

The frequency dependence of scattering imprints on pulsar observations

M. Geyer^{1,3*} and A. Karastergiou^{1,2,3}

¹*Astrophysics, University of Oxford, Denys Wilkinson Building, Keble Road, Oxford OX1 3RH, UK*

²*Physics Department, University of the Western Cape, Cape Town 7535, South Africa*

³*Department of Physics and Electronics, Rhodes University, PO Box 94, Grahamstown 6140, South Africa*

Accepted XXX. Received YYY; in original form ZZZ

ABSTRACT

Observations of pulsars across the radio spectrum are revealing a dependence of the characteristic scattering time (τ) on frequency, which is more complex than the simple power law with a theoretically predicted power law index. In this paper we investigate these effects using simulated pulsar data at frequencies below 300 MHz. We investigate different scattering mechanisms, namely isotropic and anisotropic scattering, by thin screens along the line of sight, and the particular frequency dependent impact on pulsar profiles and scattering time scales of each. We also consider how the screen shape, location and offset along the line of sight lead to specific observable effects. We evaluate how well forward fitting techniques perform in determining τ . We investigate the systematic errors in τ associated with the use of an incorrect fitting method and with the determination of an off-pulse baseline. Our simulations provide examples of average pulse profiles at various frequencies. Using these we compute spectra of τ and mean flux for different scattering setups. We identify setups that lead to deviations from the simple theoretical picture. This work provides a framework for interpretation of upcoming low frequency data, both in terms of modelling the interstellar medium and understanding intrinsic emission properties of pulsars.

Key words: Pulsars: general, scattering, ISM: structure

1 INTRODUCTION

Observed radio pulses from pulsars pass through the ionised interstellar medium (ISM) before terrestrial detections are made. As the waves propagate through the ISM, inhomogeneities in the electron densities of dense, ionised regions cause the radio waves to scatter. This scattering is observed through, amongst other effects, the frequency dependent temporal broadening of the received radio pulses. The temporal broadening typically leads to characteristic exponential *tails* in averaged pulsar profiles (e.g. Löhmer et al. 2001). Understanding the frequency dependence of the temporal broadening allows us to investigate the properties of the scattering media. Compensating for scattering effects can be used to infer the intrinsic radio emission characteristics of pulsars.

The scattering tails of the broadened profiles are routinely modelled with exponential functions that take the

form $\tau^{-1}e^{-t/\tau}$, where τ is referred to as the *characteristic scattering time*. This functional form, as described in detail in Sec. 2.1, is based on the assumption of a single, thin scattering screen extending infinitely transverse to the line of sight to the pulsar. The screen is invoked as an approximation of the combined scattering that the pulse signal undergoes in the extended ISM along the line of sight (Scheuer 1968; Cronyn 1970; Williamson 1972, 1973). The characteristic scattering time τ depends on the properties of the scattering screen as well as its location along the line of sight (see eq. 2) and it is maximised for a midway screen.

Apart from approximating the overall scattering geometry to a single scatterer, historic studies have also investigated distinct electron density distribution models for the screen. These include an isotropic Gaussian electron column density wavenumber spectrum, which leads to a temporal broadening dependence on frequency, $\tau \propto \nu^{-4}$, (Lee & Jokipii 1976; Lang 1971). Alternatively power law wavenumber spectra, characterised by an inner and outer cut-off length scale, are adopted to model the plasma inhomogeneity.

* E-mail: marisa.geyer@physics.ox.ac.uk

geneities. Most often a Kolmogorov power spectrum, which describes the turbulence of a neutral gas, is preferred (Rickett 1977; Rickett et al. 1984; Rickett 1990). A Kolmogorov spectrum will have a frequency dependence of $\tau \propto \nu^{-4.4}$, for wavenumbers lying between the inner and the outer cut-off length scales¹.

The scale at which electron density inhomogeneities lead to strong diffractive scintillation and consequently the observed pulse broadening analysed in this paper, is set by the field coherence scale, $s_0 = 1/(k\theta)$, with k the radio wavenumber (Rickett 1990). For strong diffractive scintillation $s_0 \ll L_F$, where $L_F = (D/k)^{1/2}$, is the Fresnel scale (e.g. Smith & Thompson 1988. Refer to Fig. 1 for descriptions of θ and D).

Estimations of τ values from pulsar profile observations are done by either deconvolution of the folded pulse profile with a broadening function such as a one-sided exponential function through e.g. the *CLEAN*-based algorithms (Bhat et al. 2003), or by forward modelling which convolves a template unscattered pulse profile (or alternatively an observed high frequency pulse profile) with a parametrised broadening function and compares the model to the observed profile (e.g. Löhmer et al. 2001, 2004). It should be noted that using a high frequency observation as a template will, for pulsars that exhibit intrinsic frequency evolution, introduce biases to measurements of scattering parameters. A deconvolution method does not suffer from this effect.

Alternatively τ can be estimated by means of the *scintillation bandwidth*, δf (Lyne & Rickett 1968). The scattering process imparts frequency dependent phase changes on the passing radio waves. An interference pattern will emerge for phase changes of the order $2\pi\delta f\tau \approx 1$ radian. A measurement of δf , over which the interference effects are observed in the dynamic spectrum, can therefore lead to an approximation of τ (Cordes 1986).

Using the ISM transfer function implied by the thin screen model, results in fits of τ in observed data that follow the expected spectral dependence (Armstrong et al. 1995). However, deviations from the theoretical spectral dependences have also been proposed and observed - both with spectral index α (where $\tau \propto \nu^{-\alpha}$) smaller than predicted (Bhat et al. 2004; Löhmer et al. 2001; Lewandowski et al. 2015) and larger than predicted (Rickett 1990; Lambert & Rickett 1999; Tuntsov et al. 2012). Different models will lead to different values for α , but the critical minimum for strong diffractive scattering, is $\alpha = 4$ (Romani et al. 1986).

Anisotropic scatterers, i.e. media for which the distribution of scattering angles exhibit directionality have been considered as causes for steeper spectra (Tuntsov et al. 2012). They were first invoked to explain organised patterns in the dynamic spectra of pulsar observations (Gupta et al. 1994). Stinebring et al. 2001 found that complex patterns in dynamic spectra are often related to parabolic arcs in the secondary (power) spectra of pulsar observations, which can be explained by anisotropic scattering (Walker et al.

2004). There has also been some successes in associating anisotropy with observations of Intra Day Variability (IDV) (Bignall et al. 2003; Tuntsov et al. 2012). Extreme Scattering Events (ESE's) as observed by Fiedler et al. (1987) and analysed by e.g. Walker (2007) have motivated works studying the impact of alternative scatterers, such as self-gravitating, AU-sized ionised gas clouds. A one dimensional Kolmogorov power spectrum in electron density distribution is often used to model extreme anisotropy. This model suggests $\alpha = 8$.

On the contrary, weaker dependences on frequency have been attributed to screens which have a limited physical size rather than being infinite transverse to the line of sight as most models assume (Cordes & Lazio 2001). Beyond the edges of such a screen, no radio waves coming from the pulsar will be bent back into the line of sight of the observer, leading to no further scatter broadening with increasing wavelength. Such configurations also lead to a loss in observable flux.

A physical inner scale of the scattering screen, that is smaller than the diffractive scattering scale, will restrict diffraction to a maximum scattering angle. This too will lead to less power observed at large angles, or alternatively flatter τ spectra with lower power law indices, at long wavelengths. In the case of an inner scale however, flux is preserved (Rickett et al. 2009).

Until recently, characterising the frequency dependence of temporal scatter broadening was done using pulsar observations taken at different receivers and telescopes at different epochs. With the construction of telescopes such as LOFAR (van Haarlem et al. 2013) and the MWA (Tingay et al. 2012), we have the advantage of studying pulsars at low frequencies, where scattering is most pronounced, using broad bands (see Stappers et al. 2011 for an outline of LOFAR's pulsar observing modes). This allows accurate measurements of τ spectra and simultaneous studies of the flux spectrum and profile evolution of the pulsar. Analysis of broad band data will therefore test our current understanding of pulsar scattering.

Recent works showing average pulse profiles in the LOFAR bands (30 - 90 MHz and 110 - 190 MHz) confirm the expectation of scatter broadened pulse shapes at these frequencies (Bilous et al. 2015; Pilia et al. 2015).

In this paper we provide a framework based on simulations, for interpreting observations such as the ones presented in the LOFAR works. This can be used to gain insight on the size, location and nature of scattering screens.

The paper is split into sections as follows. In Sec. 2 we describe the general scattering setup, the simulated data and the techniques by which we fit the data to retrieve τ spectra. The mathematical formalism for examples of the broadening functions for isotropic and anisotropic scattering is also presented in this section. In Sec. 3 we present the results of our experiments. These include example profile shapes as well as τ and flux spectra for infinite and finite screens respectively. We end with a discussion of our results and their relevance to broad band observations in Sec. 4, followed by a short conclusion.

¹ Note that for power-law spectra, where $\tau \propto \nu^{-\alpha}$, the fluctuation spectral index β is often quoted instead of the spectral index α . For $\alpha \geq 4$, they are related through: $\alpha = 2\beta/(\beta - 2)$. A Kolmogorov spectrum with $\alpha = 4.4$ has $\beta = 11/3$. Also note that the symbols α and β are sometimes used interchangeably.

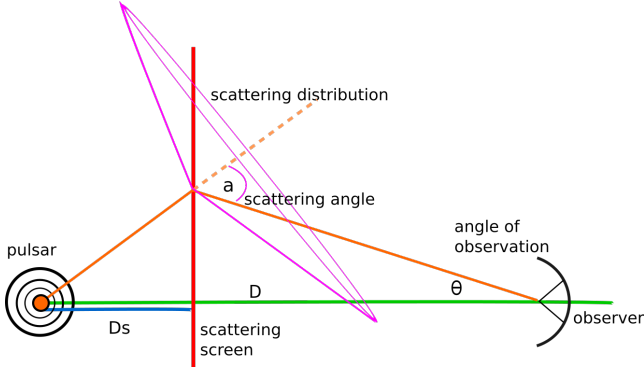


Figure 1. A schematic representation of radio waves scattered by a single scattering screen, with D_s the distance from the pulsar to the screen and D the overall distance from the pulsar to the observer. The scattering angle α , (drawn from some probability distribution, here represented by the magenta cone), is the angle by which the ray is scattered away from its straight line trajectory and θ the angle at which we observe the scattered ray.

2 SIMULATED DATA AND FITTING TECHNIQUES

2.1 Scattering geometry

The scattering geometry considered in this work is shown in Fig. 1. The radio emission from a pulsar at a distance D from the observer hits a scattering screen at a distance D_s (measured from the pulsar). The photons are scattered by the screen such that those refracted back to the observer travel along paths of different lengths. The additional path lengths will cause the observed scatter broadening.

The geometrical relationship between the delay t , with respect to the direct line of sight, and the scattered angle α for a single thin screen is given by $t = D_s(1 - D_s/D)|\alpha|^2/2c$ (Cronyn 1970). The scattered angle relates to the observed ray angle θ by $\theta = \alpha(D_s/D)$.

Following the analysis put forth in Cordes & Lazio (2001) we present general expressions for the probability density function f_t , which describes the probability that a given photon from the pulsar will be delayed by a time t , in Sec. 2.2.2. We simulate scattering imprints on pulse profiles by convolving an intrinsic pulse train with a given temporal broadening function f_t . Thereafter, we fit the simulated data with a set of methods and test how accurately they reproduce the simulated values of τ .

2.2 Simulated data

2.2.1 Simulated Intrinsic Pulse

The intrinsic pulse profile in our simulated data consists of a single Gaussian component with a chosen duty cycle and pulse period (P). When necessary, multiple pulse periods are stitched together to form a pulse train. We investigate a typical slow pulsar with $P = 1$ s and a duty cycle of 2.5% and fast pulsar with $P = 20$ ms with a duty cycle of 10%.

The mean flux spectral index for all the simulations is chosen to be $\gamma = 1.6$, where the mean flux is $S_m = \nu^{-\gamma}$. We use γ as the flux spectral index and α as the τ spectral index to avoid confusion. The mean flux of the pulsar is normalised such that it has a maximum value of unity for

the lowest studied frequency. The use of a Gaussian template is motivated by the fact that a large fraction of pulsar profiles are relatively well approximated by this shape (e.g. Karastergiou & Johnston 2007; Johnston & Weisberg 2006). The drawback is that a single Gaussian shape is not representative of all pulsar profiles, especially not of older, slower pulsars. In Sec. 3.3 we briefly study possible scattering impacts on a more complex profile shape based on the average profile of PSR B1237+25.

2.2.2 Simulated Scattering Screen

We consider nearby sources by choosing a standard distance of $D = 3.0$ kpc from the observer to the pulsar with a mid-way scattering screen at $D_s = 1.5$ kpc. We assume we can treat the scattering screen as a cold plasma for which the scattering angle $\alpha \propto \nu^{-2}$, with ν the observing frequency (Lang 1971). For a distribution in scattering angles the standard deviation will follow the relation $\sigma_\alpha \propto \nu^{-2}$. The scattering strength of the screen is fixed by choosing a proportionality constant in this relation. We use $\sigma_\alpha = 3$ mas at 1 GHz, such that $\sigma_\theta = 1$ mas at distances $D = 1$ kpc and $D_s = 1/3$ kpc, as suggested in Cordes & Lazio (2001).

We present two analytic expressions for temporal broadening functions associated with isotropic and anisotropic scattering by an infinite thin screen, and thereafter discuss a ray-tracing experiment by which numerical broadening functions for alternative configurations can be built up.

Isotropic broadening functions

The temporal broadening function for a two dimensional scattering screen that scatters photons isotropically and can be described by a circularly symmetric Gaussian probability distribution in the scattering angle, takes the following form

$$f_t = \tau^{-1} e^{-t/\tau} U(t) \quad (1)$$

$$\tau = D'_s \sigma_\alpha^2 / c \quad (2)$$

$$D'_s = D_s(1 - \frac{D_s}{D}), \quad (3)$$

where σ_α is the standard deviation in scattering angle α in each coordinate direction and c is the speed of light (Cordes & Lazio 2001). The unit step function, $U(t)$, ensures that we only consider time $t > 0$. Eq. (1) is valid for a scattering screen that is infinite transverse to the line of sight. Such an infinite screen will conserve flux coming from the pulsar; every ray scattered out of the line of sight is compensated by another ray that hits the screen and is scattered back into the line of sight. For finite screens the total flux of the source may not reach the observer depending on the physical extent of the screen and the observing frequency. (See discussion in the Sec. 3.3)

Anisotropic broadening functions

An anisotropic scattering screen would, in general, lead to the observation of elongated images of point sources behind such screens. These effects have, for example, been

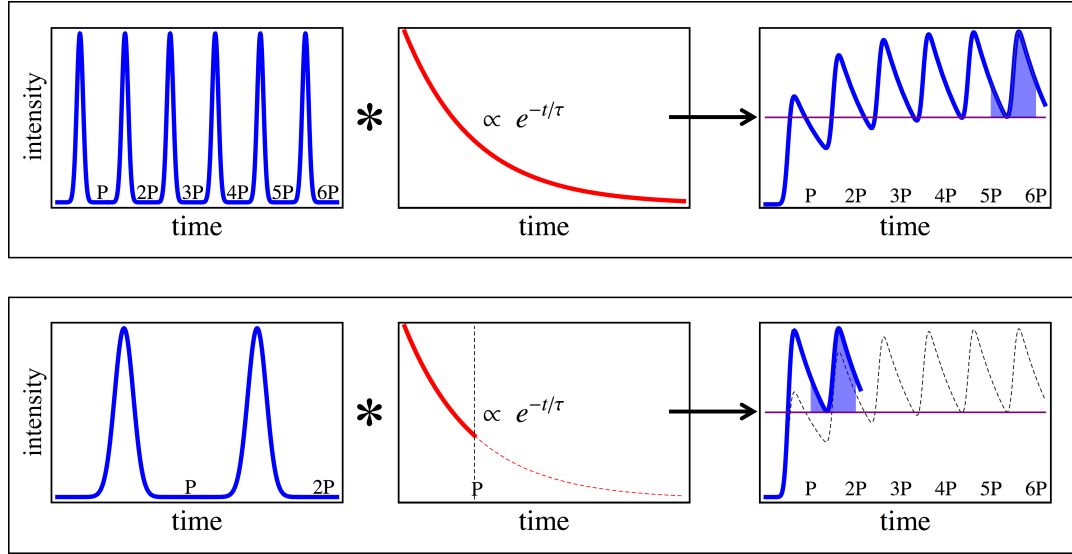


Figure 2. A schematic of the different methods for generating noiseless scattered profiles, for simulations or forward fitting. The top panel shows the *long train* method, and the bottom panel the preferred *train* method.

observed in the imaging of pulsar B0834+06 at 327 MHz (Briskin et al. 2010).

We consider an anisotropic scattering mechanism that can be modelled by an asymmetric Gaussian distribution in scattering angles. The temporal broadening function will now be dependent on two different characteristic scattering times, say τ_x and τ_y , corresponding to the two dimensions that define the scattering screen, and has the form

$$f_t = \frac{1}{\sqrt{\tau_x \tau_y}} e^{-\frac{t}{2}(\frac{1}{\tau_x} + \frac{1}{\tau_y})} I(0, \frac{t}{2}(\frac{1}{\tau_x} - \frac{1}{\tau_y})), \quad (4)$$

where I is the Bessel function of the first kind. The scattering times are defined as in eq. (2) with $\tau_{x,y} = D'_s \sigma_{a,x,y}^2 / c$ which reduces to the isotropic case for $\tau_x = \tau_y$.

Numerical broadening functions

In Cordes & Lazio (2001) it was shown that the broadening function for a disk-shaped isotropic scattering screen centred on the line of sight can be created by amending eq. (1) with a unit step function, such that the broadening function becomes $f_t U(t_{max} - t)$. This means that beyond a maximum observation angle θ_{max} , no photons are scattered back into the observer's line of sight. The maximum angle has a geometric path length correspondence to a maximum time delay t_{max} , as measured with respect to the direct line of sight.

We expand on these effects by means of a ray-tracing code that allows us to place a screen of any size between the source and the observer. We simulate a rectangular screen as a toy model: by specifying a screen's width and height, its position along the line of sight (D_s) and its two-dimensional offset with respect to the line of sight, the scattering geometry is defined.

The temporal broadening function is constructed by calculating the probability that a given photon hitting the truncated screen will reach the detector. Binning these probabilities according to the associated arrival times for each pho-

ton allows us to build up the broadening function with a required time resolution.

2.2.3 Scattered profiles

Fig. 2 gives a schematic overview of two methods for producing scattered profiles. The method shown in the top panel represents the most intuitive way of creating folded scattered pulse profiles: a train of Gaussian pulses are convolved with a broadening function resulting in a train of broadened pulse shapes. A broadened pulse towards the end of the train (shaded) is representative of the average pulse profile. We refer to this method as the *long train* method.

An alternative method is shown in the bottom panel of Fig. 2. This method, which we simply call the *train* method, convolves a Gaussian pulse train of two pulses with a truncated temporal broadening function defined over a single pulse period (solid line). The second broadened pulse is then extracted as the averaged pulse profile (shaded). The outcome of the long train method (scaled by a constant factor and overplotted in dashed lines in the right-most frame) converges to the same average profile shape after a sufficient period of time. Fig. 3 shows the convergence of the two methods. The residuals between the last broadened pulse resulting from the train method and the long train method are shown to decrease as the pulse train length (N) in the latter is increased.

We note that the long train method requires care in the choice in the length of both the pulse train and the exponential, to ensure that the convolution converges to a constant set of pulses. Typically the required pulse train length increases with an increase in scattering, thereby also increasing the computation time. The advantage of the train method is that it's both computationally less expensive and its accuracy is not dependent on the length of the broadening function or the pulse train (beyond 2 pulses). This motivates us to make use of the *train* method throughout this paper.

Finally, after generating the noiseless scattered profiles,

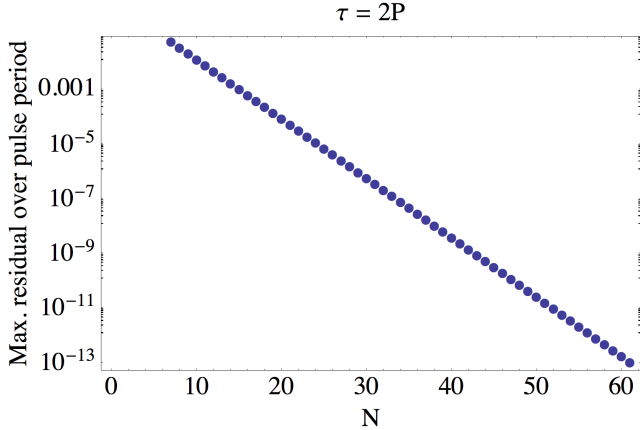


Figure 3. The residual differences of the last broadened pulses, as calculated from the analytic expressions for the convolutions in the train method and the long train method respectively, after having corrected for a constant scaling factor equal to the integral of the broadening function up to $t = P$ (viz. $1 - e^{-P/\tau}$). The residuals are normalised to the area of the pulse and expressed as a function of the number of pulses (N) used in the pulse train of the long train method. The broadening function for the long train method is defined for all t , whereas for the train method the broadening function is truncated at the pulse period. In this example $P = 0.4$ s and $\tau = 2P$. The residual can be made arbitrarily small for larger scattering values, but will require ever larger N values.

Gaussian noise is added to simulate an observation with a chosen peak signal to noise ratio (SNR). Starting from an un-scattered pulse with an off-pulse baseline of zero, all methods described above will, in the case of high levels of scattering (large τ compared to the period P), result in profiles with a raised off-pulse baseline. We use a Gaussian kernel to smooth the noisy scattered profiles. From the smoothed profile we determine the raised baseline value to subtract. The impact of this on the fitting will be discussed in Sec. 3.1.

The frequency range over which we simulate data is chosen to create severe scattering imprints at the lowest considered frequency, where τ is a substantial fraction or even larger than the pulse period. In these cases special care must be taken in the forward fitting process, as discussed the next section. Although not presented in this paper, the simulations can be extended to include scattering screens with Kolmogorov (turbulent) scattering statistics.

2.3 Fitting techniques

Forward fitting techniques create models to fit the data by convolving an estimated intrinsic pulse shape with a parameterised broadening kernel that describes the response of the ISM (Ramachandran et al. 1997). In general, it is difficult to obtain a model for the intrinsic pulse shape as pulsars often exhibit frequency evolution. In this paper we approximate the intrinsic pulse shape with a Gaussian function, with the caveats already stated.

An example of a forward fitting method is an Exponentially Modified Gaussian (EMG) method. In its most basic form this gives the analytic expression for a single Gaussian pulse convolved with a one-sided exponential function. Such

a simplistic implementation would not work well for large scattering times compared to the pulse period.

As mentioned earlier, inverse techniques perform a deconvolution of the broadened pulse shape through e.g the *CLEAN* algorithm (Bhat et al. 2003). We focus on forward techniques. As shown in eq. (2), the spectrum of a Gaussian isotropic scattering model has the form $\tau \propto \nu^{-4}$. We test the efficiency of fitting methods by determining how accurately they replicate this theoretical spectrum.

We use the *lmfit* package in python for our fits (Newville et al. 2014). We use our methods of generating noiseless scattered profiles, described above, as the fitting functions. The method fitting parameters are the underlying Gaussian intrinsic pulse components, i.e. the width (σ), mean (μ) and amplitude (A), as well as the characteristic scattering time (τ) and a DC offset, described in the next section.

The aim is to fit the simulated data for a range of frequencies to thereby obtain a spectrum of the characteristic scattering time, τ . In doing so we bear in mind that all observed scattered broadened profiles will typically result from integrating over some frequency band. The τ value obtained for fitting a monochromatic simulated scattered profile at frequency f_m corresponds to the τ value associated with the central observing frequency f_c and a bandwidth δf by

$$f_c = 10^{[\log_{10}(f_m + \delta f/2) + \log_{10}(f_m - \delta f/2)]/2} \quad (5)$$

or inversely,

$$f_m = \frac{1}{2} \sqrt{\delta f^2 + 4f_c^2}. \quad (6)$$

This equation provides the relevant frequency for a τ measurement given any f_c and δf , assuming a power law spectrum. Therefore computed τ values should be plotted against calculated f_m values to obtain meaningful fits. We will revisit the impact of this in the next section. We note that in our methodology we fit all frequency bands independently, whereas more constraints could be placed by simultaneously fitting for τ and the pulse parameters across a broad bandwidth, if that were available.

3 RESULTS

3.1 A simple end-to-end experiment

Before moving onto detailed results, we describe a simple end-to-end experiment. We simulate an intrinsic pulse shape scattered by an isotropic infinite scattering screen as described in Sec. 2.2. The frequency range over which we conduct the simulations is 60 - 140 MHz. Within this frequency range the pulse profiles of the pulsar with $P = 1.0$ s and a duty cycle of 2.5 %, are notably scattered and in the severest of cases (lowest frequencies) the scattering tail wraps around.

Initially, we generate profiles monochromatically at 60 and 90 MHz with a peak SNR of 20. We fit the noisy profiles using the train method. We find that the determination of the off-pulse baseline of the noisy data has an important effect on the fit, which we investigate as follows. As mentioned, we use a Gaussian kernel to smooth the simulated data and from this smoothing function estimate the off-pulse baseline, denoted in the following as b . The effect of the error in accurately determining the baseline, on the

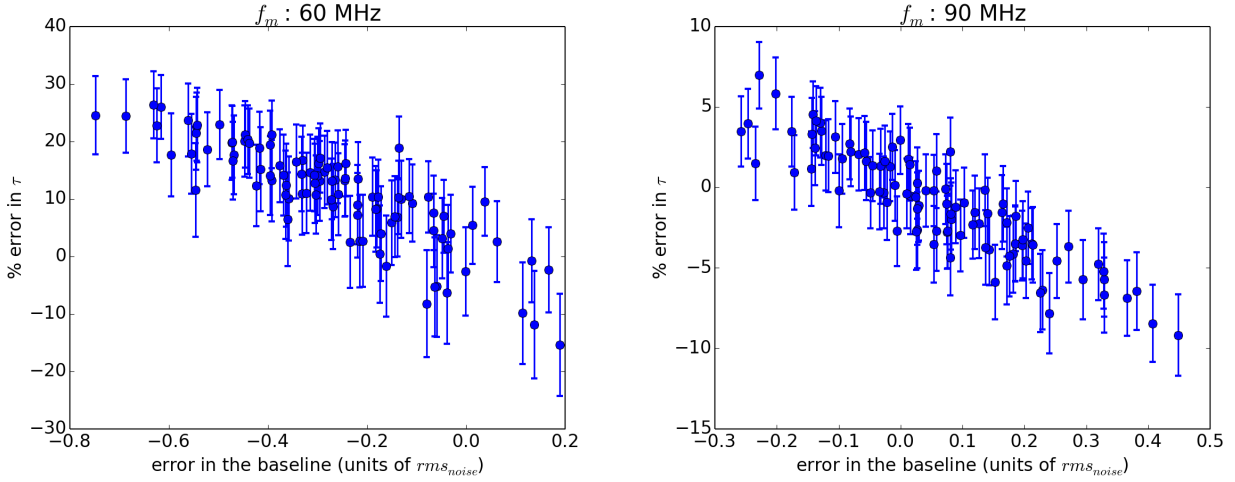


Figure 4. Uncertainties in the obtained τ values as a function of the error in estimating the baseline of the scattered profile, shown for two distinct frequency values. Each panel represents 100 independent realisations of Gaussian noise. The data are simulated as described in the text.

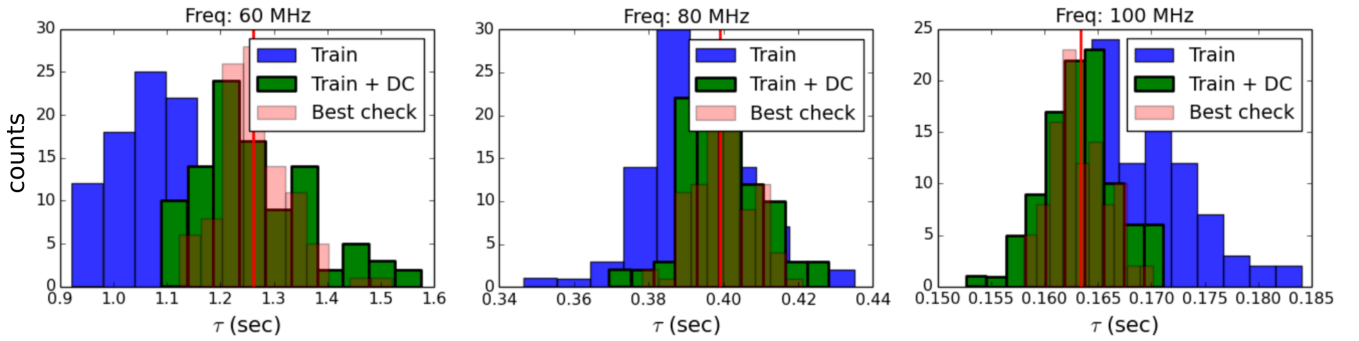


Figure 5. Histograms showing the obtained spreads in τ values at three sample frequencies, using different methods. The red line indicates the true value of τ . The *best check* histogram shows the distribution obtained from fits where the noiseless minimum of the profile is known exactly, and is therefore an idealised reference case. The *train* and *train + DC* methods show the spreads in τ using the train method with our without fitting for a DC-offset. In the case where the method does not fit for a DC-offset, a Gaussian smoothing kernel is used to estimate and subtract the raised off-pulse baselines.

subsequently fitted τ value, is shown in Fig. 4. Here, we compare the fitted τ values to the outcome when the off-pulse baseline is known exactly. The plots show a percentage error in τ , $100(\tau_{\text{known}} - \tau_{\text{fit}})/\tau_{\text{known}}$ plotted against the error in the determined baseline in units of the noise, $(b_{\text{known}} - b_{\text{est}})/\text{rms}_{\text{noise}}$, for 100 realisations of the noise at 60 and 90 MHz.

The error in the τ value increases systematically as the error in the baseline value grows, and is greater at the lower of the two frequencies. At 60 MHz, an error of 10 - 20% in the τ value is already possible for baseline errors below 1σ of the noise. As the frequency increases, the error in τ for a given error in the baseline drops significantly. We also note that at lower frequencies the τ value is more likely to be underestimated, whereas at higher frequencies τ is more likely to be overestimated.

This result prompted us to investigate an improved model. In Fig. 5 we show the impact of adding a DC-offset to the parameter set for which the model fits. In the *train + DC* model we therefore no longer estimate the off-pulse baseline

in order to subtract it from the profile before fitting the data, but we allow the model to fit for a best off-pulse baseline value. The spread in τ values in Fig. 5 represents 100 noise realisations, and the data are simulated as for Fig. 4. Introducing a DC offset to the train method, ensures that the estimated τ values are centred on the true value (red line). The histogram labelled *best check* gives the spread in obtained τ values when the noiseless off-pulse baseline of the simulated data is known exactly. We note that the *train + DC* model approximates this idealised distribution well. However, at both the low and high end of the frequency range shown, the *train + DC* distributions are more asymmetric than this reference case.

Fig. 6 shows the correlation between *DC* and τ and the associated confidence intervals, at 60 MHz. The parameters are anti-correlated, with a correlation coefficient of $C(\text{DC}, \tau) = -0.73$. This correlation allows an improved accuracy of the obtained τ value, when including *DC* as a fitting parameter.

For the remainder of this paper we will mostly make

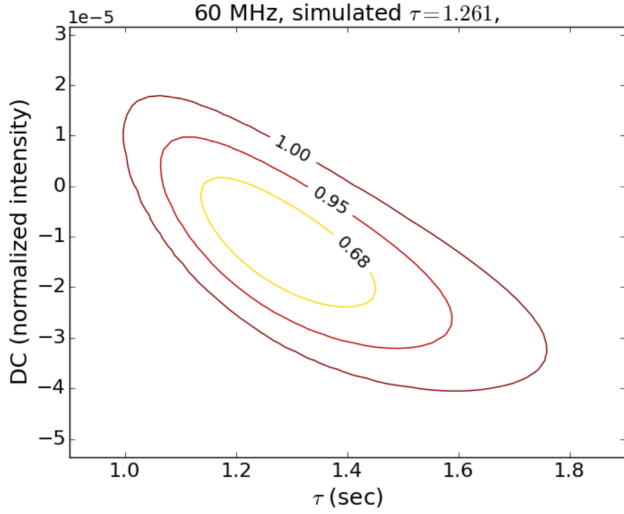


Figure 6. The confidence intervals ($1\sigma - 3\sigma$) associated with the obtained fits for the vertical offset parameter, DC , and characteristic scattering time, τ .

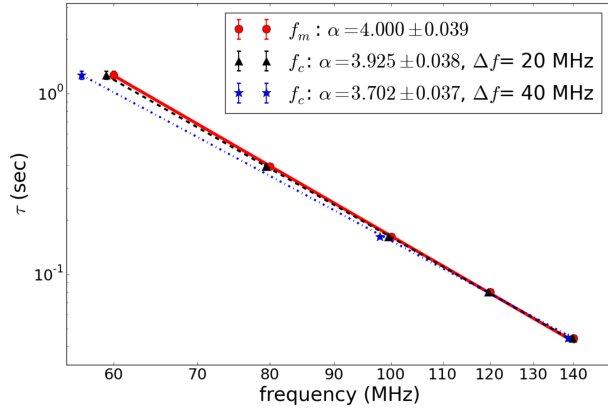


Figure 7. Deviations from the theoretical characteristic τ spectrum are observed as the bandwidth of the simulated data increases. The obtained τ associated with a central frequency f_c and bandwidth δf relate to the simulated result produced by a monochromatic frequency input f_m through eq. (5).

use of the train method with a DC offset. The obtained τ values are therefore subject to uncertainties as inferred from the distributions in Fig. 5. We consider this to best match real observational data, where the baseline determination will always be affected by instrumental noise.

In this end-to-end experiment we also investigate the typical impact of the bandwidth of an observation on the determined τ spectrum. As shown in eq. (5) the τ value associated with a fit of an integrated profile must be remapped to the corresponding frequency f_m , before the spectrum is computed.

Fig. 7 shows the impact on the τ spectrum fit when the central frequency values f_c are used instead of f_m . Simulating the data monochromatically at 20 MHz intervals from 60 to 140 MHz leads to a measured spectral index of $\alpha = 4.00 \pm 0.04$, reproducing the theoretical spectrum well. In contrast a bandwidth of 20 MHz at 20 MHz in-

Table 1. The mean spectral indices and standard deviations as obtained from 1000 τ and subsequent spectra fits. Two pulsars are investigated: a slow pulsar and a millisecond pulsar. For each pulsar the spectral index is determined over two sets of frequencies. The fits and their associated standard deviations improve with a decrease (increase) in scattering (frequency). The different methods are as described in the text.

Slow pulsar: 1.0 sec, 2.5% duty cycle				
SNR: 20				
Frequency Range: 60 – 120 MHz				
	EMG	Train	Train +DC	Train +DC asymm errors
Mean Spectral Index	2.90	3.80	4.00	4.01
Standard deviation	0.05	0.09	0.09	0.09
% Error	-27.6%	-4.9%	-0.1%	0.2%
Frequency Range: 90 – 150 MHz				
Mean Spectral Index	3.78	3.89	4.00	4.00
Standard deviation	0.06	0.06	0.06	0.03
% Error	-5.4%	-2.8%	0.0%	0.0%
Fast pulsar: 20 ms, 10% duty cycle				
SNR: 20				
Frequency Range: 150 – 210 MHz				
Mean Spectral Index	1.11	3.71	3.87	4.01
Standard deviation	0.06	0.42	0.32	0.17
% Error	-72.2%	-7.2%	-3.1%	0.2%
Frequency Range: 200 – 260 MHz				
Mean Spectral Index	2.38	3.85	3.98	3.99
Standard deviation	0.09	0.18	0.14	0.07
% Error	-40.4%	-3.7%	-0.5%	-0.3%

tervals leads to a spectral index of $\alpha = 3.93 \pm 0.04$, when plotted against f_c . An even larger bandwidth of 40 MHz gives $\alpha = 3.70 \pm 0.04$. To produce this figure, data were fitted with the train method and the baseline of the simulated data subtracted before adding noise - to eliminate the effects of baseline errors as mentioned above.

3.2 Results for Infinite Screens

We now move on to infinite scattering screen setups. We first present the obtained τ spectra, followed by the flux spectra. In each case we consider the two scattering mechanisms discussed in Sec. 2.2.2, which serve as examples of either an isotropic or anisotropic scattering process.

3.2.1 Tau Spectra

3.2.1.1 Isotropic Scattering Tau-spectra

Isotropically scatter broadened profiles of the two simulated pulsars described in Sec. 2.2.1 are shown in Figs. 8 (1.0 s pulsar) and 10 (20 ms pulsar). The simulated data, as seen in yellow, again have peak SNR = 20. The scattering setup is as described in Sec. 2.2.2. A range of observing frequencies is chosen such that the scattering tail wraps around at the low end and a typical exponential scattering tail is observed at the high end.

The profiles are fitted with the train + DC method as described in the previous section. Fits without the DC-offset parameter (i.e the train method in which the off-pulse

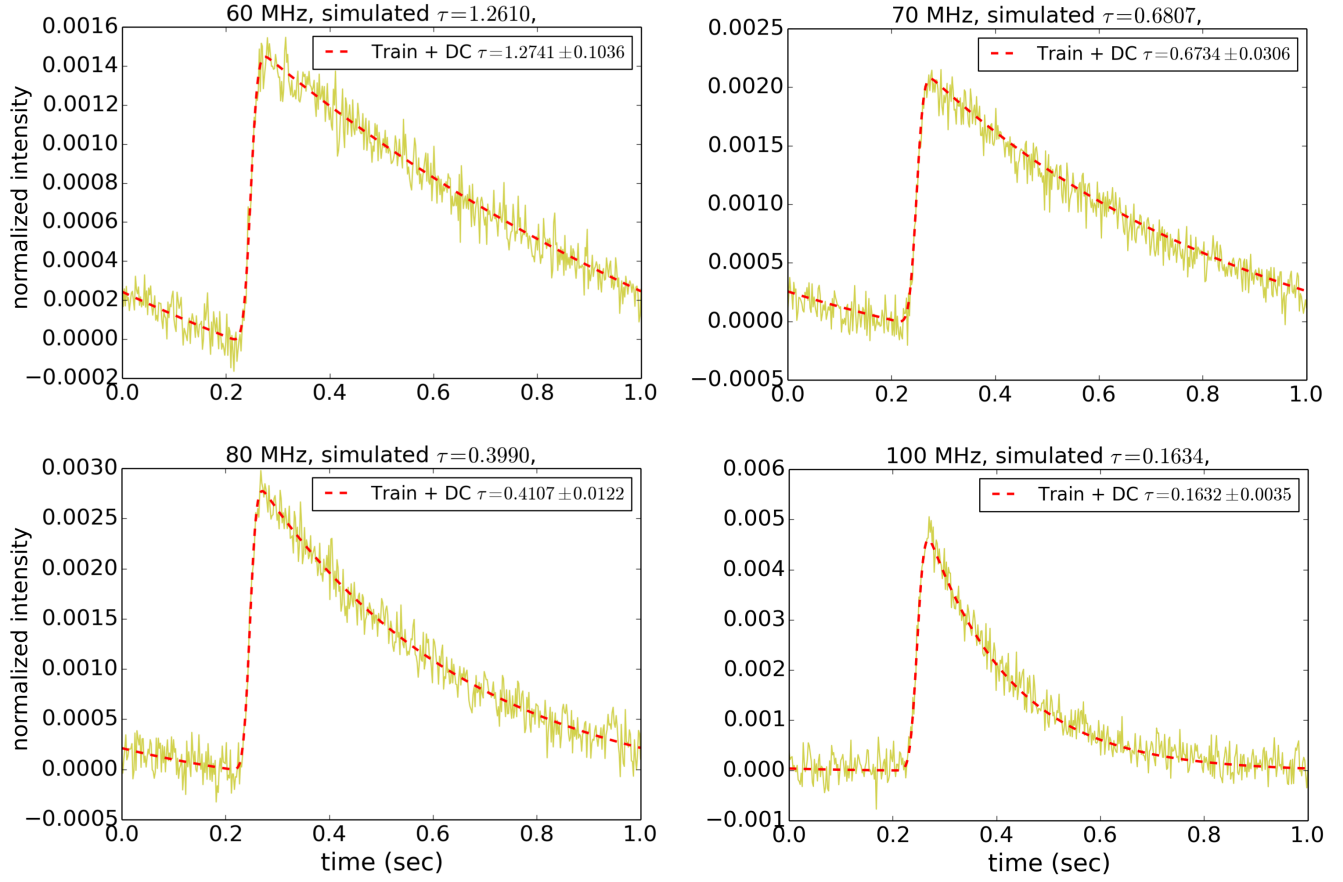


Figure 8. Best-fits as produced by the train + DC method for simulated scattered pulse profiles of a $P = 1.0$ s pulsar. Fits without the DC-offset parameter (i.e. the train method in which the off-pulse baseline is determined using a smoothing kernel) look similar to the DC fits by eye, but produces skewed values for τ , as was shown in Fig. 5. The impacts of these offsets on the power law fit to the τ spectra can be seen in Table 1. Profiles are plotted normalised to the area under the pulse - such that at the lowest frequency (viz. 60MHz) in the absence of a non-zero off-pulse baseline, the area under the pulse is equal to one.

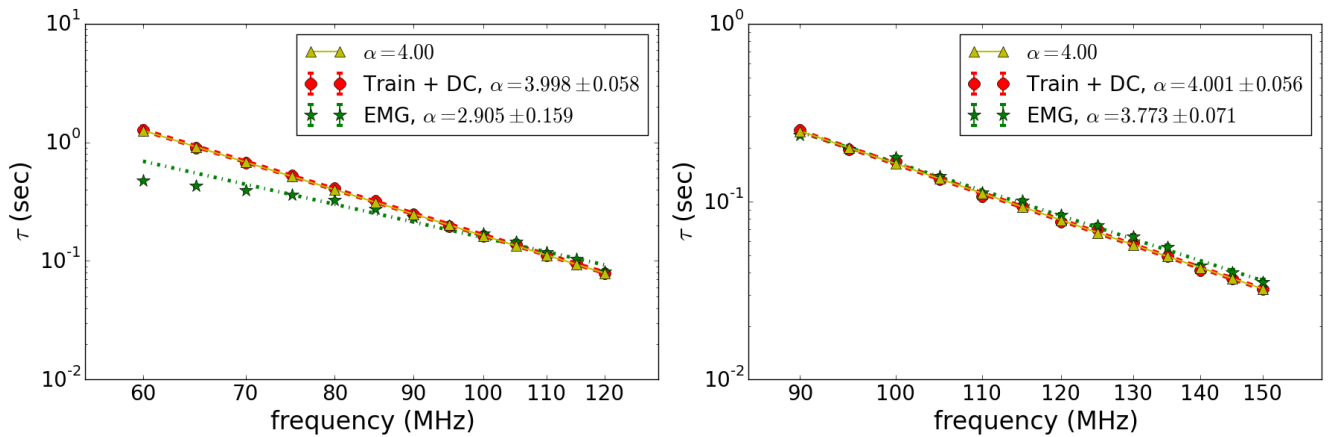


Figure 9. The τ spectra for the 1.0 s pulsar as obtained by the train + DC method. The left-hand panel shows the spectrum at lower frequency values (high scattering), and the right-hand panel at higher frequency values. Even at the levels of scattering associated with the left hand panel, the train + DC method is able to accurately reproduce the theoretical τ spectrum with a spectral index close to 4. The obtained 1σ error bars are of the same order as the marker size and therefore not clearly visible. The EMG method is included as an example of a method that does not model wrap around scattering tails and can therefore not accurately reproduce spectral index values in data where high levels of scattering are present. Table 1 gives additional information about the standard deviations in the α values when rerunning the experiment multiple times and fitting with various methods.

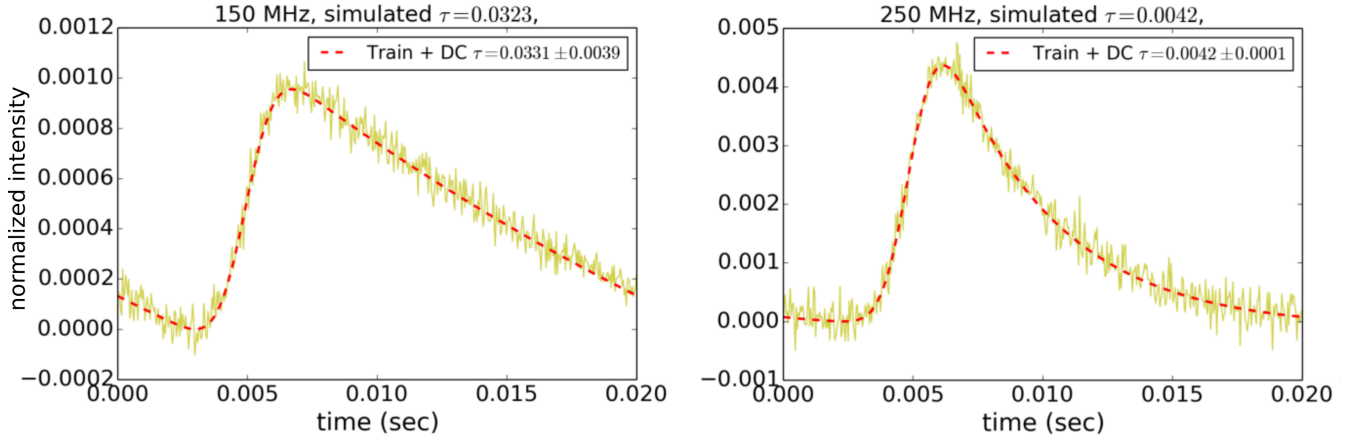


Figure 10. Similar to Fig. 8, but for a pulsar with period $P = 20$ ms and a duty cycle of 10%.

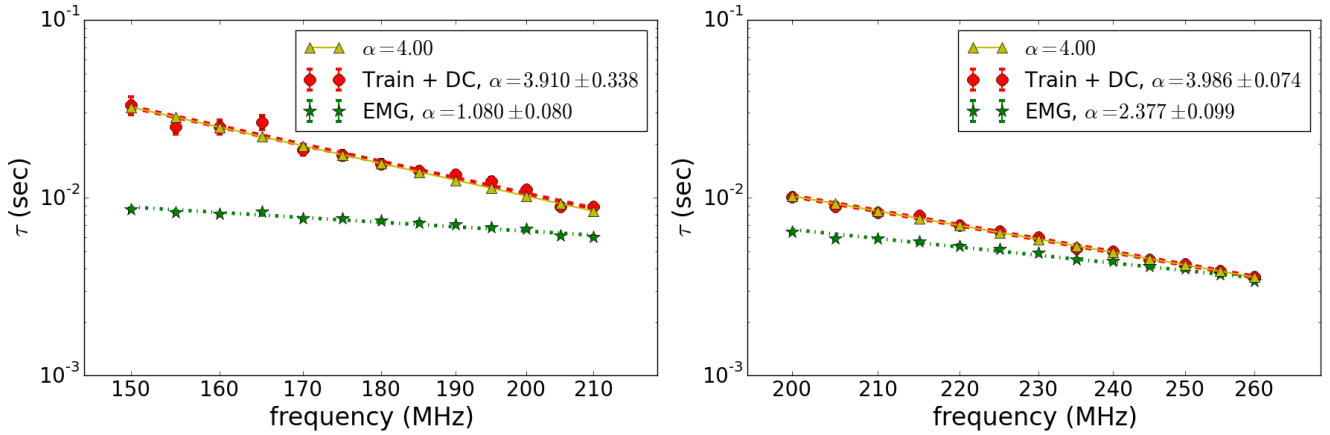


Figure 11. The τ spectra of the modelled millisecond pulsar. The uncertainty in the spectral index obtained by the train + DC method grows as the frequency range is lowered. At these lower frequencies the EMG method deviates most severely from the theoretical spectrum, and continues to underestimate the value of α even for the higher frequency range. The frequency at which the choice of method matters is much higher for the simulated millisecond pulsar than for the slow pulsar in Fig. 9.

baseline is determined using a smoothing kernel) look very similar to the DC fits by eye, but produce skewed values for τ as was shown in Fig. 5. The impacts of these offsets on the power law fit to the τ spectra can be seen in Table 1.

As described in Sec. 2.2.3 the train and the long-train methods, with added DC offsets, can perform equally well for all the simulations. A long train method with a too short pulse train will lead to an underestimation of τ .

In the case of severely scattered profiles it is vital to use an accurate fitting method as the outcome becomes increasingly sensitive to the model parameters. The train + DC method presents the simplest (and computationally fastest) way to obtain accurate τ values.

The τ spectra for the simulated pulsars are shown over two frequency ranges in Figs. 9 and 11. As the frequency increases, the obtained spectral index value more accurately approaches the simulated index of $\alpha = 4$. Here we also include a basic implementation of the EMG model not accounting for the pulse period P . This method underestimates τ significantly at high levels of scattering. At higher

frequencies where the scattering tail no longer wraps around, the choice of method becomes less important.

We generate 1000 datasets, of simulated scattered profiles at intervals of 5 MHz within our frequency ranges, per pulsar and run the fits for τ and α . The outcomes of this process are summarised in Table 1.

For each realisation of the experiment, the train + DC method performs best, and in the case of the millisecond pulsar (for which the scattering is more severe overall) has a lower standard deviation in its estimation of α , than the standard train method. The average spectral index is calculated over two sets of frequencies for both the slow and millisecond pulsar. At higher frequencies (lower scattering) α approaches the simulated value of 4 more accurately. This can for example be seen in the tabulated error values, which represent the error of the mean spectral index with respect to the true value of 4 expressed as a percentage.

In Fig. 5 we showed the spread of τ values obtained by the train methods and by a *best check* reference case. As noted even the train + DC model shows a skewed distribution of τ - especially at low frequencies. To take this skewness

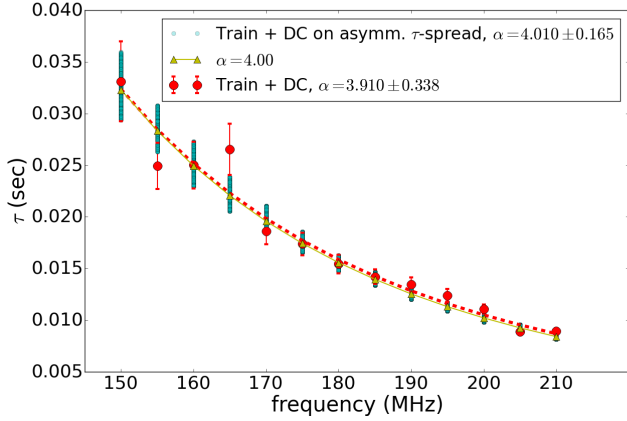


Figure 12. The power law fit of the millisecond pulsar is improved when we take the skewness of the τ distribution (resulting from 1000 noise realisations and subsequent fits) into account at each frequency. The τ values that fall within the 1σ uncertainty of this distribution are shown as cyan circles at each frequency. The train + DC fit (red) is an example of a single noise realisation and subsequent fit. Data are plotted on a linear scale for the τ spread values to be visible.

Table 2. Spectral indices similarly obtained as for Table 1. The simulated data have SNR values of 10 and 50, and show the expected improvements in fits as the SNR values increases. The tabulated values represent a mean α and associated standard deviation after 1000 executions of the experiment.

Slow pulsar: 1.0 sec, 2.5% duty cycle		
Frequency Range: 60 – 120 MHz		
	Train + DC	
	SNR 10	SNR 50
Mean Spectral Index	4.0	4.0
Standard deviation	0.18	0.04
% Error	-0.5%	-0.1%

into account, we update the power law fit as follows. From a distribution of τ values (obtained from several noise realisations at a given SNR level, and subsequent profile fits with the train + DC method), we subtract 16% of the values from both the lower and upper end of the distribution to be left with the 1σ (68%) central distribution. We now use the values within this range along with their corresponding frequency values to update the τ spectrum, as shown in Fig. 12. As can be seen from Table 1 under the column heading *Train + DC, asymm errors*, this process mainly provides a small improvement at extreme levels of scattering, such as for the millisecond pulsar at 150–210 MHz. For the most part of this article we will therefore not show these improvements.

As is expected the success of the τ spectra fits and the associated uncertainties are also impacted by the peak SNR value, and will improve as the SNR increases. In Table 2 we show the achieved spectral index values after repeating the experiment above at SNR values of 10 and 50. At a SNR value of 50 the theoretical spectrum is best reproduced and the standard deviation in the spread of individual measurements has decreased significantly.

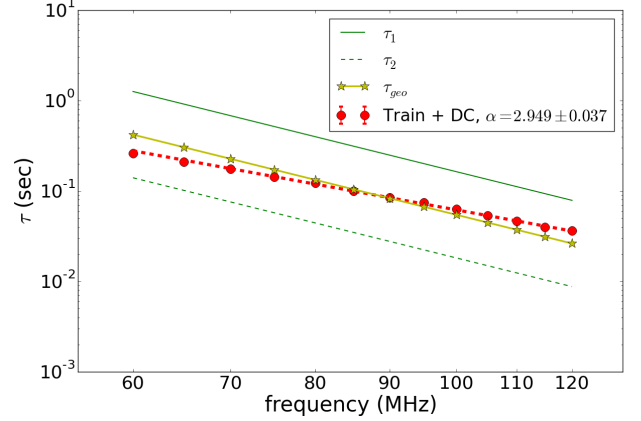


Figure 14. The τ spectrum for an anisotropic scattering event fitted with a single τ power law is shown in red (thick dashed line, circles). The green lines (solid and dotted without markers) show the simulated values for τ_1 and τ_2 . The geometric mean for each of these values, $\tau_{geo} = \sqrt{\tau_1 \tau_2}$ is shown in yellow (stars). The τ_1 , τ_2 and τ_{geo} values have a modelled frequency dependence $\tau \propto \nu^{-4}$. At low frequencies the fits deviate most significantly from the τ_{geo} values.

3.2.1.2 Anisotropic Scattering Tau-spectra

In the case of the anisotropic scatterer discussed here, the broadened profile is dependent on two characteristic scattering times, each with a frequency dependence $\tau_{x,y} \propto \nu^{-4}$, as described in Sec. 2.2.2. In general a more complex frequency dependence could be at play, since different observing frequencies will interact with anisotropic scattering ellipses of varying sizes. Sampling different regions of the scattering screen in this way, can lead to additional frequency dependencies not included in our model.

We show how an anisotropic scattering process will create apparent deviations from the expected τ spectrum when fitted with an isotropic scattering model. The scattering strengths are chosen to be $\sigma_{ax} = 3$ mas and $\sigma_{ay} = 1$ mas at 1 GHz. That is to say the screen scatters 3 times more weakly in the chosen y -dimension than in the x -dimension. Setting $\sigma_{ay} = 3$ mas will reproduce the isotropic scatterer considered in Sec. 3.2.1.1. The scattering geometries are as before and the pulsar is the same 1 sec Gaussian profile from the previous section.

The simulated data are fitted with the train + DC method. The resultant profile fits for a selection of frequencies are shown in Fig. 13. The corresponding τ spectrum is shown in Fig. 14. As before these are chosen to represent the mean profiles and spectrum after repeating the experiment with Gaussian noise of peak SNR = 20 multiple times. The mean and standard deviation of the spectral index as calculated after 1000 repetitions are given in Table 3.

In Fig. 14 the input values of τ_1 and τ_2 (green solid and dashed lines) are shown as a reference. Both have a frequency dependence with spectral index $\alpha = 4$. The geometric mean,

$$\tau_{geo} = \sqrt{\tau_x \tau_y} \quad (7)$$

which would reduce the anisotropic scattering to isotropic

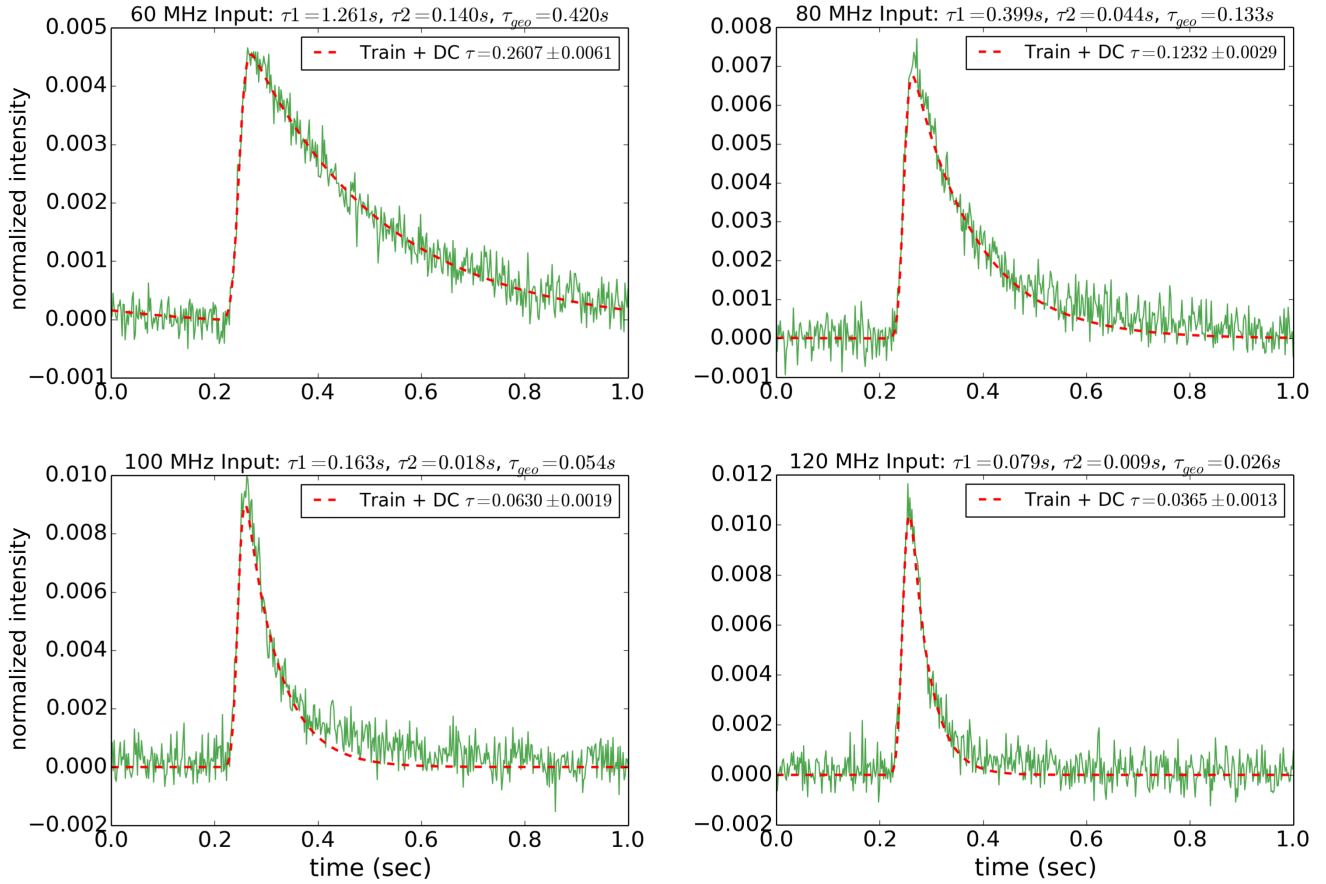


Figure 13. Profile fits as obtained by applying the isotropic train + DC method to anisotropically simulated pulse profiles of a $P = 1.0$ s pulsar. The anisotropy is created by choosing a thin screen that scatters 3 times more strongly in one dimension than in the other. At higher frequencies the isotropic fit to anisotropic data leads to a rough estimation of the geometric mean scattering time (τ_{geo}). At low frequencies τ_{geo} is significantly underestimated. Profiles are plotted normalised to the area under the pulse, such that at the lowest frequency (viz. 60MHz) in the absence of a non-zero off-pulse baseline, the area under the pulse is equal to one.

Table 3. The mean spectral indices as obtained by fits to profiles produced through an anisotropic scattering process. The tabled spectral indices represent the average index as obtained from 1000 fitting procedures.

Anisotropic Scattering

Slow pulsar: 1.0 sec, 2.5% duty cycle SNR: 20	
Frequency Range: 60 – 120 MHz	
	Train +DC
Mean Spectral Index	2.95
Standard deviation	0.05
% Error	-26.3%

scattering for $\tau_1 = \tau_2$, is computed from the input τ_1 and τ_2 and shown in yellow (stars). The single τ fit, shown in red, provides an estimate of the τ_{geo} value at mid frequencies. At low frequencies however the method fits deviate from τ_{geo} significantly. Such an estimate of τ_{geo} will worsen as the degree of anisotropy increases.

These results show that scatter broadened profiles caused by anisotropic scattering mechanisms require more sophisticated fitting models that can account for the de-

gree of anisotropy. Such a model would make explicit use of eq. (4) in its description. For individual pulsars, screen anisotropy might be revealed through elongated images (Briskin et al. 2010) or features of the secondary spectra, which will be instructive in determining the choice of fitting function, and in improving the fit.

3.2.2 Flux-spectra

The mean flux (S_m) of a pulse profile is calculated by integrating the intensity of the profile averaged over the pulse period. This is illustrated schematically on the right-hand side of Fig. 2, where the mean observable flux is shaded in blue. Scatter broadening that causes the power of pulses to be smeared into succeeding pulses will lower the observable flux, as intrinsic flux is lost to the raised baseline. Since this is an artefact of the temporal domain, we can in principle circumvent the flux loss by imaging pulsars directly (e.g. Dembska et al. 2015).

We compute the flux for the setups in Sec. 3.2.1 using the 1 sec pulsar, and show that a turnover in the flux spectrum occurs at low frequencies (high scattering). The

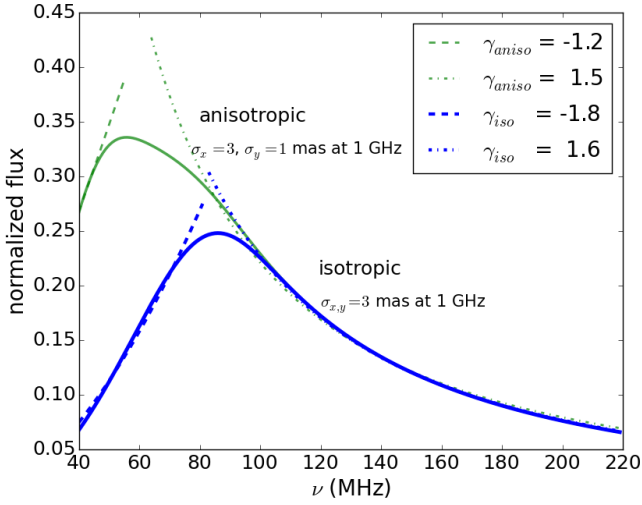


Figure 15. Spectral turnovers in the pulsar flux spectra associated with infinite isotropic and anisotropic scattering screens. The turnovers at low frequencies are due to the raise in off-pulse baselines which leads to a loss in observable flux. These effects cause positive spectral slopes over the associated low frequency ranges. The spectrum related to anisotropic scattering here has a turnover at a lower frequency than the isotropic case. This is because the mean scattering strength in the chosen anisotropic case is weaker.

location of this turnover is dependent on scattering mechanism and geometry.

The mean flux values quoted here are calculated from the noiseless scatter broadened pulse shapes. It is normalised such that in the absence of the baseline being raised the lowest considered frequency (viz. 40 MHz) will have a mean flux of unity.

Fig. 15 show the turnovers in spectra for both the isotropic and anisotropic scattering setup. The turnover in the flux spectrum for the anisotropic scattering case occurs at a lower frequency and causes a smaller loss in flux since the mean scattering is weaker. The turnover also has an asymmetric shape.

Fits to the high-frequency end of these spectra approximate the intrinsic spectral index of $\gamma = 1.6$. For low frequencies however the spectral turn-around leads to a positive slope such that the spectral index is estimated at $\gamma = -1.8$ for the isotropic case and $\gamma = -1.2$ for the anisotropic case.

In Fig. 16 we illustrate how changing the scattering geometry impacts on the flux spectra. We show that as the position of the screen (D_s) is moved for a fixed overall distance (D) the frequency at which we expect to observe spectral turnover due to baseline effects will change. This phenomenon is symmetric about the mid-point ($D_s/D = 0.5$) such that for example the spectra for $D_s/D = 0.8$ and 0.2 are equal. The observed scattering effects are at its maximum for $D_s/D = 0.5$ and therefore the frequency at which flux loss is first observed will be highest for a midway screen.

The overall distance D is increased from 3 kpc to 6 kpc from the top panel to the bottom panel of Fig. 16. This overall increase in distance increases the scattering and subsequently the frequency at which spectral turnover appears. Simultaneously as D is increased the frequency window

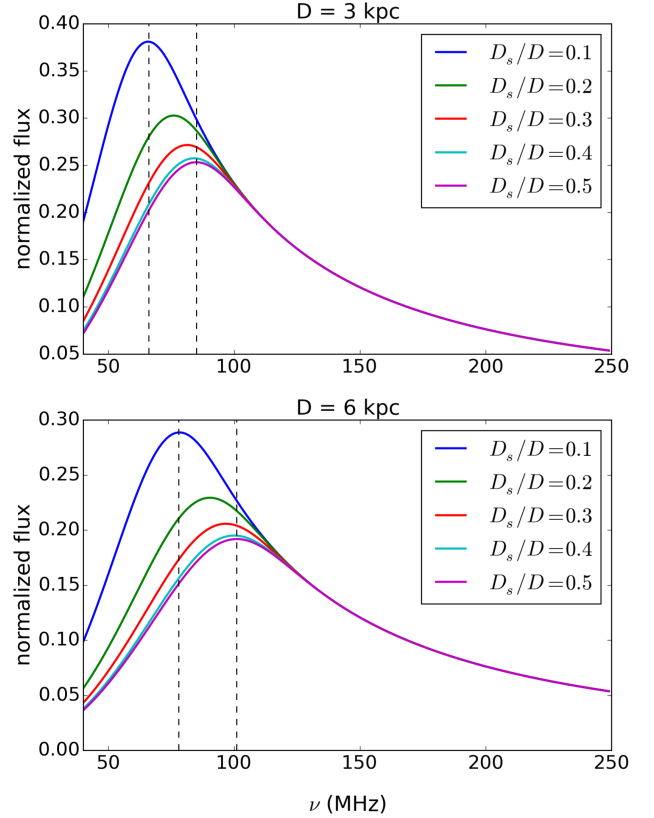


Figure 16. Flux spectra plotted for different ratios of D_s/D and overall distance D . The highest frequency at which a spectral turnover occurs increases from 86 MHz to 102 MHz as D is increased from 3 to 6 kpc. Ever higher values of D does not increase this value significantly ($D = 12$ kpc has a turnover at ~ 120 MHz).

within which spectral turnover takes place for all plotted ratios of D_s/D (indicated with dashed lines) grows slowly. This result shows that for distances within our galaxy and a scattering strength of $\sigma_\theta = 3$ mas at 1 GHz, as taken from Cordes & Lazio (2001), turnovers in mean flux spectra will occur below 100 MHz. Spectral turnovers at higher frequencies are either due to different scattering mechanisms (deviations from isotropically Gaussian scattering) or different scattering strengths than the typical value used here. However since a wide range of scattering strengths are expected, it is unclear whether or not a coherent picture is expected throughout the population of pulsars with measured spectra.

At very large distances, $D > 200$ kpc, (while still using the above scattering strength) the simulated pulse is completely scattered out and a turnover in the flux spectrum no longer detectable in this way.

3.3 Effects of screen geometry

In Sec. 2.2.2 we discussed computing the temporal broadening functions (f_{ts}) of truncated screens by means of a ray-tracing code. We expect these finite screens to cause a loss in the observable flux of the pulsar. Furthermore we investigate how the shape of an average pulse profile, and

consequently the τ spectrum, is impacted by the presence of such a screen.

By defining the screen in terms of its height and width we can create both square screens that will have a comparable impact to a circular screen or rectangular (elongated) screens. In the theoretical limit where screens become infinitely more long than wide (i.e. 1D filaments) the broadening function will tend to,

$$f_t = e^{-t/\tau} / (\sqrt{\pi t \tau}) U(t). \quad (8)$$

We stick to screens where the ratio of its dimensions has not yet reached this theoretical limit.

3.3.1 Profile effects

The broadening function for a source at a distance of $D = 3$ kpc with a finite isotropic scattering screen halfway between the source and the observer is presented in Fig. 17. The modelled screen has size 400 AU by 600 AU, with the pulsar offset to the middle of the screen by 100 AU horizontally and 50 AU vertically. The scattering strength is as before, $\sigma_a = 3$ mas at 1GHz.

The dips in the broadening function are associated with the four edges of the screen and show the associated loss in flux at different frequencies. In this instance we see dips at 0.02, 0.10, 0.15 and 0.20 s associated with distances to the edge of 100, 250, 300 and 350 AU.

The numerical broadening functions are convolved with the same $P = 1.0$ s pulsar as used for the infinite scattering screens in Sec. 3.2. A selection of the resulting broadened profiles are shown on the left-hand side of Fig. 18.

The broadening effects can be split into three regimes. At high frequencies (above 210 MHz, not shown) the effect of scattering is negligible and the observed pulse resembles the intrinsic Gaussian pulse. At mid-frequencies (here 120 – 210 MHz) the pulse shapes have the typical exponential tail as would be associated with an infinite scattering screen. At low frequencies (below 120 MHz) the finite nature of the scatterer becomes visible in the form of bumps associated with the loss in flux at various path lengths (or times). These pulse shapes have a much steeper trailing edge than normally associated with a scatter broadened pulse shape.

As an example of these effects on a more complicated profile we use a high frequency profile of B1237+25 as a multi-component template and consider how this template would change with frequency when it is scattered by a truncated screen. In our simulations we use the pulsar's period ($P = 1.38$ s) and distance ($D = 0.85$ kpc), while acknowledging that the pulsar itself, being a high galactic latitude source, is not an ideal candidate for scattering studies. We consider the impact of an isotropic screen, size 200 by 200 AU, placed midway along the line of sight. The scattering strength is as before. The resulting broadened pulse shapes at several frequencies are shown on the right-hand side of Fig. 18.

Note how the truncation effect of the screen can cause the height of the two profile peaks to change relative to each other. Between 110 and 130 MHz the left peak is lower than the right one, at 140 MHz the peaks are roughly equal, and beyond 140 MHz the left peak is the higher peak. The change

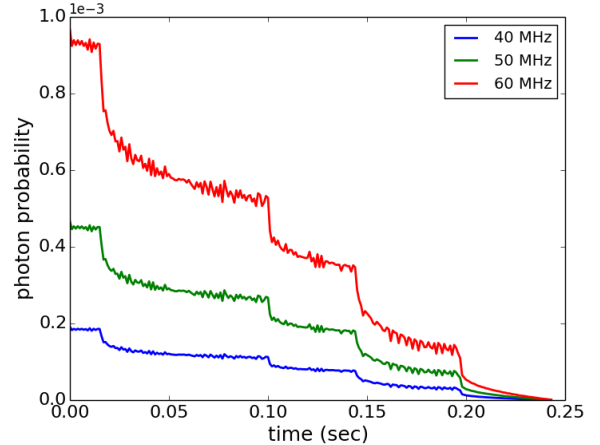


Figure 17. The broadening function associated with a truncated scattering screen of size 400 AU by 600 AU, with the pulsar offset to the centre of the screen as discussed in the text. The functions are generated numerically by means of a ray tracing experiment that calculates the probability of photons arriving at the observer at a given time.

in features in these profiles could easily be misinterpreted as features of the intrinsic pulse shape.

The above scattering setups are theoretical constructs, however by investigating the impact of truncated screens on both the τ and flux spectra, we can build a more complete picture of all the possible sources of deviations to the otherwise theoretically predicted spectral trends.

3.3.2 Effects on τ spectra

To estimate the τ spectra we added Gaussian noise to the simulated profiles, after which we fitted them with the train method of Sec. 2.3. (Since the truncation prevents wrap around profiles here we do not have to include a DC offset parameter.) The SNR value is modelled to scale with the mean flux of the profile, and is therefore notably smaller at low frequencies for which the flux loss is greater. (We discuss the flux spectrum in more detail in the next section.) The resulting profile fits for a selection of frequencies are shown in Fig. 19.

The corresponding τ spectrum is shown in Fig. 20. From the τ spectrum it is clear that a model based on the assumption of an infinite screen, greatly underestimates the true value of τ . A weighted fit to all the obtained τ -values produces a spectral index of $\alpha = 3.17$. If we identify a single frequency at which the observations appear to become sensitive to the edges of the screen, we can fit the spectrum with two power laws instead. Using 100 MHz as this breakpoint, we find that $\alpha = 2.85$ above this frequency. Below 100 MHz we see a flattening of the τ spectrum, such that the obtained spectral index here is $\alpha = 0.34$.

3.3.3 Effects on flux-spectra

We expect to measure a loss in the observable flux for a pulsar behind a truncated scattering screen of such size that at low frequencies not all the photons from the pulsar will be

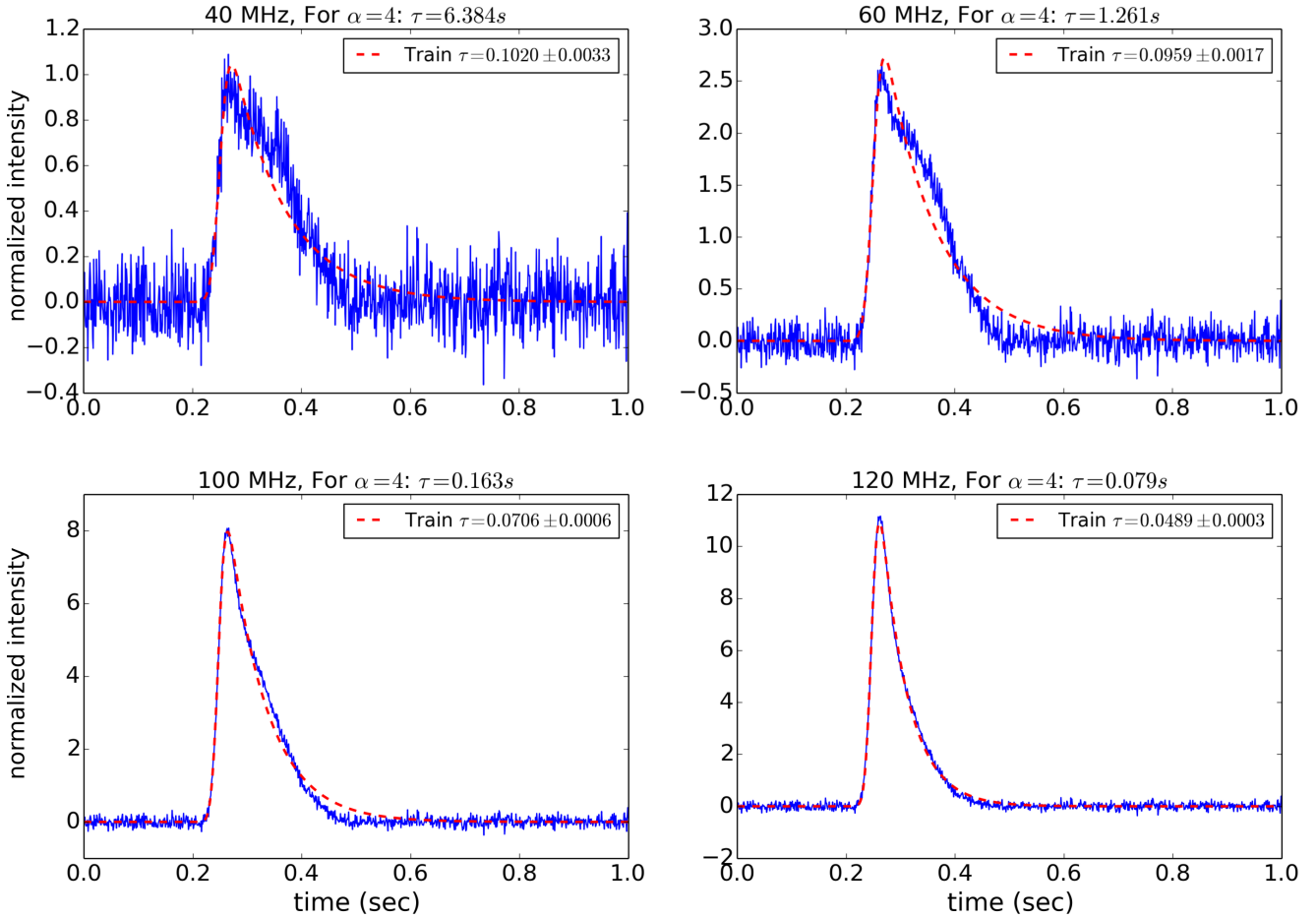


Figure 19. Profiles as generated by the truncated broadening function of Fig. 17 with added Gaussian noise. The profiles plotted here are normalised relative to the peak of the profile at 40 MHz (which has suffered the greatest flux loss). The input τ values at different frequencies are given, along with the values obtained from the fits with the train method. In all instances the τ values are underestimated by more than 45%.

refracted back to our line of sight. To model this loss we calculate the probability that a photon hitting the truncated screen at some angle and location on the screen, will reach the observer. We then compare this to the probability associated with a sufficiently infinite screen. *Sufficiently infinite* is defined in terms of the width of the Gaussian distribution from which the scattering angle is drawn. A distribution that runs from $-5\sigma_\theta$ to $5\sigma_\theta$ is considered to be an approximation of an infinite screen.

The ray-tracing code is normalised such that the temporal broadening function of a sufficiently infinite screen at the lowest studied frequency produces an impulse flux of unity. This means when convolving a simulated pulse with an infinite screen the resulting flux spectrum will equal the intrinsic flux spectrum of the pulsar with $S_m = 1$ at the lowest observed frequency (provided that no flux is lost due to high levels of scattering that have lead to wrap around scattering tails).

The flux spectrum in Fig. 21, associated with isotropic scattering, shows a turnover close to 120 MHz. Above this frequency the screen is effectively infinite such that the flux of the pulsar behind it is conserved, and the intrinsic pulsar spectrum with $\gamma = 1.6$ becomes visible. Below 120 MHz, the

spectral index is positive as more flux is lost towards lower frequencies (longer wavelengths) when the scattering process becomes sensitive to the edges of the scattering screen.

Fig. 21 also contains the flux spectrum for a truncated anisotropic scatterer of the same size and location. The anisotropy is created by again choosing that the screen scatters 3 times more weakly in one dimension. Since this leads to an altogether weaker scatterer the loss in flux is less severe and we observe the onset of flux loss at a lower frequency than for the isotropic example.

The turnovers in both Figs. 15 and 21 occur at frequencies below 150 MHz for the current scattering setup. Pushing the overall distance (D) to 6 kpc in Fig. 16 did not greatly increase the frequency at which the spectral turnover due to baseline effects occur. An alternative way to produce turnovers at higher frequencies, while keeping the scattering strength, could be by means of ever smaller scattering clouds. Fig. 21 includes an example (as a grey line) of a smaller screen with a higher turnover frequency.

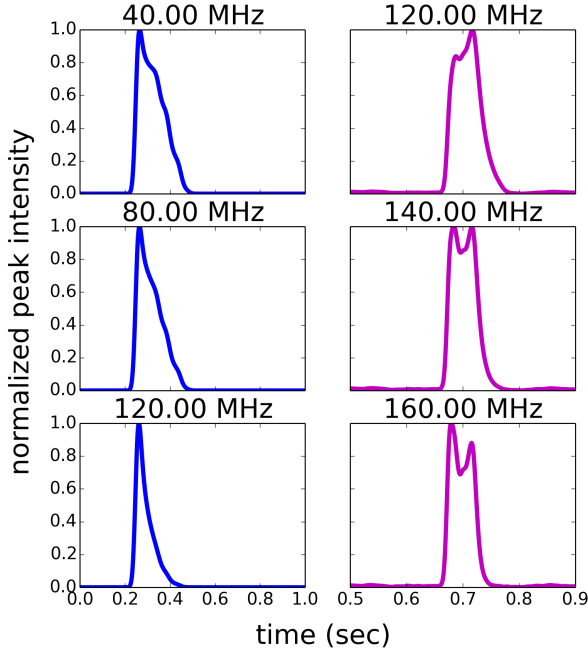


Figure 18. *Left:* Convolutions of an intrinsic Gaussian pulsar profile with the temporal broadening functions shown in Fig. 17. The scattering screen has size 400 AU by 600 AU, offset with respect to the line of sight as described in the text. The distances are $D = 3$ kpc and $D_s = D/2$ kpc. At 40 MHz and 80 MHz the effect of a truncated screen is seen as bumps and an overall steepening of the profile shape. At 120 MHz the profile starts to resemble the typical exponential scattering tail. *Right:* Broadened pulses using B1237+25 as a profile template and a screen of size 200 AU by 200 AU, centred on the line of sight, with $D = 0.85$ kpc and $D_s = D/2$. The relative heights between the two peaks change as the observing frequency changes. All profiles are plotted normalised to their height.

4 DISCUSSION

Reflecting on the results in the Sec. 3 we make the following statements.

In Sec. 3.1 we show that errors related to estimating the off-pulse baseline can be circumvented by treating the baseline level (DC offset) as a parameter for which our methods fit. This ensures that the estimated τ value is centred on the true simulated value for τ .

Table 1 summarises the impact of using different methods to estimate the spectral indices. Errors in τ increase towards lower frequencies (or higher scattering) for all the fitting methods. Given a pulse shape and a duty cycle, τ in millisecond pulsars will become comparable to the pulse period at higher frequencies than for normal pulsars.

Larger errors in τ lead to larger errors in α . Our simulations show that at high levels of scattering there is an underlying asymmetry in the measurement uncertainties of τ . Correctly accounting for the asymmetry results in a more accurate measurement of α . The impact of this correction is seen for the millisecond pulsar case, where the mean estimate of α changes from 3.87 to 4.01 over the low frequency range, while the standard deviation in the estimate simultaneously drops from 0.32 to 0.17. Having introduced this

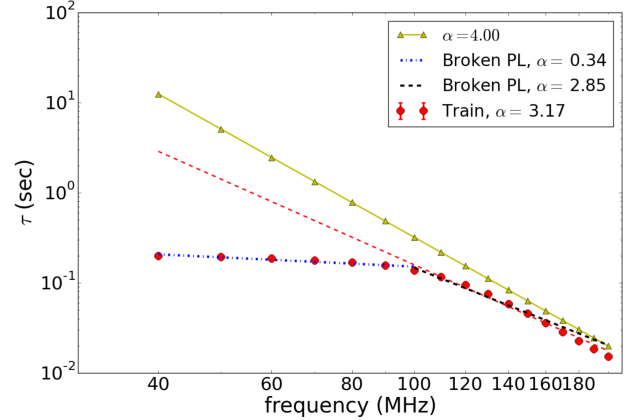


Figure 20. The τ spectrum associated with a truncated screen. At low frequencies (high scattering) the τ values as obtained by the train method (red circles, with error bars smaller than the marker) greatly underestimates the input scattering (yellow triangles). This is due to the fact that at large wavelengths the scattering process will become sensitive to the edges of the screen leading to imprints in the temporally broadened pulse profiles for which our method does not account. The short black (dashed) and blue (dot-dash) lines show a broken power law fit with spectral indices as indicated in the legend, and the long red (dashed) line is the weighted power law fit of all the obtained τ values.

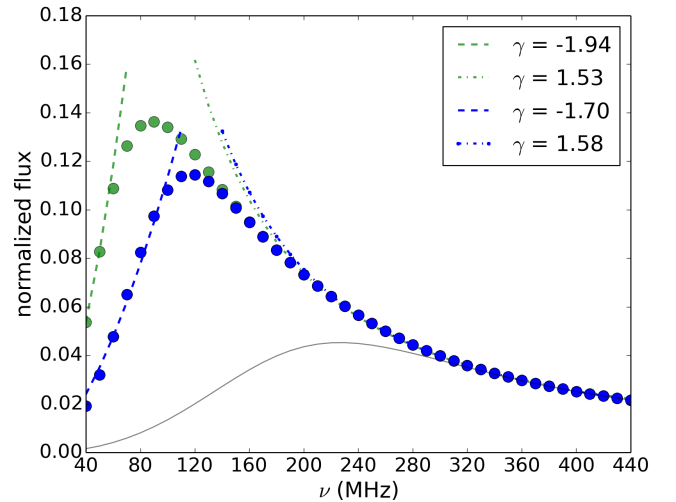


Figure 21. The flux spectrum of the truncated scattering screen setup discussed in the text is shown in blue. The turnover in the spectrum is due to the longer wavelengths being emitted beyond the boundaries of the screen such that not all of the pulsar's flux fill be received back at the detector. The equivalent anisotropic case for a screen that scatters 3 time more weakly in one dimension is presented in green. The turnover in this case will occur at a lower frequency compared to the stronger isotropic scatterer. The grey line shows the flux spectrum for a smaller screen (100 by 200 AU). Decreasing the screen size leads to turnovers at higher frequencies.

correction, the α values for all our simulated data have a mean value within 0.3% of the true value, and a standard deviation below 0.2.

Increasing the SNR also leads to a decrease in the standard deviation of α estimates. In the case of the 1.0 sec pulsar at 60 – 120 MHz, an increase of SNR from 10 to 20 to 50 decreases the standard deviation by a factor of 2 and 2.25, as shown in Table 2.

The EMG method in this paper portrays the most basic model, where the wrap around of a scattering tail is not modelled. For such an implementation deviations from the theoretically expected values at low frequencies can be as large as 30% – 70% for the pulsars and scattering scenarios we simulate. The failure of this model is made explicit to encourage the revision of fits in the literature where wrap around scattering tails may not have been modelled, e.g. Fig. 1 in Lewandowski et al. (2013).

We show in Fig. 7 that it is important to attribute the measured τ values to the associated monochromatic frequency (f_m , computed from the central frequency and bandwidth of the observation as in eq. 6) to obtain accurate estimates of the spectral index (α) from a power law fit.

We investigate the scatter broadening of pulsars due to anisotropic scattering, which we model as an asymmetric Gaussian distribution, and show that fitting these profiles with isotropic models leads to systematic errors in the measurement of τ . Assuming the scattering strength in one dimension to be weaker than what we have considered as typical in our simulations, such fits will lead to apparent τ spectra with $\alpha < 4$ (in our example $\alpha = 2.95$). Improvements to such fits will rely on using more appropriate broadening functions, such as eq. (4) or (8), in the fitting model. We are currently investigating the use of anisotropic fitting techniques to analyse scattered profiles from LOFAR observations. Given the recent evidence for highly anisotropic scattering our simulations show that the measured α (and goodness of fit for τ) could serve as a diagnostic in the time domain for the type of scattering screen.

From the simulated truncated screens in Sec. 3.3, we show that an isotropic thin screen with dimensions 400 AU by 600 AU midway between the Earth and a pulsar at 3 kpc, changes the shape of the pulse profile with frequency and leads to a flattened τ spectrum below 200 MHz. In Fig. 20 the spectral index was estimated to be $\alpha = 0.34$ below 100 MHz. Adjustments to the modelled screen size and location will impact the shape and features of the τ spectra, resulting in possible breaks.

For a complex pulse profile the frequency evolution due to this kind of scattering may change the component amplitude ratio. We use the profile of B1237+25 to demonstrate this effect, although this pulsar is not observed to exhibit such behaviour. The effects caused by scattering can in principle be disentangled from the intrinsic emission, as it is unlikely that both will have the same dependence on frequency.

Fig. 22 provides a cartoon style summary of deviations from the theoretical τ spectrum we use (a power law spectrum with spectral index 4) and the underlying causes. Subscripts I and A refer to isotropic and anisotropic scattering mechanisms, and ∞ or tr to either infinite or truncated thin screens.

For the range of scattering mechanisms and geometries

discussed above, we also investigated the impact on pulsar flux spectra. Observationally, pulsar flux spectra have steep negative spectral indices, typically within the range $\gamma = -1.8$ to -1.4 (Lorimer et al. 1995; Maron et al. 2000; Bates et al. 2013). Spectral turnovers at low frequencies are a known feature of the population of slow pulsars (e.g. Lorimer et al. 1995). Thermal absorption by dense surrounding environments (e.g. Rajwade et al. 2016) may be partially responsible for spectral turnovers, which may also be intrinsic to the radio emission of the pulsar. Some evidence suggests that the efficiency of the radio emission shuts off at lower frequencies (Malofeev et al. 2000). Alternatively, under the assumption that the lower frequencies are emitted higher above the neutron star surface than the higher frequencies (Radhakrishnan & Cooke 1969), the intersection of the line of sight with the radio emitting regions in the magnetosphere will also be frequency dependent, with the possibility of the line of sight gradually missing the emitting regions altogether at low frequencies. A small group of pulsars show turnovers at much higher frequencies. These are known as the gigahertz peaked spectrum (GPS) pulsars. The GPS turnovers are also attributed to either the radio emission process becoming inefficient above a threshold frequency, or dense ISM environments (Kijak et al. 2007, 2011).

Here we concentrate on the effects of scattering on the observed flux, which should occur in addition to the above mechanisms. In our simulations we show that high levels of scattering by infinite screens, result in a measurable integrated flux loss, due to the rise in the off-pulse baseline. In the case of slow pulsars this will lead to spectral turnovers typically below 100 MHz, as shown in Fig. 16. The turnover frequency due to this effect in millisecond pulsar flux spectra, would be higher than for slow pulsars. Older studies of millisecond pulsar flux spectra have mostly not found turnovers (Kuzmin & Losovsky 2000; Kuzmin 2002), however more recent work by Kuniyoshi et al. (2015) have observed turnovers in MSP flux spectra, albeit in imaging observations. This could point to anomalous scattering by e.g. finite scattering clouds. The lack of observed spectral turnovers from time-domain studies suggest that most MSP profiles are not substantially scattered in current low frequency observations.

Fig. 15 shows the difference in turnover frequencies for an infinite isotropic and anisotropic scattering screen. In our example of an anisotropic screen, the spectral turnover occurs at ~ 50 MHz. The turnover frequency for the corresponding isotropic case was ~ 90 MHz, i.e. higher by a factor of $\sqrt{\sigma_{ax}\sigma_{ay}}/\sqrt{\sigma_i^2}$ (where subscripts a and i refer to anisotropic and isotropic respectively and (x, y) are the two dimensions describing the thin screen, with $\sigma_i = \sigma_{ax} = 3$ and $\sigma_{ay} = 1$ mas at 1GHz).

Similar spectral turnovers can be caused by the loss of flux associated with truncated scattering screens. In Fig. 21 we show that a finite screen of size 400 AU by 600 AU at the same location as the infinite screen, would lead to turnover frequencies in approximately the same range, with the isotropic flux spectrum peaking at 125 MHz and the anisotropic spectrum at 90 MHz. Keeping the scattering strength constant, the turnover frequency can be pushed up to 230 MHz for a screen of size 100 AU by 200 AU, centred on the line of sight. In this paper, we have concentrated on screens with infinitely small inner scales, i.e. the scales of

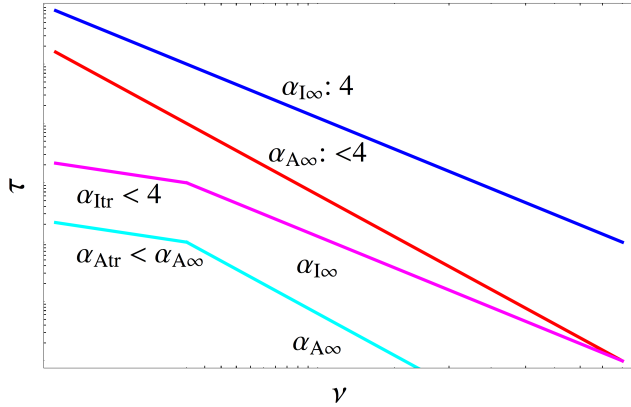


Figure 22. A cartoon summary of the deviations from the theoretical τ spectrum and their underlying causes. Subscripts I and A refer to isotropic and anisotropic scattering mechanisms, and ∞ or tr to either infinite or truncated thin screens. The anisotropic scatterer shown here scatters more weakly in one dimension than what is considered a typical isotropic scattering strength.

structure in the screen that result in the largest scattering angles. A larger inner scale, as discussed in Rickett et al. (2009), would result in a maximum scattering angle, similar to the case of a screen of finite size, and hence similar spectra for τ . The crucial difference between the two is that the existence of an inner scale does not reduce the total observed flux, as is the case for a finite screen.

To reproduce GPS spectra with our modelling would require a fine tuning of the truncated screen size and or the simulated scattering strength. Studying the correlated changes in flux spectra, pulsar profiles and their associated τ -spectra with frequency can reveal the complex interplay between the phenomena that set the observed pulse profile and its dependence on frequency, and should lead to a better understanding of the effects of each.

5 CONCLUSIONS

We have showed that our train+DC method is an efficient forward fitting technique for scattered profiles. Our investigations show flatter τ spectra (than what is theoretically expected) when using an isotropic model to fit profiles shaped by an anisotropic scattering mechanism. Flatter τ spectra can also be the result of low frequency profile shape changes caused by a truncated scattering screen. Over a large enough sampled frequency range, such a τ spectrum will exhibit a break. Analysing the impact of scattering on flux spectra, has revealed turnovers at low frequencies due to baseline effects. Similarly truncated screens can cause turnovers in flux spectra. These turnovers are a function of the screen size and its scattering strength.

We anticipate new data from instruments such as LOFAR and the MWA in the frequency range studied here. This paper will serve as a framework to compare real data to the simulated results obtained here. Our analysis of scattering imprints is by no means complete, as we have only investigated simple profile shapes and Gaussian scattering mechanisms. However, our results indicate where attention should be paid in terms of creating accurate fitting meth-

ods as well as describing possible sources of deviations to theoretically predicted results.

6 ACKNOWLEDGEMENTS

M. Geyer is a Commonwealth Scholar funded by the UK government. The authors also express their gratitude to SKA South Africa for their continued support. We would like to thank Jayanth Chennamangalam and Mark Walker for valuable discussions and insights relating to this work. We thank the anonymous reviewer for their valuable comments to improve the quality of the paper.

References

- Armstrong J. W., Rickett B. J., Spangler S. R., 1995, *ApJ*, 443, 209
- Bates S. D., Lorimer D. R., Verbiest J. P. W., 2013, *MNRAS*, 431, 1352
- Bhat N. D. R., Cordes J. M., Camilo F., Nice D. J., Lorimer D. R., 2004, *ApJ*, 605, 759
- Bhat N. D. R., Cordes J. M., Chatterjee S., 2003, *ApJ*, 584, 782
- Bignall H. E., Jauncey D. L., Lovell J. E. J., Tzioumis A. K., Kedziora-Chudczer L., Macquart J.-P., Tingay S. J., Rayner D. P., Clay R. W., 2003, *ApJ*, 585, 653
- Bilous A., Kondratiev V., Kramer M., Keane E., Hessels J., Stappers B., Malofeev V., Sobey C., Breton R., Cooper S., Falcke H., Karastergiou A., Michilli D., 2015, *ArXiv eprints* (astro-ph/1511.01767)
- Briskin W. F., Macquart J.-P., Gao J. J., Rickett B. J., Coles W. A., Deller A. T., Tingay S. J., West C. J., 2010, *ApJ*, 708, 232
- Cordes J., Lazio T., 2001, *ApJ*, 549, 997
- Cordes J. M., 1986, *ApJ*, 311, 183
- Cronyn W. M., 1970, *Science*, 168, 1453
- Dembska M., Basu R., Kijak J., Lewandowski W., 2015, *MNRAS*, 449, 1869
- Fiedler R., Dennison B., Johnston K., Hewish A., 1987, *Nature*, 326, 675
- Gupta Y., Rickett B. J., Lyne A. G., 1994, *MNRAS*, 269, 1035
- Johnston S., Weisberg J. M., 2006, *MNRAS*, 368, 1856
- Karastergiou A., Johnston S., 2007, *MNRAS*, 380, 1678
- Kijak J., Gupta Y., Krzeszowski K., 2007, *A&A*, 462, 699
- Kijak J., Lewandowski W., Maron O., Gupta Y., Jessner A., 2011, *A&A*, 531, A16
- Kuniyoshi M., Verbiest J. P. W., Lee K. J., Adebahr B., Kramer M., Noutsos A., 2015, *MNRAS*, 453, 828
- Kuzmin A., 2002, in Pramesh Rao A., Swarup G., Gopal-Krishna eds, *The Universe at Low Radio Frequencies* Vol. 199 of *IAU Symposium, Low-frequency Observations of Millisecond Pulsars*. p. 389
- Kuzmin A. D., Losovsky B. Y., 2000, *Astronomy Letters*, 26, 500
- Lambert H. C., Rickett B. J., 1999, *ApJ*, 517, 299
- Lang K. R., 1971, *ApJ*, 164, 249
- Lee L. C., Jokipii J. R., 1976, *ApJ*, 206, 735
- Lewandowski W., Dembska M., Kijak J., Kowalińska M., 2013, *mnras*, 434, 69

- Lewandowski W., Kowalińska M., Kijak J., 2015, MNRAS, 449, 1570
- Löhmer O., Kramer M., Mitra D., Lorimer D. R., Lyne A. G., 2001, ApJ Letters, 562, L157
- Löhmer O., Mitra D., Gupta Y., Kramer M., Ahuja A., 2004, A&A, 425, 569
- Lorimer D. R., Yates J. A., Lyne A. G., Gould D. M., 1995, MNRAS, 273, 411
- Lyne A. G., Rickett B. J., 1968, Nature, 218, 326
- Malofeev V. M., Malov O. I., Shchegoleva N. V., 2000, Astronomy Reports, 44, 436
- Maron O., Kijak J., Kramer M., Wielebinski R., 2000, Astron. Astrophys. Suppl., 147, 195
- Newville M., Stensitzki T., Allen D. B., Ingargiola A., 2014, <http://dx.doi.org/10.5281/zenodo.11813>
- Pilia M., Hessels J. W. T., Stappers B. W., Kondratiev V. I., Kramer M., van Leeuwen J., Weltevrede P., Lyne A. G., Zagkouris K., more. 2015, ArXiv eprints (astro-ph/1509.06396)
- Radhakrishnan V., Cooke D. J., 1969, ApJ Letters, 3, 225
- Rajwade K., Lorimer D. R., Anderson L. D., 2016, MNRAS, 455, 493
- Ramachandran R., Mitra D., Deshpande A. A., McConnell D. M., Ables J. G., 1997, MNRAS, 290, 260
- Rickett B., Johnston S., Tomlinson T., Reynolds J., 2009, MNRAS, 395, 1391
- Rickett B. J., 1977, ARA&A, 15, 479
- Rickett B. J., 1990, ARA&A, 28, 561
- Rickett B. J., Coles W. A., Bourgois G., 1984, A&A, 134, 390
- Romani R. W., Narayan R., Blandford R., 1986, MNRAS, 220, 19
- Scheuer P. A. G., 1968, Nature, 218, 920
- Smith F. G., Thompson J. H., 1988, Optics. John Wiley and Sons Ltd.
- Stappers B. W., Hessels J. W. T., Alexov A., Anderson K., Coenen T., Hassall T., Karastergiou A., 2011, A&A, 530, A80
- Tingay S., Goeke R., Hewitt J. N., Morgan E., Remillard R. A., Williams C. L., Bowman J. D., 2012, in Resolving The Sky - Radio Interferometry: Past, Present and Future Realisation of a low frequency SKA Precursor: The Murchison Widefield Array. p. 36
- Tuntsov A. V., Bignall H. E., Walker M. A., 2012, ArXiv eprints (astro-ph/1212.0633)
- van Haarlem M. P., Wise M. W., Gunst A. W., Heald G., McKean J. P., Hessels J. W. T., de Bruyn A. G., more. 2013, A&A, 556, A2
- Walker M. A., 2007, in Haverkorn M., Goss W. M., eds, SINS - Small Ionized and Neutral Structures in the Diffuse Interstellar Medium Vol. 365 of Astronomical Society of the Pacific Conference Series, Extreme Scattering Events: Insights into the Interstellar Medium on AU-Scales. p. 299
- Walker M. A., Melrose D. B., Stinebring D. R., Zhang C.-M., 2004, MNRAS
- Williamson I. P., 1972, MNRAS, 157, 55
- Williamson I. P., 1973, MNRAS, 163, 345