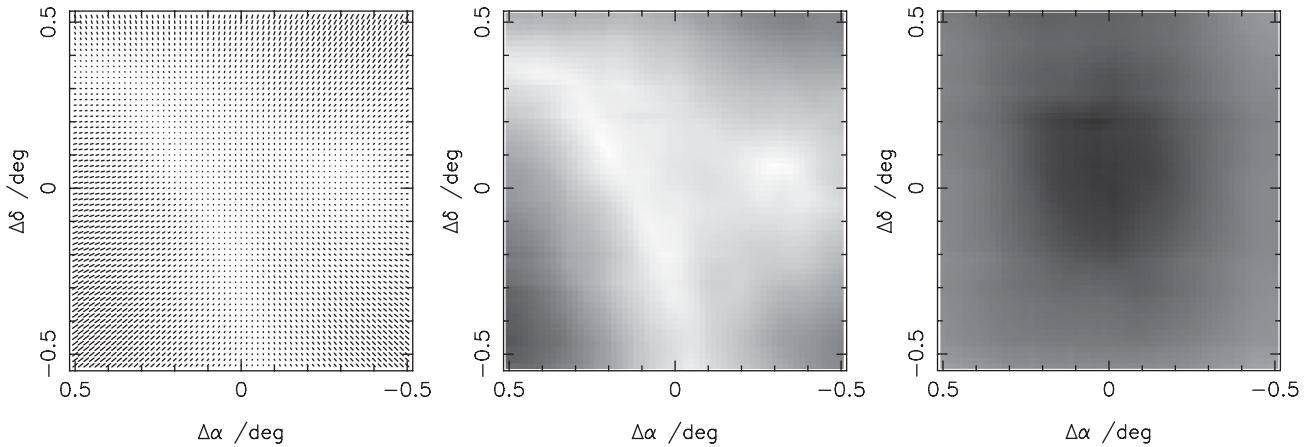


Figure 5. The variation of measured PSF over one of the CFHTLenS fields with a highly variable PSF, W1m1p2. The plot has been made by combining information from seven exposures, which are somewhat offset with respect to each other. PSF ellipticities have been simply averaged. Left: stick plot showing the variation of measured PSF ellipticity across the field. The maximum stick length on the plot corresponds to ellipticity of 0.13. Centre: grey-scale showing the ellipticity amplitude, with black having an ellipticity of 0.16. Right: grey-scale showing the fraction of the light in the central pixel of the PSF, with black indicating a fraction 0.1.

distortion variations are typically arcmin or greater, so that the typical accuracy in distortion correction by this method is shear error of the order of 10^{-3} or better.

Having measured the distortion on an image, this information was incorporated into the galaxy models by calculating the local linearized distortion transformation and applying that as an affine transformation to the models being fitted to each galaxy on each individual exposure. In other words, the effect of distortion was corrected by including it in the forward modelling of each galaxy on each exposure, rather than by trying to remove it from the data. Because the PSF models had not had distortion removed from them (the stars were not distortion-corrected on input to the models), the distortion was applied to the galaxy models prior to PSF convolution to ensure that the distortion and convolution corrections to the models correctly mirrored these effects in the data. In practice, the models were fitted in Fourier space, so the affine transformation was also applied in Fourier space, using a simple matrix operation to convert the affine transformation between real and Fourier space (Bracewell et al. 1993). The affine transformation allowed for scale, rotation and shear distortion, and ensured that the ellipticities of all galaxies were measured with respect to the celestial coordinate system, and not any arbitrary pixel grid (i.e. galaxy position angle was defined with respect to the local north–south, east–west axes at each galaxy). When calculating shear correlation functions or power spectra from these data on large scales, it may be necessary to correct for the translation of the celestial coordinate system around the celestial sphere (i.e. there needs to be a ‘course angle’ correction to the differences in position angles of pairs of galaxies separated on large angular scales, when not working on the celestial equator; see also Kilbinger et al. 2012).

7 TESTS FOR SYSTEMATIC ERRORS

We postpone detailed analyses for systematic biases to a companion paper (Heymans et al. 2012); however in this section we discuss the chief effects which may bias the shear measurement results and illustrate the tests that we have carried out, showing specific examples from individual CFHTLenS fields.

There are a number of possible effects which could lead to the creation of systematic biases in the measured shear.

- (i) Inaccurate PSF modelling.
- (ii) Use of galaxy models and prior distributions for the model parameters that are not representative of the true population.
- (iii) Bias in the shear estimator (such as that described in Appendix C).
- (iv) Inaccurate relative astrometry.

PSF modelling errors could be investigated by comparing the models with the stars used to generate the models (i.e. a quantitative assessment such as shown in Fig. 4): however, it is difficult to assess how large a deviation between model and reality can be tolerated, what is really needed is an assessment on the effect of the final measured galaxy shear. Such an analysis has been presented by Paulin-Henriksson et al. (2008) and Paulin-Henriksson, Refregier & Amara (2009), giving a good qualitative picture of the effects expected from PSF modelling errors. It is difficult to make a quantitative prediction of the effect on shear without repeating the analysis with realistic galaxy sizes and shapes, PSF surface brightness distributions that match the CFHTLenS PSFs and likely PSF modelling errors. However, the results of Paulin-Henriksson et al. (2008) show that we expect modelling errors to result in significant cross-correlation between the PSF model and the galaxy ellipticities, and furthermore that the cross-correlation is expected to be a strong function of galaxy size. The use of incorrect galaxy models or prior distributions is likewise expected to produce size-dependent PSF–galaxy cross-correlation. Bias in the shear estimator is expected to produce PSF–galaxy cross-correlation with a strong signal-to-noise ratio dependence (Fig. C2).

To test for these effects, we first measured, in each individual CFHTLenS field, the cross-correlation between galaxy ellipticity and the mean ellipticity of the PSF model, measured at the location of each galaxy, with PSF ellipticity averaged over the exposures used. The cross-correlation is shown for three typical fields, as a function of signal-to-noise ratio ν_{SN} and fitted galaxy size, in Fig. 6. Statistical uncertainties were calculated by randomly dividing the observed samples into 10 subsamples, and measuring the variance in the statistic between the subsamples. The reported error is then $\sqrt{10}$ smaller than the subsample rms. However, on scales of 1° , there is a significant additional term that arises from the random cross-correlation of the cosmic shear signal with the PSF. In Heymans

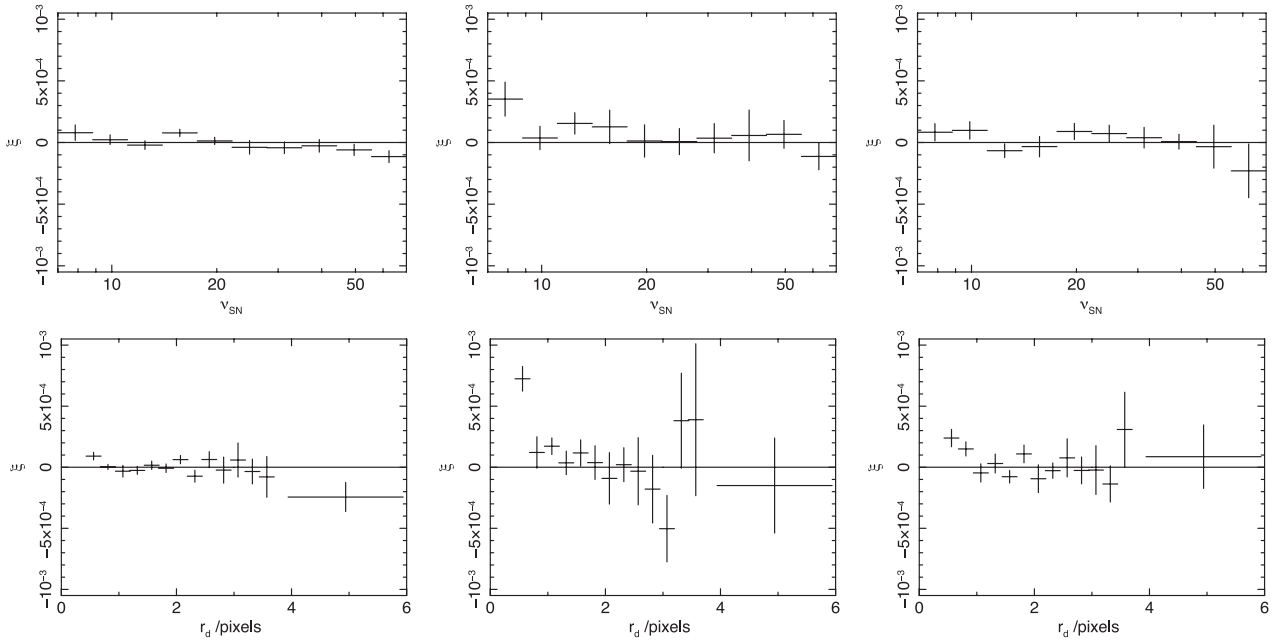


Figure 6. Examples of cross-correlation between PSF model ellipticity and galaxy ellipticity in three CFHTLenS fields, from left to right, W1m2p1, W1p3m0 and W1p2m3. Upper panels show the cross-correlation as a function of signal-to-noise ratio v_{SN} , lower panels as a function of fitted galaxy major-axis scalelength r_d . Vertical bars indicate 68 per cent confidence intervals.

et al. (2012) this term is explicitly estimated using shear derived from the dark matter simulations of Harnois-Deraps, Vafaei & van Waerbeke (2012). In this paper, this additional term is ignored. It manifests itself as an uncertainty, which is highly correlated between the plotted points, to be added in quadrature to the statistical uncertainties shown in Fig. 6.

The test presented may not be sensitive to all conceivable forms of PSF modelling error. For example, the PSF model may be too small in some regions, leading to positive PSF–galaxy correlation, but too large in others, leading to negative PSF–galaxy correlation. Averaging over the field may lead to no net detection of an effect. However, we are not able to analyse the fields in smaller subsamples, as the statistical and cosmic shear uncertainties become large.

The results from the full-field analyses are generally encouraging, with no obvious dependence of a cross-correlation signal on either signal-to-noise ratio or fitted galaxy size. The possible exception to this statement is at the lowest fitted galaxy scalelengths, less than 0.7 pixels, where often a positive cross-correlation is observed. This is very likely due to reaching the limits of the sampled models that were generated as described in Appendix A. It should not be surprising that systematic effects may be seen for galaxies with major axis scalelengths smaller than one pixel in size, and considerably less than the PSF size. Fortunately, the number of such galaxies is relatively small, and they all have small weights (Fig. 2).

If there are residual errors in the relative astrometry, this would result in systematic shear patterns being induced with scalelengths of variation determined by the scalelength of the astrometric errors (probably many arcminutes). The resulting systematic signal would likely not be correlated with any other measurable quantity. We tested for this in a subset of CFHTLenS fields by switching off the relative astrometry correction described in Section 4, remeasuring those fields and generating shear-shear correlation functions. Only small differences, much smaller than the statistical uncertainties in the correlation function, were observed, indicating that for CFHTLenS the absolute astrometric ac-

curacy already is adequate for shear measurement from individual exposures.

An in-depth analysis of the survey for systematic errors is presented by Heymans et al. (2012).

8 CFHTLENS IMAGE SIMULATIONS

The *lensfit* shape measurement algorithm has previously undergone a series of verification tests using simulated sheared images that test a range of different observing conditions, galaxy and PSF types (Paper II). These tests have used both sets of simulations from the Shear TEsting Programme (STEP; Heymans et al. 2006; Massey et al. 2007), demonstrating sub-per-cent level accuracy. The philosophy of *lensfit*, to analyse data from individual exposures rather than a co-added image, requires a suite of specialized image simulations that mimic this main feature of the CFHTLenS analysis. We describe in this section how these simulations compare to previous image analysis challenges and the subsequent verification and calibration analysis.

As well as confirming that the algorithms described above do work on multiple exposures, the simulations also allow us to measure the calibration term that is required to correct for noise bias, as discussed in Section 3.5, in Appendix C and by Refregier et al. (2012) and Melchior & Viola (2012). Kacprzak et al. (2012) also describe the use of simulations for calibration of noise bias.

8.1 GREAT CFHTLenS image simulations

Our philosophy followed the strategy of the GREAT challenges⁷ (Bridle et al. 2010; Kitching et al. 2012) to design the simplest simulation that nonetheless addresses the specific question: how accurate is *lensfit* when analysing realistic galaxy distributions,

⁷ www.greatchallenges.info

measured with multiple, low-signal-to-noise ratio exposures? We therefore chose a constant PSF across the field of view, and we postpone discussion of our tests on the robustness of the measurements to a spatially varying PSF in the actual data to the companion paper (Heymans et al. 2012). In the simulations, we did, however, vary the PSF between the multiple exposures simulated for each field, based on the variation between exposures that we see in the CFHTLenS data.

We simulated the full CFHTLenS area with each simulated 1 deg^2 field consisting of seven dithered exposures, with a linear dithering pattern that was matched to the average linear dither pattern in the survey, ignoring higher order astrometric corrections. Galaxies and stars were positioned within the field with the same distribution as in the real data and with matching magnitude and signal-to-noise ratio. We then created a population of disc-plus-bulge galaxies such that the distributions of galaxy brightness, size, intrinsic ellipticity and bulge fraction matched the distributions described in Section 3.2 and Appendix B. As in those models, the disc and bulge ellipticities were set equal. The unlensed orientation of each galaxy was random and each galaxy was assigned a cosmological shear based on its original position and redshift in the CFHTLenS data, as simulated by our N -body lensing simulation of CFHTLenS (Harnois-Deraps et al. 2012). In order to potentially allow reduction of the impact of noise from the intrinsic galaxy ellipticity, we followed Massey et al. (2007) by creating two image simulations for each field, with the intrinsic ellipticities rotated by 90° between the two images, although in practice we found that such shape noise mitigation was itself problematic and was therefore not used, as discussed in Section 8.4. Camera distortions were not included in the simulations but the *lensfit* pipeline still applied a correction as it would on real data.

The model prior distributions differed significantly from the distributions assumed in previous image challenges, as shown in Fig. 7. We simulated the galaxy population as comprising two types: 90 per cent of galaxies were disc-dominated, and the remainder were bulge-dominated. No irregular galaxies were included. The upper panel compares the intrinsic ellipticity distributions in our custom CFHTLenS simulations to those simulated in the first STEP challenge and both GREAT challenges. It can be seen that the distributions of ellipticity for the disc-dominated population differ significantly. The bulge-dominated galaxy population ellipticity distribution was more similar to the ellipticity distribution assumed in previous challenges, although the fraction of bulge-dominated galaxies in the STEP simulations was negligible. The lower panel of Fig. 7 shows the size distribution for both populations, which varied with galaxy magnitude in our custom simulations. This can be compared to the somewhat larger size distribution for the first STEP challenge, scaled according to the relative pixel sizes in STEP and CFHTLenS, and the set of three discrete galaxy-to-PSF size ratios simulated in the GREAT08 challenge, which for the average CFHTLenS PSF corresponds to galaxy scalelengths of $r = [0.2, 0.3, 0.6] \text{ arcsec}$. Comparing these distributions indicates how challenging our custom simulations were compared to previous image simulations, owing to the more realistic, broad ellipticity distributions and high fraction of small galaxies. We therefore would expect neither *lensfit* nor other methods to perform as well on these custom simulations as in previous challenges.

The image simulations were produced using a modified version of the GREAT challenges image simulation software package detailed in appendix A.3 of Bridle et al. (2010). Aside from the differences in galaxy populations described above, our simulations also differed from the GREAT08 and GREAT10 simulations by allowing for

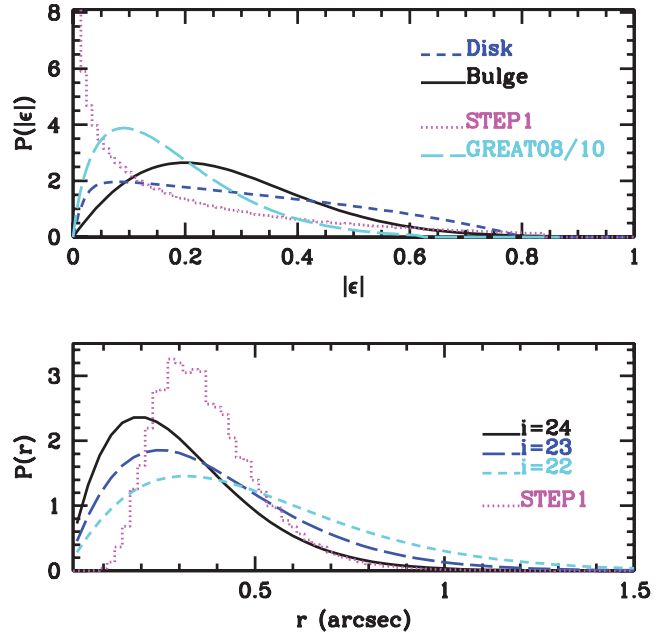


Figure 7. Upper panel: the intrinsic ellipticity distribution for disc-dominated galaxies in the STEP (dotted curve), GREAT (long-dashed) and CFHTLenS (short dashed) simulations and for bulge-dominated galaxies in the CFHTLenS simulations (solid curve). Lower panel: the magnitude-dependent galaxy size distribution, where r is the major-axis half-light radius for the bulge component and the major-axis exponential scalelength for the disc component. The CFHTLenS simulated distributions are shown at magnitudes $i' = 22$ (short dashed curve), $i' = 23$ (long dashed) and $i' = 24$ (solid). The distribution of the first STEP challenge is also shown (dotted curve).

overlapping object isophotes, by having each exposure containing both galaxy and stellar objects and by simulating dithered data. We also modified the definition of galaxy size to be the size of the semi-major axis of a galaxy's isophotes, instead of the geometric mean of the semi-major and semi-minor axes. This redefinition is consistent with the convention used for measurements of data (e.g. Simard et al. 2002) and avoids introducing an incorrect correlation between size and ellipticity (see also Appendix A). The definition of signal-to-noise ratio was chosen such that it was defined within a circular aperture of radius of 0.935 arcsec (5 pixels) for all galaxies. Using a matched aperture flux, as measured by *SEXTRACTOR* on the data, then allowed us to directly replicate the signal-to-noise ratio distributions within each CFHTLenS exposure.

The choice of simulated PSF model was motivated by the data. To each of the 36 MegaCam CCDs in each CFHTLenS exposure, we fitted an elliptical-isophote Moffat (1969) profile to the images of stellar objects. A circular light profile at distance r from the centre of a star was given by

$$I(r) = I_0 \left[1 + \left(\frac{r}{r_s} \right)^2 \right]^{-\beta_{\text{Moffat}}} \quad (9)$$

where I_0 is the peak flux, and the scale radius r_s and the shape parameter β_{Moffat} were related to the FWHM of the PSF as

$$\text{FWHM} = 2r_s \sqrt{2^{1/\beta_{\text{Moffat}}} - 1}. \quad (10)$$

We created a series of circular model PSFs on a pixel grid for a range of scale radii r_s and shape parameters β_{Moffat} and applied a linear transformation of the co-ordinates to shear the model PSF to have elliptical isophotes. We then calculated the best-fitting PSF

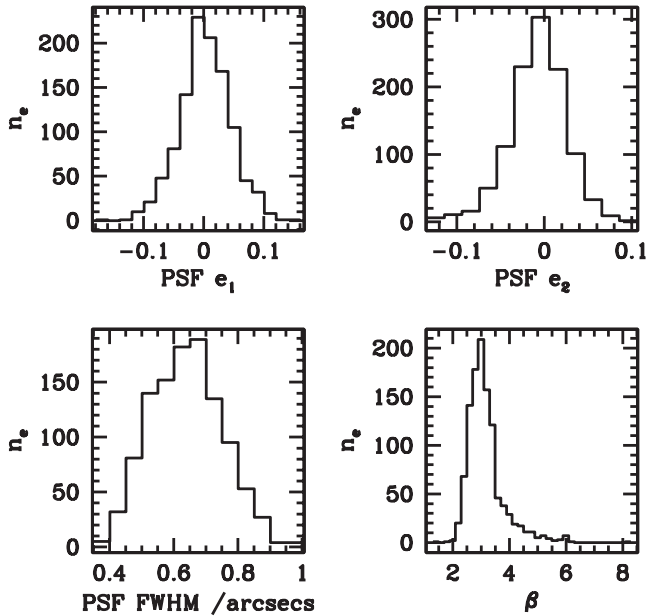


Figure 8. The distribution of PSF parameters measured from the CFHTLenS data and replicated in the CFHTLenS simulations. The upper panels show the PSF ellipticity. The lower panels show the PSF FWHM in arcsec (left) and the shape parameter β_{Moffat} (right).

ellipticity and the Moffat profile parameters β_{Moffat} and α for each CCD chip. For each simulated exposure we adopted a PSF model for the full field of view by selecting a CCD PSF at random. This was to ensure that we fully represented the range of PSF models at all locations in the images, rather than taking the often more circular average over the full field of view. We then grouped the PSF model exposures into fields following the grouping of exposures in the real data. The distribution of PSF parameters measured from the CFHTLenS data, and replicated in the simulations, is shown in Fig. 8.

8.2 SkyMaker CFHTLenS image simulations

When constructing calibration corrections, it is of crucial importance to verify the base image simulations used to ensure the calibration is not dependent on the simulation software used. To this end we created a second suite of 140 MegaCam image simulations generated using SkyMaker⁸ (Bertin 2009). The SkyMaker PSFs were not Moffat profiles, but instead had radial surface brightness profiles that took into account the telescope optics, atmospheric seeing and scattering into wings of the PSF, although those adopted here had circular symmetry. Unlike the GREAT simulations, only a single exposure was simulated, for reasons of computational time, and the signal-to-noise ratio distribution of galaxies on those single exposures was matched to that of the CFHTLenS galaxies on the co-added data. The signal-to-noise ratios were generally smaller than in the SkyMaker simulations tested in the first STEP challenge (Heymans et al. 2006). Furthermore, the size and ellipticity distributions were chosen to match those in the custom CFHTLenS simulations described above, unlike the SkyMaker simulations tested in Heymans et al. (2006), which follow the dotted curves in Fig. 7.

The galaxy and star positions were chosen at random and circular PSFs with an FWHM of 0.7 arcsec were used. A cosmic shear signal from a redshift plane at $z = 1$ was assumed in the simulation.

8.3 Model bias

The primary aim of the simulations was to measure the bias caused by noise (Appendix C), and for this aim the simulated galaxies should match the real population as closely as possible. The strategy of matching the simulations to the models and priors used in the *lensfit* measurements is reasonable, as these distributions have been derived from observations and are therefore a good representation of the galaxy population. The tests presented here do not test for the effect of deviations of the true population from the model assumptions, but they do test the efficacy of the method taking into account the implementation of the statistical algorithm in Section 3.3, including the effects of convolution with appropriate PSFs, the effects of finite signal-to-noise ratio and noise bias, the combination of multiple dithered images and the effect of masks applied to the data.

The difficulty of attempting also to test for the effect of deviations of the true population from the assumed population is that we have little prospect of obtaining an independent assessment of how different the real Universe might be from our assumptions, even at the level of knowing the distributions of surface brightness of individual galaxies in the local Universe (Graham & Worley 2008). Notwithstanding that difficulty, this question has been considered by other authors. Voigt & Bridle (2010) considered the multiplicative biases that might arise from fitting multiple Gaussian components to exponential or de Vaucouleurs surface brightness distributions, convolved with a PSF. It was found that fitting a single Gaussian to an exponential distribution caused a multiplicative bias of 3 per cent and an additive bias of ~ 0.001 , but that these biases fell to 0.3 per cent and $< 10^{-4}$ when two Gaussian components were adopted. Larger biases were found when fitting Gaussians to a pure de Vaucouleurs profile (3–6 per cent and 0.001, respectively, when fitting two Gaussians), but this is not surprising given the significant difference in surface brightness profile between Gaussian and de Vaucouleurs distributions. We note that pure de Vaucouleurs-profile galaxies are expected to be a small fraction of the true population. Multiplicative biases of 2–5 per cent were found when fitting two Gaussians to composite galaxy profiles comprising exponential plus de Vaucouleurs components: in this case we expect our fitting of exponential plus de Vaucouleurs models to work well and therefore to lead to significantly smaller biases.

Voigt & Bridle (2010) argue that these biases arise because a Gaussian profile is too flat to reproduce well the central regions of realistic galaxy profiles. In this paper, we fit two components, exponential plus de Vaucouleurs, to the real population of galaxies, and these should be far better representations of the true surface brightness distributions than multiple Gaussians, with sufficient flexibility to model well a wide range of Sersic indices. We therefore expect model bias to be significantly smaller than found by Voigt & Bridle (2010) when fitting Gaussians, and to be subdominant compared with noise bias, as evaluated in the following sections, in the CFHTLenS analysis.

Future surveys covering larger sky areas will have more stringent requirements on the acceptable level of systematics than CFHTLenS, and it is likely that the question of model bias will need to be revisited for those surveys.

⁸ www.astromatic.net

8.4 Analysis and calibration

We processed both sets of image simulations with *lensfit* using the known galaxy and star positions as an input catalogue. Fig. 9 shows the measured galaxy size, signal-to-noise ratio, bulge fraction and ellipticity distributions from both sets of image simulations and the CFHTLenS data. In all cases the distributions incorporate the weights assigned by *lensfit* (equation 8). The simulated signal-to-noise ratio distribution shown in the upper left-hand panel was constructed to match the CFHTLenS depth. We found reasonable agreement between the ellipticity distribution measured in the data and simulations, although in the weighted distributions of CFHTLenS there is an excess of galaxies at $|e| \simeq 0.45$. We do not expect perfect agreement between the measured distributions and priors, because only the summed posterior distribution is expected to match the prior distribution (Paper I), whereas Fig. 9 shows the values obtained from the expectation values of the probability distribution (after marginalizing over the other parameters). The excess is only present in the weighted distribution, and not in the unweighted distribution, which implies that, on average, galaxies with non-zero, middling ellipticities have more sharply defined likelihood surfaces than their peers at lower or higher ellipticity. This effect is not fully understood, but it is probably linked to the issue of noise bias that is discussed in Appendix C. In comparison to the data, in the simulations there are also lower fractions of small galaxies for which an ellipticity was determined. As these small galaxies do exist in the simulation (see Fig. 7) this might indicate a mismatch between the PSF in the data and simulations, such that small galaxies were harder to measure in the simulations than in the real data. Either pixelization and noise from the data led us to overestimate the PSF size which was then used in the simulation, or the Moffat profile used in the GREAT simulations and the SkyMaker PSF were poor models of the MegaCam PSF. There are, however, sufficient numbers of small galaxies to calculate a calibration as described below, and as these small galaxies were assigned the lowest weights by application of equation (8), it should not be problematic that their

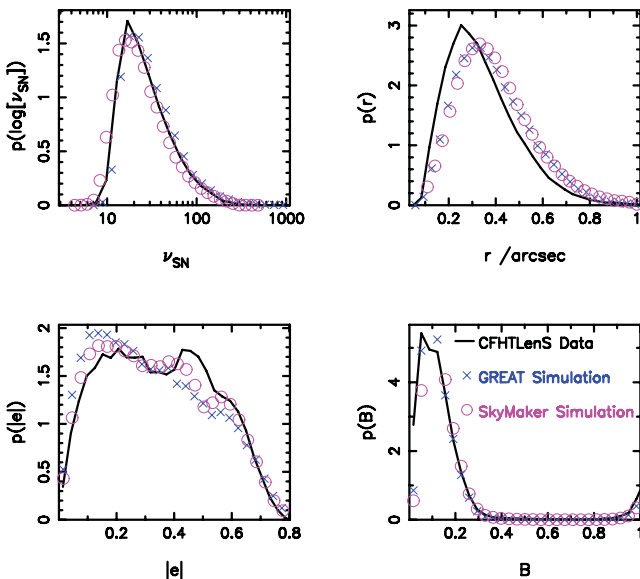


Figure 9. Comparison of the weighted distributions of measured signal-to-noise ratio (upper left), galaxy size (upper right), ellipticity distributions (lower left) and bulge fraction (lower right) from the custom GREAT image simulations (crosses), custom SkyMaker image simulations (circles) and the CFHTLenS data (solid).

calibration was based on relatively fewer galaxies than for other sizes in the data.

We follow Heymans et al. (2006) in modelling our shape measurement as a linear combination of a multiplicative error m and an additive error c such that

$$e^{\text{obs}} = (1 + m)e^{\text{true}} + c + \Delta e \quad (11)$$

where e^{obs} is the observed complex ellipticity as measured by *lensfit*, Δe is the noise on the measure and e^{true} is the sheared intrinsic ellipticity defined as in equation (1) and for simplicity we treat m as an orientation-independent scalar. At this point we are faced with several alternatives to determine m and c . Comparing e^{obs} directly with e^{true} would lead to a strong noise bias in the measurement of m and c , as the observed ellipticity is bounded with $|e^{\text{obs}}| < 1$ such that $\langle \Delta e \rangle$ is non-zero for a galaxy with non-zero ellipticity. Comparing e^{obs} directly with shear γ would also lead to an intrinsic ellipticity noise bias, as for a finite number of galaxies $\langle e^{\text{true}} \rangle$ may be non-zero. This is particularly relevant when galaxies are binned in order to look for size- and signal-to-noise-ratio-dependent calibrations, increasing the random shot noise in each bin. Massey et al. (2007) developed a method to circumvent this issue by including rotated pairs of galaxies in the simulations such that $\langle e^{\text{int}} \rangle = 0$ by construction, where e^{int} is the intrinsic ellipticity prior to application of any lensing shear. We generated such galaxy pairs in the custom GREAT simulations, but found that, when galaxies within a pair were assigned different weights, which at the low signal-to-noise ratio regime was typical, $\langle w e^{\text{int}} \rangle$ was significantly non-zero. Thus we did not use this method of shape-noise cancellation. Instead, our preferred analysis method was to first bin galaxies by size r and signal-to-noise ratio ν_{SN} , such that there were equal numbers of objects in each bin, and then to subdivide those bins by the input cosmic shear of each galaxy. Each galaxy was entered into two bins of shear, treating each component as being an independent measure of shear for the purposes of deriving the calibration correction. Then, a calibration was derived by minimizing χ^2 given by

$$\chi^2 = \sum_i \frac{1}{\sigma_{\gamma}^2} (\bar{\gamma}_i^{\text{obs}} - [1 + m(\nu_{\text{SN}}, r)] \bar{\gamma}_i^{\text{true}} - c(\nu_{\text{SN}}, r))^2, \quad (12)$$

where for each bin i

$$\bar{\gamma}_i^{\text{obs}} = \frac{\sum_j w_j (e_{\alpha j}^{\text{obs}} - e_{\alpha j}^{\text{int}})}{\sum_j w_j} \quad (13)$$

and where for galaxy j in bin i , $e_{\alpha j}$ is either the e_1 or e_2 component of the ellipticity, with $\alpha = 1, 2$ being set by the component of the input cosmic shear for that galaxy in that bin.

The error σ_{γ} on $\bar{\gamma}_i^{\text{obs}}$ was estimated through a bootstrap error analysis within each bin i . σ_{γ} was found to be uncorrelated between bins. The summation over both ellipticity components in equation (13) was motivated by an initial component-dependent minimization, which showed negligible differences in the calibration measured for each ellipticity component.

This analysis method ensured that we were unbiased to measurement or intrinsic ellipticity noise, but it does disguise biases introduced by the weights. In Fig. 10 we use the GREAT CFHTLenS simulations to show the weighted average true intrinsic ellipticity $\langle w e^{\text{int}} \rangle$ as a function of the PSF ellipticity e^* for four sets of galaxies binned by galaxy size and signal-to-noise ratio. For each galaxy sample we find a preferential weighting of galaxies oriented in the same direction as the PSF. This weight bias is independent of which ellipticity component we choose and affects bright large galaxies to a similar extent to small, faint galaxies. It is an undesirable

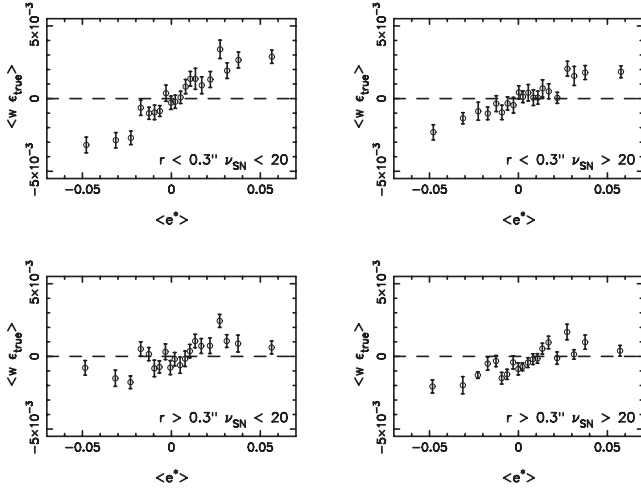


Figure 10. The weighted average true intrinsic ellipticity ($w\epsilon^{\text{int}}$) as a function of the PSF ellipticity e^* for four sets of GREAT-simulated galaxies binned by galaxy size and signal-to-noise ratio. Small galaxies with major-axis scalelengths $r < 0.3$ arcsec are shown in the upper panels; large galaxies with $r > 0.3$ arcsec are shown in the lower panels. The left-hand panels show the faint sample with $\nu_{\text{SN}} < 20$; the brighter sample with $\nu_{\text{SN}} > 20$ are shown on the right-hand panels. Each bin contains roughly the same number of galaxies.

outcome of the weighting scheme, which correctly identifies that it is easier to measure the shapes of galaxies oriented with respect to the PSF in contrast to galaxies oriented perpendicular to the PSF, and may be viewed as a form of the orientation bias discussed by Bernstein & Jarvis (2002, Section 3.6). We investigated isotropization schemes to remove the PSF-related weight bias by assigning galaxies an averaged weight based on its size relative to the PSF size and the signal-to-noise ratio of the galaxy. We found, however, that the isotropization reduced the effectiveness of the weights, which also decrease the noise bias in the shear measurement (Appendix C), and resulted in an even stronger weight bias in the opposite sense. We therefore concluded that weight bias, which impacts shape measurement at the 10^{-3} level, is an undesirable feature of applying inverse-variance weighting, but is a small effect for the study of CFHTLenS. In Heymans et al. (2012) we describe a method to identify and remove fields which show any significant correlation with the PSF, which can arise both from this weight bias and from the general effect of noise bias.

Fig. 11 shows some example results of the measured multiplicative error $m(\nu_{\text{SN}}, r)$ in the GREAT simulations as a function of ν_{SN} for four size bins selected to be the smallest, and the three discrete galaxy sizes simulated in GREAT08. The solid curve is a fit to the full data set, which comprises 20 bins in each of ν_{SN} and size, given by

$$m(\nu_{\text{SN}}, r) = \frac{\beta}{\log_{10}(\nu_{\text{SN}})} \exp(-\alpha r \nu_{\text{SN}}) \quad (14)$$

with $\alpha = 0.057$ and $\beta = -0.37$, for r measured in pixels. We find that for all galaxies with $r > 0.5$ arcsec, or $\nu_{\text{SN}} > 35$, m is consistent with zero on average, but that the calibration correction increases for smaller and lower signal-to-noise ratio galaxies, as expected for noise bias (Appendix C). The *lensfit* weighting scheme already downweights these noisy galaxies, such that even though the calibration correction can be significant, these galaxies contain very little weight in our scientific analyses. This can be seen in the left-hand panel of Fig. 12, which shows the multiplicative calibration

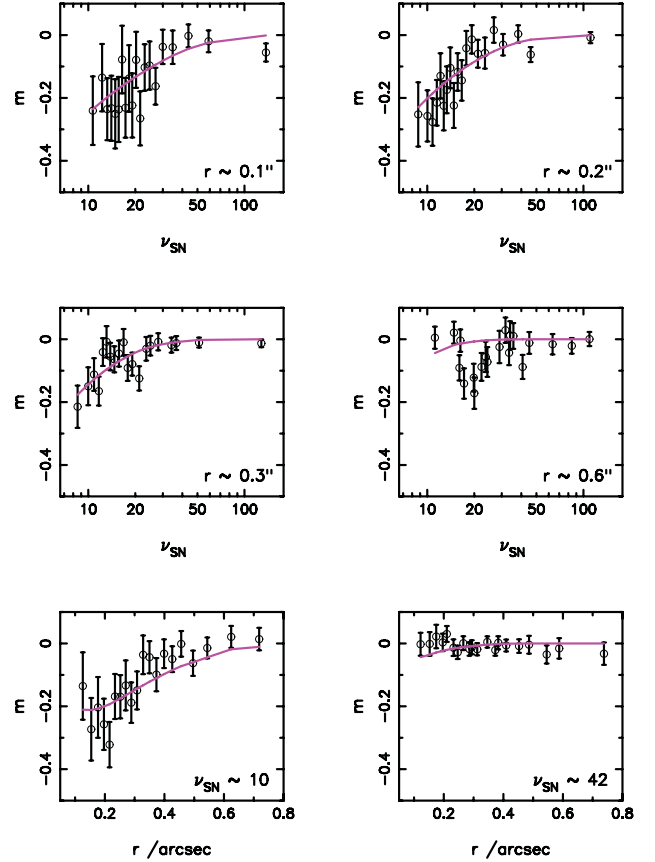


Figure 11. Example results of the measured multiplicative error $m(\nu_{\text{SN}}, r)$ in the GREAT simulations as a function of signal-to-noise ratio for four size bins selected to be the smallest scalelength (0.1 arcsec, upper left), and the three discrete galaxy sizes equivalent to those simulated in GREAT08 (0.2, 0.3, 0.6 arcsec). The solid curve is the best-fitting functional form to the full data set which comprises 20 bins in both signal-to-noise ratio and size. The lower two panels show $m(\nu_{\text{SN}}, r)$ as a function of r for two bins of ν_{SN} .

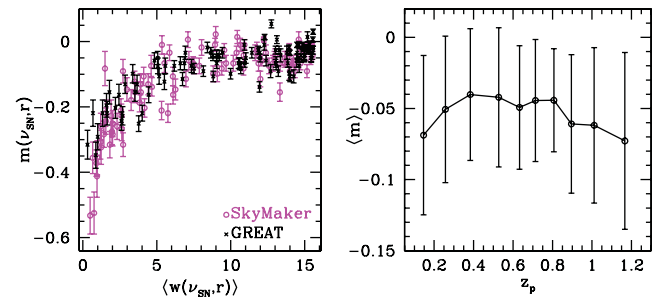


Figure 12. Left: the multiplicative calibration m for each signal-to-noise ratio and size bin as a function of the average *lensfit* weight $\langle w \rangle$ of the simulated galaxies in the bin. Results are shown for the GREAT (crosses) and the SkyMaker (open circles) simulations. Right: the mean applied multiplicative calibration $\langle m \rangle$ for CFHTLenS galaxies binned by photometric redshift z_p . Vertical bars indicate the rms of the distribution of m values within each bin.

tion for each bin of signal-to-noise ratio and size as a function of the average weight of the galaxies in the bin, for both the GREAT and SkyMaker simulations. It can be seen that the net effect is very similar in the two sets of simulations, although the SkyMaker simulations display a slightly larger bias than the GREAT simulations for galaxies with low signal-to-noise ratio and thus lower

lensfit weight. As the GREAT simulations have anisotropic PSFs that closely match the CFHTLenS PSFs, and as they incorporate the effect of having seven multiple exposures, whereas the Sky-Maker simulations are single exposures in each field, we use the results from the GREAT simulations to calibrate the CFHTLenS shear measurements. The differences between the results from the simulations only become apparent for galaxies which are significantly downweighted in CFHTLenS, and overall it is encouraging to note the close match between the two sets of simulations. The right-hand panel of Fig. 12 shows how the mean applied calibration correction $\langle m \rangle$ varies with the photometric redshift of the galaxies in CFHTLenS, showing primarily the larger correction that needs to be applied at redshifts $z_p \gtrsim 0.8$ owing to the decreasing signal-to-noise ratios and sizes of galaxies at higher redshifts. We note also the larger correction deduced for galaxies at low redshifts, $z_p \lesssim 0.3$, which probably arises from contamination of the low-redshift population by galaxies with higher spectroscopic redshifts (see Heymans et al. 2012, for further discussion).

We find no evidence for a significant additive error c term in the *lensfit* analysis of the simulations, indicating that the multiplicative error arising from noise bias is the dominant term to be corrected.

8.5 Simulated cosmic shear tomography

To test the application of such a calibration, we estimated the calibration-corrected two-point shear correlation function (other papers will describe the correction to estimators such as the three-point function and others). We first calculated the uncalibrated two-point function,

$$\xi_{\pm}(\theta) = \frac{\sum w_i w_j [\epsilon_i(\mathbf{x}_i) \epsilon_j(\mathbf{x}_j) \pm \epsilon_r(\mathbf{x}_i) \epsilon_r(\mathbf{x}_j)]}{\sum w_i w_j}, \quad (15)$$

where the weighted sum is taken over galaxy pairs with $|\mathbf{x}_i - \mathbf{x}_j| = \theta$. The two-point correction term was then estimated as

$$1 + K(\theta) = \frac{\sum w_i w_j [1 + m(v_{\text{SN},i}, r_i)][1 + m(v_{\text{SN},j}, r_j)]}{\sum w_i w_j} \quad (16)$$

for the same galaxy pairs, thus taking into account potential astrophysical correlations between galaxy size, brightness and clustering scales. The calibration corrected two-point function was then given by

$$\xi_{\pm}^{\text{cal}}(\theta) = \frac{\xi_{\pm}(\theta)}{1 + K(\theta)}. \quad (17)$$

The alternative of applying the calibration correction directly to individual galaxies could result in instability in the rare events where $1 + m(v_{\text{SN},i}, r_i) \rightarrow 0$. In addition, using this method automatically removes any unwanted correlation between the calibration correction and intrinsic ellipticity which could arise from the measured galaxy signal-to-noise ratio or size being weakly correlated with shape, which is a possible outcome of any shape measurement method. We show a comparison of the measured and input two-point correlation functions in Fig. 13, for the raw uncalibrated data and the calibrated data. The error bars are representative of the signal-to-noise ratio expected for a CFHTLenS-like survey and the data points are strongly correlated. To investigate the impact of the calibration as a function of redshift we show the signal for three tomographic configurations; a full sample with $0.1 < z < 1.3$ similar to the sample used in Kilbinger et al. (2012) and two tomographic bins with $0.5 < z < 0.85$ and $0.85 < z < 1.3$, similar to the two tomographic bins used in Benjamin et al. (2013). Note that a very low

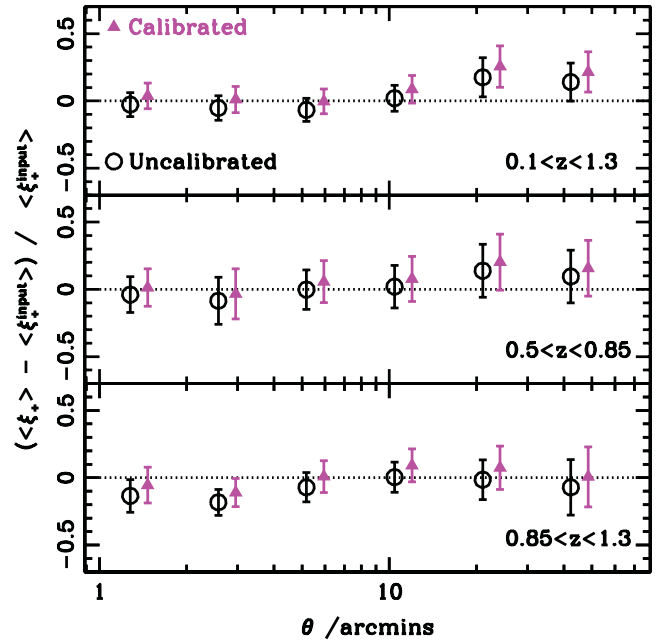


Figure 13. Comparison of the measured to input two-point correlation function for the raw uncalibrated data (circles) and the calibrated data (triangles) from the GREAT CFHTLenS simulations. The error bars are representative of the signal-to-noise ratio expected for a CFHTLenS-like survey. To investigate the impact of the calibration as a function of redshift we show the signal for three tomographic configurations: $0.1 < z < 1.3$ (upper panel), $0.5 < z < 0.85$ (middle panel) and $0.85 < z < 1.3$ (lower panel).

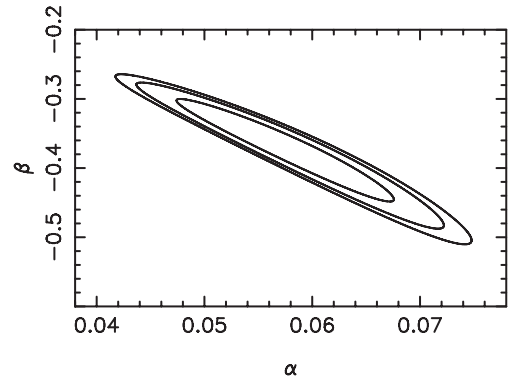


Figure 14. Confidence regions of the multiplicative calibration correction parameters in equation (14) used to infer systematic error contribution to the calibration corrected two-point function $\xi_{\pm}^{\text{cal}}(\theta)$.

redshift bin with $0.1 < z < 0.5$ has statistical error bars that encompass the full figure. Fig. 13 shows that for a CFHTLenS-like survey the correction is well within the statistical errors. We do see that the calibration correction is more significant for the higher redshift sample, as expected given that the average size and signal-to-noise ratio decreases with redshift.

We determined the systematic error contribution to the calibration-corrected two-point function $\xi_{\pm}^{\text{cal}}(\theta)$ as follows. First, we took the likelihood distribution of the multiplicative calibration correction parameters in equation (14), as constrained by a fit of the calibration correction to the simulated data. This likelihood distribution is shown in Fig. 14 with the 68 per cent, 90 per cent and 95 per cent confidence intervals. We randomly sampled this distribution such that we had a sample of $N = 100$ (α, β) values that fully

represented the probability of the fit (i.e. 68 per cent of the pairs lay within the 1σ contour). We then calculated the two-point correction term $K_N(\theta)$ for each of the N resampled values and determined the variance of the calibration-corrected two-point function $\xi_{\pm}^{\text{cal}}(\theta)$ between the N differing corrections. This method allows us to isolate a systematic error covariance matrix for the two-point shear correlation function, allowing marginalization over the uncertainty on the calibration correction. Kilbinger et al. (2012) show that this has a negligible effect on cosmological parameter estimates from the two-point shear correlation function. The concept can be easily modified to cover any number of the weak lensing statistics that will be used to analyse CFHTLenS.

9 CONCLUSIONS AND FURTHER DEVELOPMENTS

We have presented the method of shear measurement developed for CFHTLenS, based on the earlier *lensfit* algorithm of Paper I and Paper II. The method fits PSF-convolved galaxy models to multiple galaxy images to estimate galaxy lensing shear. Significant enhancements have been made to the basic algorithm in order to create a method that can be applied to the survey as follows.

- (i) Creation of pixel models of the PSF on individual exposures, allowing for spatial and temporal variation of the PSF and for discontinuities between detectors in the imaging camera.
- (ii) Use of two-component models, disc plus bulge, with a prior distribution for bulge fraction estimated from published work.
- (iii) Measurement of likelihood from multiple dithered exposures, optimally combining information when the PSF differs between exposures, and automatically allowing any one galaxy to have differing numbers of useful exposures, where that number may vary because of the effect of detector and field boundaries coupled with the dither pattern. The multiple exposures are neither interpolated nor co-added, avoiding systematic effects that are otherwise introduced by those processes.
- (iv) Measurement of camera distortion from astrometric measurements, and correction of fitted models to optimally remove the effects of distortion from the shear measurement, rather than correction of the data using interpolation (i.e. interpolation of data is avoided).
- (v) Use of realistic priors on galaxy size, ellipticity and bulge fraction, based on published *HST* and other data.
- (vi) Improved algorithms for Bayesian marginalization over galaxy position, size, flux and bulge fraction.
- (vii) Modification of the shear estimator from a posterior probability estimator to a likelihood estimator to avoid anisotropy and bias caused by implementation of the sensitivity correction described by Paper I, but at the cost of needing to calibrate the effect of noise bias on the estimator.

We regard the issue of making measurements from non-interpolated individual exposures as being particularly important. If images are interpolated and co-added, the resulting non-stationary convolution kernel that has been introduced can have a severely detrimental effect on lensing shear measurement. In principle, the varying convolving kernel could be evaluated at each galaxy location and a correction applied to the PSF, but working from individual exposures avoids that complexity. Furthermore, the resulting co-added PSF is likely to have even more complexity than the PSF on individual exposures, and in general would be expected to lead to co-added PSFs with ‘twisted’ isophotes (ellipticity and centroid varying as a function of distance from the peak of the PSF), which

would cause problems for many published methods of shear measurement. As multiple exposures are usually dithered, and as the PSF may be discontinuous between detectors (as in MegaCam) and as there are gaps between detectors, any co-added PSF inevitably has discontinuous variations across the blurred regions between detectors in the dithered images. The ‘gap’ regions are not generally sufficiently large to allow independent measurement of the PSF from stars in those regions, so either such regions would need to be excised, or the PSF would need to be measured on individual exposures and the local PSF reconstructed appropriately. Excision of such problematic regions would lead to significant reduction in area for surveys with a large dither pattern. These problems are all avoided with the joint-fitting approach described here.

We have tested the method using two, independent, extensive simulations of CFHTLenS based on the GREAT and SkyMaker image simulation codes, but with galaxy parameters more closely matched to those found in the faint galaxy populations than have been assumed in previous lensing challenges. The faint galaxy population has galaxies of smaller size and a broader range of ellipticity than has been assumed in the STEP and GREAT challenges, which leads to a degradation in performance compared with what would have been naively expected from those simulations. When creating image simulations to test shear measurement methods, it is important to include real observational problems such as those described in this paper, and also to ensure that the simulated galaxy population mimics closely the observed distribution. With the current generation of surveys, we have only been able to match the coarse features of the population, namely size and ellipticity, and even there uncertainties remain over issues such as the redshift distribution of galaxy properties. Future, large-area surveys which aim to achieve significantly lower systematic uncertainties than in CFHTLenS should be concerned about our current lack of knowledge in this area.

Finally, the issue of noise bias has been discussed by a number of recent authors, and with current estimators we do need to calibrate the noise bias from realistic simulations. One hope for the future is that a properly formulated Bayesian estimator may be able to rigorously correct for noise bias without the need for calibration simulations.

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APPENDIX A: GALAXY MODELS

Galaxy models were created for CFHTLenS according to two basic types, bulge- or disc-dominated, as described in Section 3.2. Thus two surface brightness profiles were assumed, for each of the bulge and disc components of galaxies. For computational speed, these were two-dimensional distributions with radial surface brightness dependencies having Sérsic index equal to either 1 or 4, which were given elliptical isophotes by a 2D shear transformation. The transformation was defined so that the semi-major axis of a model component was invariant (unlike, for example, the GREAT08 or GREAT10 simulations; Bridle et al. 2010, Kitching et al. 2012). The rationale for this choice was that the ellipticity of disc galaxies is determined primarily by the inclination angle of the disc to the line of sight, in which case the semi-major axis is the best estimate of an

inclination-independent intrinsic size of the galaxy. When defining a prior based on the distribution of intrinsic size of galaxies, that prior should be applied to the distribution of semi-major axes.

Exponential and de Vaucouleurs surface brightness distributions are sharply peaked towards their centres, so if the surface brightness distributions are generated on a coarse pixel grid, the central pixel may become incorrectly dominant. The problem is worse for the de Vaucouleurs profile, $\rho(r) \propto \exp(-7.67(r/r_h)^{1/4})$: the central pixel is a factor of 2143 brighter than pixels at $r = r_h$. A partial solution to the problem is to generate surface brightness distributions on a finer-resolution grid; however for moderate oversampling factors (a factor of 11 oversampling was adopted for CFHTLenS), even that is not good enough. For CFHTLenS this problem was overcome by applying an integral constraint, requiring the sum of the model pixel values to equal the expected integral under the de Vaucouleurs profile, adjusting the central pixel value so that this condition was met. An alternative approach would be to use the ‘photon-shooting’ Monte Carlo method employed in the GREAT simulations.

Model components were convolved with the local PSF pixel-basis model at the location of each galaxy being fitted. Because both the galaxy model and the PSF contain Fourier modes above the Nyquist sampling limit of the observed image, we should seek to generate models at high resolution, convolve with a similarly high-resolution PSF, and only after that convolution, downsample to the observed sampling (noting that the downsampling should depend on the galaxy’s precise location). However, as the PSF had to be estimated from the stars on each exposure, it was not straightforward to generate reliable models with high-frequency modes correctly represented. In principle, such high-frequency information may be recoverable from the data because stars are observed at a variety of positions with respect to the pixel grid; however the PSF also varies between stars, and each star is itself a noisy measure of the downsampled PSF, so it was decided not to attempt to reconstruct oversampled PSF pixel-basis models. In effect, the PSF models adopted had their high-frequency modes aliased into the range of observed frequencies. In an attempt to make an approximate correction for that effect, an approximate PSF model was created, assuming circular isophotes, which comprised a Moffat profile fitted to the PSF near the centre of the field of view. The high-resolution galaxy models were first convolved by the Moffat model, before being downsampled. This downsampled galaxy model was then convolved with the actual PSF pixel-basis model for each galaxy, but with a multiplicative Fourier-space correction applied to remove the effect of the Moffat model. In the limit of band-limited data, an exact convolution by the true PSF model would have been achieved by this process. For the actual data, which may contain higher-frequency modes, some level of correction for those modes was obtained by this approximate method, but clearly a future improvement would be to reconstruct a high-resolution PSF pixel-basis model from the data and deal with the high-frequency modes more rigorously.

It is worth noting that, because the observed stars have been convolved with the detector’s pixel response function (usually assumed to be a top-hat function), the PSF model employed already included the effect of the mean pixel response. No further averaging across pixels was done.

A further problem with generating surface brightness distributions from sheared 2D distributions is that disc galaxies change their surface brightness profiles with inclination angle: edge-on galaxies have a surface brightness distribution along their minor axes which are best described as a sech^2 function (van der Kruit & Searle 1981). This is thought to be partly a result of the distribution of stars perpendicular to the disc and partly a result of obscuration associated

with the disc (Graham & Worley 2008). In order to attempt to address this, in one set of tests, the 2D disc models were replaced by a 3D model in which the model surface brightness had a sech^2 distribution perpendicular to the plane. Obscuration was assumed to be distributed the same as the stars (probably an oversimplification of the true obscuration), and the projected 2D distribution was obtained by integrating through the 3D distribution. The obscuration was set to match the fits reported by Graham & Worley (2008). However, when the tests for systematics, described in Section 7, were made, there appeared to be no significant improvement in the results. Inspection of typical PSF-convolved 3D models showed that they were indistinguishable from 2D models (once allowance for the apparent increase in observed scalelength had been made; see Graham & Worley 2008) and it was concluded that the extra computational expense of the 3D models was not justified. The likely reason for the lack of distinction between 2D and 3D models is that, in the CFHTLenS data, most galaxies are only marginally resolved, and the semi-minor axis in particular is much smaller than the size of the PSF. This test did, however, provide some reassurance that the results would not be sensitive to this aspect of the galaxy models.

The 2D models of disc-dominated galaxies that were used here had the ellipticity and orientation of the bulge component set equal to that of the disc. The 3D models that were also tested had realistic ellipticity gradients, caused by the bulge component being less flattened than the disc. The lack of sensitivity to the 2D or 3D galaxy structure arises because in the CFHTLenS data most galaxies are only marginally resolved, and hence the bulge ellipticity has little influence on the overall measurement, which is dominated by the more extended disc component. In future surveys with better resolution, such as the Euclid⁹ mission, ellipticity gradients must be taken account of (Bernstein 2010) and more complex galaxy models will likely be needed.

APPENDIX B: PRIOR DISTRIBUTIONS

B1 Galaxy scalelength

The disc and bulge scalelength priors are derived from the fits to *HST* WFPC2 observations of galaxies in the Groth strip by Simard et al. (2002), shown in Fig. 1 for the disc-dominated galaxies. The relation between the log of the median major-axis scalelength, r_d and i -band magnitude appears linear. A least-squares fit to the median values in the range $18.5 < i_{814} < 25.5$ yields

$$\log_e [r_d / \text{arcsec}] = -1.145 - 0.269(i_{814} - 23). \quad (\text{B1})$$

In applying this relation to CFHTLenS we assumed $i_{814} = i_{\text{CFHT}}$. Those measurements do not strongly constrain the functional form of the scalelength prior: we adopted a prior of the form

$$p(r) \propto r \exp(-(r/a)^\alpha)$$

where the value of a was adjusted to match the magnitude-dependent median of the Simard et al. fits. The value of α was also not strongly constrained by those data, so we adopted a value in the middle of a range that appears to give a good representation of the distribution, $\alpha = 4/3$. With this choice of α , $a = r_d/0.833$. One concern about the use of the WFPC2 data might be that large, low-surface brightness galaxies may have been missed owing to the limited surface brightness sensitivity of the observations. While we cannot exclude the possibility that some galaxies may be missed, we note that the

⁹ sci.esa.int/euclid

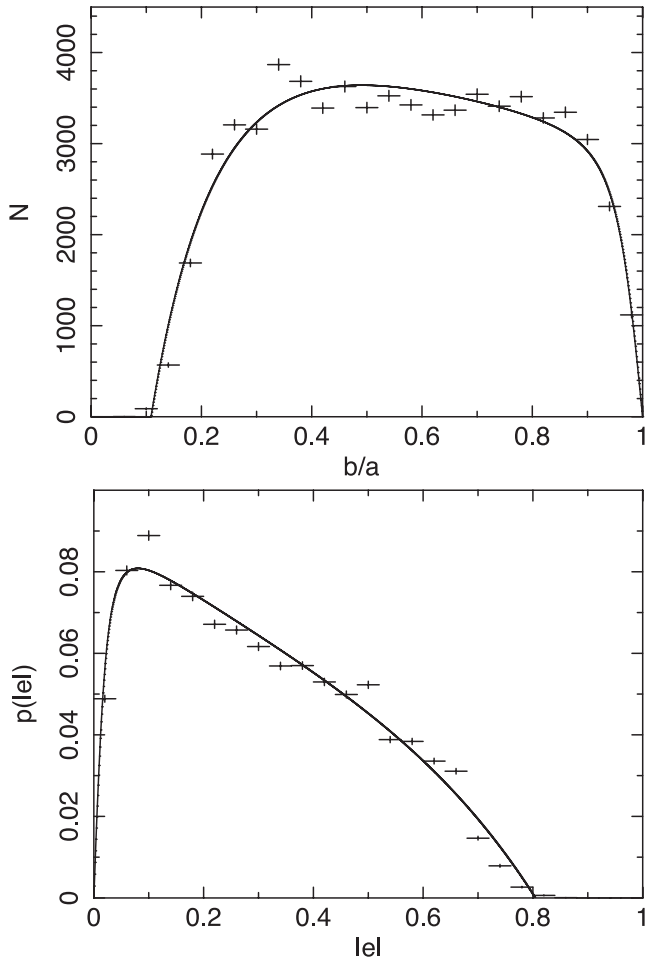


Figure B1. The observed distributions of axis ratio (upper) and ellipticity (lower) for disc galaxies with $M_g < -19$ in SDSS, compared with the model function adopted for the CFHTLenS ellipticity prior (solid curves).

WFPC2 data extend to significantly fainter galaxy magnitudes than CFHTLenS (Fig. 1) which mitigates against the possible loss of low surface brightness galaxies at the CFHTLenS depth. The use of the median galaxy size also guards against the loss of low surface brightness galaxies. *HST* observations are currently the only way of reliably probing the size distribution of galaxies on scales ~ 0.1 arcsec.

The distribution of bulge scalelengths (or equivalent half-light radii) is more problematic. In principle, we need to consider the disc- and bulge-dominated populations separately. The sizes of bulges assumed in the models for the disc-dominated galaxies have been discussed in Section 3.2. The distribution of sizes of bulge-dominated galaxies might, in principle, be obtainable from the Simard et al. (2002) analysis; however these galaxies are a minority of the population, and given the degeneracies that exist between scalelength, ellipticity and surface brightness profile, in noisy data, it is difficult to obtain a robust estimate of the size distribution with present data. In the local Universe, half-light radii of the most luminous early-type galaxies are about 50 per cent larger than those of late-type galaxies, but at the galaxy luminosities appropriate for CFHTLenS, $M \sim -19$, they are about a factor of 2 smaller (Shen et al. 2003). The distribution of early-type scalelengths is not well-known, but we note that the details of the early-type prior are not too important as they only account for about 10 per cent of the galaxies. We

therefore assign the same distributions to the half-light radius of the bulge-dominated population as to the exponential scalelength of the disc-dominated population, so that the half-light radii of early-type galaxies are 0.6 times the half-light radii of late-type galaxies, consistent with Shen et al. (2003). Future work should seek to refine our knowledge of these distributions, particularly beyond the local Universe.

B2 Galaxy ellipticity

Fits to faint galaxies do not strongly constrain the distributions of galaxy ellipticity, so we make the assumption that faint galaxies have the same distribution of disc ellipticity as low-redshift galaxies of the same luminosity in the Sloan Digital Sky Survey (SDSS): disc ellipticities are determined largely by the disc's inclination, thickness and optical depth at the waveband of observation, the latter being luminosity dependent (Unterborn & Ryden 2008). Accordingly, we fit the ellipticity distribution of 66 762 SDSS disc-dominated galaxies selected from DR7 (Abazajian et al. 2009) to have g -band absolute magnitude $M_g \leq -19$, to correspond in luminosity and waveband to galaxies selected at $z \leq 24.5$ in CFHTLenS, for which the median redshift is estimated as $z \approx 0.66$ (Hildebrandt et al. 2012). The functional form adopted was

$$p(e) = \frac{Ae(1 - \exp[-\frac{(e - e_{\max})}{a}])}{(1 + e)(e^2 + e_0^2)^{1/2}} \quad e < e_{\max} \quad (\text{B2})$$

where $e \equiv |e|$ and where the maximum ellipticity cutoff $e_{\max} = 0.804$, obtained from fitting to the distribution, arises primarily because of the finite thickness of galaxy discs. The circularity parameter $e_0 = 0.0256$, also obtained from fitting, reflects the intrinsic non-circular symmetry of face-on disc galaxies: even face-on galaxies only rarely appear to have circular isophotes. The fitted value for the dispersion was $a = 0.2539$. The normalization parameter A was determined numerically. The distributions of axis ratio and ellipticity for disc galaxies are shown in Fig. B1.

In principle, the ellipticity distribution should be a function of galaxy luminosity. However, the prior distribution only has any significant effect for the lowest signal-to-noise ratio galaxies: high signal-to-noise ratio galaxies have well-defined likelihood values and the posterior probability distribution is largely unaffected by the prior. Accordingly we chose a galaxy luminosity appropriate for the faintest galaxies in CFHTLenS selected at the median redshift, and fixed that prior distribution to be the same for all galaxies.

A prior for the ellipticity distribution of bulge-dominated galaxies is not readily obtained. In the absence of better information, we adopted a fit to the ellipticity distribution of bulge-dominated galaxies in Simard et al. (2002) of the form

$$p(e) \propto e \exp[-be - ce^2] \quad (\text{B3})$$

where the best-fitting values were $b = 2.368$, $c = 6.691$ and where the normalization was determined numerically.

B3 Galaxy bulge fraction

Likewise, the bulge fraction is not well-constrained by current data, and even in the nearby universe the bulge fractions deduced for bright galaxies are dependent on the models used to fit them (Graham & Worley 2008). A reasonable distribution to assume, and which is also consistent with fits to faint galaxy samples such as Schade et al. (1996), is that there are two populations. Bulge-dominated galaxies were fit by purely de Vaucouleurs profiles, and

account for around 10 per cent of the galaxy population. Disc-dominated galaxies were assumed to make up the remainder, and we assumed the bulge fraction $f = B/T$ to have a truncated normal distribution in f with its maximum at $f = 0$ and with $\sigma = 0.1$, truncated so that $0 \leq f < 1$. The bulge fraction was marginalized over as described in the main text. The bulge fraction prior was independent of other parameters.

B4 Prior iteration

In Paper II we proposed an iterative scheme to improve the accuracy of the prior probability distributions using the observed sample of galaxies. In the STEP simulations the procedure worked well. However, when applied to the CFHTLenS data, the iterations did not readily converge. It appeared that the main cause of this problem was that in CFHTLenS, unlike STEP, the galaxies are only marginally resolved, so that there are strong degeneracies between size, ellipticity and surface brightness profile. Given that the data are also noisy, it is not surprising that the iterations tended to converge to arbitrary solutions.

Fortunately for CFHTLenS, the available *HST* and SDSS data were sufficient to achieve adequate estimates of prior distributions. In future surveys such as Euclid, it may be necessary to use the survey itself to define the prior distributions. However, in that survey, galaxies are expected to be rather well resolved, and it is likely that iterative improvement of the prior would be effective.

APPENDIX C: BIAS IN THE MEAN LIKELIHOOD ESTIMATOR

Here we discuss the expected bias in the mean ellipticity measured for galaxies for a simplified example of a galaxy with a Gaussian surface brightness profile convolved with a Gaussian PSF. Because of the unrealistic Gaussian surface brightness profiles, the results calculated cannot be transferred to the actual CFHTLenS analysis: however, the Gaussian model provides a useful illustration of the origin of bias in the mean ellipticity estimate. This work was done independently of that of Refregier et al. (2012) and Melchior & Viola (2012) but reaches similar conclusions. The work described here investigates particularly the bias on the mean likelihood estimator of ellipticity, rather than the maximum likelihood estimator, and discusses also the correlation with the PSF that is induced by the presence of noise. Rather than calculate bias in the Gaussian model using a Fisher matrix approach, we explicitly calculate the shape of the likelihood surface for some simple examples.

An elliptical object may be viewed as a sheared version of an isotropic object, and we define the isotropic coordinates \mathbf{x}'_i of pixel i in terms of the observed coordinates \mathbf{x}_i as $\mathbf{x}'_i = \mathbf{A}\mathbf{x}_i$, where

$$\mathbf{A} = \begin{pmatrix} 1 - e_1 & -e_2 \\ -e_2 & 1 + e_1 \end{pmatrix} \quad (\text{C1})$$

and where e_1 and e_2 are the two components of ellipticity.

We define a Gaussian surface brightness profile as $f(\mathbf{x}_i) = \exp[-\frac{1}{2}\mathbf{x}_i^T \mathbf{C}^{-1} \mathbf{x}_i]$ where the pixel surface brightness covariance matrix is given by

$$\mathbf{C} = \frac{s^2}{(1 - e^2)^2} \begin{pmatrix} 1 + 2e_1 + e^2 & 2e_2 \\ 2e_2 & 1 - 2e_1 + e^2 \end{pmatrix} \quad (\text{C2})$$

where s is a size parameter and $e = |e| = \sqrt{e_1^2 + e_2^2}$.

If a Gaussian galaxy with pixel surface brightness covariance matrix \mathbf{C}_G is convolved with a Gaussian PSF with covariance matrix

\mathbf{C}_P then the resulting surface brightness distribution is Gaussian with covariance matrix $\mathbf{C}_G + \mathbf{C}_P$.

If we fit a model surface brightness distribution, f^m to some pixelated data f , the log(likelihood) of the fit may be written as

$$\log \mathcal{L} = \frac{1}{2\sigma^2} \frac{[\sum_i f_i^m f_i]^2}{\sum_i f_i^{m2}} \quad (\text{C3})$$

where the sum is over pixels i , the background noise has rms σ and is assumed stationary, the model galaxy's amplitude has been marginalized over and an arbitrary additive constant has been neglected (Paper I).

If the data correspond to some Gaussian galaxy convolved with a Gaussian PSF, with independent, stationary Gaussian noise added to each pixel, we may write $f_i = \mathcal{N} \exp[-\frac{1}{2}\mathbf{x}_i^T (\mathbf{C}_G + \mathbf{C}_P)^{-1} \mathbf{x}_i] + n_i$ for pixel i , where \mathcal{N} is a normalizing constant and n_i is a Gaussian random variate representing the pixel noise. Neglecting the bias caused by the term n_i , the galaxy's profile-weighted signal-to-noise ratio is given by $v_{\text{SN}} = (\sum_i f_i^2)^{1/2}/\sigma$ which allows us to define the normalization constant as

$$\mathcal{N} = \frac{\sigma v_{\text{SN}}}{\pi^{1/2} |\mathbf{C}_G + \mathbf{C}_P|^{1/4}}. \quad (\text{C4})$$

Hence the log(likelihood) of the data given some model galaxy with covariance matrix \mathbf{C}_M , convolved with the PSF, may be written as

$$\log \mathcal{L} = \frac{\left[\sum \left(\mathcal{N} e^{-\frac{1}{2}\mathbf{x}_i^T (\mathbf{C}_G + \mathbf{C}_P)^{-1} \mathbf{x}_i} + n_i \right) e^{-\frac{1}{2}\mathbf{x}_i^T (\mathbf{C}_M + \mathbf{C}_P)^{-1} \mathbf{x}_i} \right]^2}{2\sigma^2 \sum e^{-\mathbf{x}_i^T (\mathbf{C}_M + \mathbf{C}_P)^{-1} \mathbf{x}_i}} \quad (\text{C5})$$

and so the expectation value of the log(likelihood) for this particular model M is

$$\langle \log \mathcal{L} \rangle = \frac{2\pi \mathcal{N}^2 \left| \left((\mathbf{C}_G + \mathbf{C}_P)^{-1} + (\mathbf{C}_M + \mathbf{C}_P)^{-1} \right)^{-1} \right|}{\sigma^2 |\mathbf{C}_M + \mathbf{C}_P|^{1/2}} + \left\langle \frac{[\sum n_i f_i^M]^2}{\sigma^2 \sum f_i^{M2}} \right\rangle \quad (\text{C6})$$

where $f_i^M = e^{-\frac{1}{2}\mathbf{x}_i^T (\mathbf{C}_M + \mathbf{C}_P)^{-1} \mathbf{x}_i}$ and $\langle \sum n_i f_i^M \rangle = 0$. We may use the identity $(\mathbf{A}^{-1} + \mathbf{B}^{-1})^{-1} = \mathbf{A}(\mathbf{A} + \mathbf{B})^{-1}\mathbf{B}$ in the first term. The second term has a value of unity, since $\langle [\sum f_i^M n_i]^2 \rangle = \sigma^2 \sum f_i^{M2}$, and ignoring this irrelevant additive constant, we obtain

$$\begin{aligned} \langle \log \mathcal{L} \rangle &= \frac{2\pi \mathcal{N}^2 |\mathbf{C}_G + \mathbf{C}_P| |\mathbf{C}_M + \mathbf{C}_P|^{1/2}}{\sigma^2 |\mathbf{C}_G + 2\mathbf{C}_P + \mathbf{C}_M|} \\ &= \frac{2v_{\text{SN}}^2 |\mathbf{C}_G + \mathbf{C}_P|^{1/2} |\mathbf{C}_M + \mathbf{C}_P|^{1/2}}{|\mathbf{C}_G + 2\mathbf{C}_P + \mathbf{C}_M|} \\ &= 2v_{\text{SN}}^2 |(\mathbf{C}_G + \mathbf{C}_P)^{-1}|^{1/2} |(\mathbf{C}_M + \mathbf{C}_P)^{-1}|^{1/2} \\ &\quad |(\mathbf{C}_G + \mathbf{C}_P)^{-1} + (\mathbf{C}_M + \mathbf{C}_P)^{-1}|. \end{aligned} \quad (\text{C7})$$

We can repeat this calculation for a set of models M , each with a different ellipticity or size, and hence build up the expected likelihood surface as a function of model ellipticity. Fig. C1 shows contours of $\langle \log_e \mathcal{L} \rangle$ for an example where the galaxy ellipticity was $e_g = \{0.3, 0\}$ and the PSF ellipticity was $e_{\text{PSF}} = \{0, 0.3\}$, with size $\sigma = 3$ pixels for both galaxy and PSF. The choice of PSF ellipticity is extreme, much larger than the PSF ellipticity in CFHTLenS, to better illustrate the bias. Although the location of the maximum

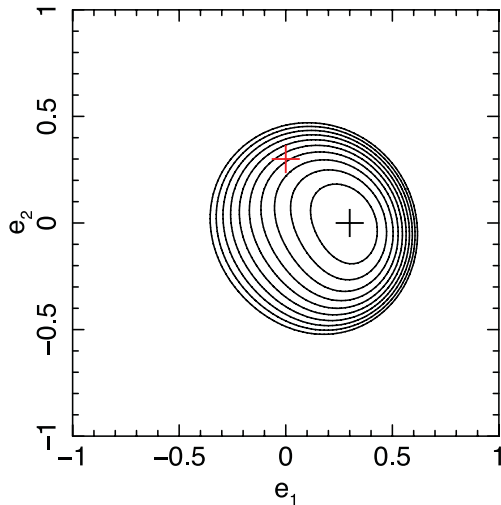


Figure C1. Contours of $\log_e \mathcal{L}$ for the example Gaussian model and PSF described in the text. The cross at $e_1 = 0.3$, $e_2 = 0$ denotes the ellipticity of both the test galaxy and the maximum likelihood. The cross at $e_1 = 0$, $e_2 = 0.3$ denotes the ellipticity of the PSF. Contours are plotted at intervals of 0.5 in $\log_e \mathcal{L}$ below the maximum, for a galaxy with signal-to-noise ratio $\nu_{\text{SN}} = 20$. The contour interval scales as ν_{SN}^2 .

is unbiased, the likelihood distribution has a substantial asymmetric distortion, both towards $e = 0$ and towards the PSF ellipticity, which results in the mean ellipticity being biased. In these illustrations, the size of the model galaxy has been held fixed (i.e. the galaxy size is assumed to be known). If this assumption is relaxed, and the galaxy size is also treated as a free parameter, the likelihood surface changes shape, with a bias in both the mean and the maximum likelihood, and an excess of likelihood at extreme ellipticity can be produced (see also Melchior & Viola 2012; Refregier et al. 2012). We note that the application of a Bayesian prior would suppress that excess likelihood, but we do not consider that effect here.

The upper panel of Fig. C2 shows the shear bias, plotted as $1 + m$, expected in this simple model, for fixed galaxy size, for a simplified case where a population of randomly oriented Gaussian galaxies, all with intrinsic ellipticity $e_g = 0.25$ and constant shear $\mathbf{g} = \{0.05, 0\}$, have been measured with a Gaussian PSF, the same size as the galaxy, with $\mathbf{e}_{\text{PSF}} = \{0, 0.1\}$. The lower panel shows the expected cross-correlation between the galaxy and PSF. It can be seen that both measures show significant departures from zero bias at low signal-to-noise ratio, $\nu_{\text{SN}} \lesssim 20$. In this model, even larger biases are seen for Gaussian galaxies that are smaller than the PSF.

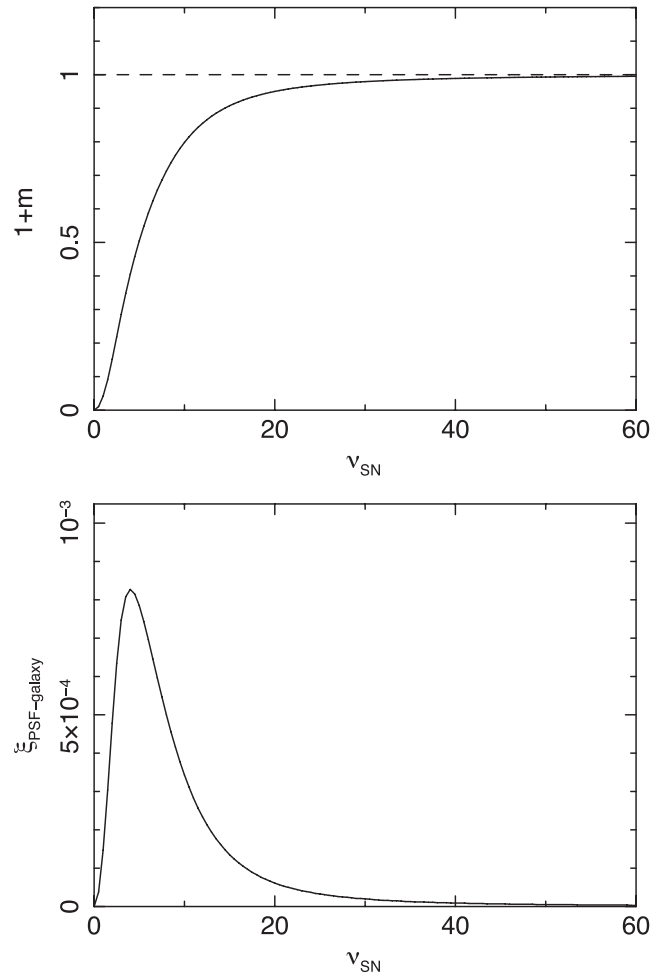


Figure C2. Upper panel: the expected shear bias, $1 + m$, arising from the bias in the mean ellipticity for the Gaussian model described in the text, as a function of signal-to-noise ratio ν_{SN} . Lower panel: the expected PSF–galaxy cross-correlation signal, $\xi_{\text{PSF-galaxy}}$, in the absence of shear, as a function of ν_{SN} .

In CFHTLenS, many galaxies are smaller than the PSF; however the significant difference in surface brightness profiles of both the galaxies and the PSF means that these results cannot be transferred quantitatively to the actual data. Our approach, described in Section 8, was instead to measure the effects from realistic simulations designed to match the data.

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